

HELICAL BLOWOUT JETS IN THE SUN: UNTWISTING AND PROPAGATION OF WAVES

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ABSTRACT

We report on a numerical experiment of the recurrent onset of helical “blowout” jets in an emerging flux region. We find that these jets are running with velocities of $\sim 100\text{--}250\text{ km s}^{-1}$ and they transfer a vast amount of heavy plasma into the outer solar atmosphere. During their emission, they undergo an untwisting motion as a result of reconnection between the twisted emerging and the non-twisted pre-existing magnetic field in the solar atmosphere. For the first time in the context of blowout jets, we provide direct evidence that their untwisting motion is associated with the propagation of torsional Alfvén waves in the corona.

Key words: magnetohydrodynamics (MHD) – Sun: activity – Sun: interior – Sun: magnetic fields

1. INTRODUCTION

Solar jets have been observed at various wavelengths, e.g., $H\alpha$ (Schmieder et al. 1995; Canfield et al. 1996), EUV, and X-ray (Shibata et al. 1992; Alexander & Fletcher 1999). They occur in emerging flux regions (EFR; e.g., Shibata et al. 1992), active regions (e.g., Canfield et al. 1996; Shimojo et al. 1996), coronal holes (e.g., Wang et al. 1998; Cirtain et al. 2007), etc. Recently, a dichotomy of jets was suggested by Moore et al. (2010). About two-thirds of the observed jets fit the “standard” reconnection picture, which invokes reconnection between oppositely directed magnetic fields, e.g., an emerging and a pre-existing magnetic field (e.g., Heyvaerts et al. 1977). The other one-third have been called “blowout” jets, which are triggered by an eruption. More precisely, it has been observed (e.g., Moore et al. 2010; Sterling et al. 2010; Shen et al. 2012) that the precursor of a “blowout” jet is a twisted and/or a sheared arch, which often carries a (small) filament or flux rope within it. When this structure becomes unstable and erupts, it blows out the envelope field producing an untwisting ejection of cool (e.g., chromospheric) and hot material. This “blowout” jet is a broad, curtain-like structure in opposition to the “standard” jet, which is more elongated and it is not commonly associated with an eruption event.

Three-dimensional (3D) simulations on the formation of blowout jets, during the emergence and eruption of solar magnetic fields, have been carried out by Archontis & Hood (2013), Moreno-Insertis & Galsgaard (2013), and Fang et al. (2014). Archontis & Hood (2013) have shown that the interaction between an emerging (twisted) magnetic field and an ambient (non-twisted) magnetic field in the solar corona can trigger both standard and “blowout” jets. Their experiments reproduced some of the observed characteristics of the blowout jets, their internal helical structure, and their overall curtain-like shape. Moreno-Insertis & Galsgaard (2013) studied the recurrent onset of eruptions in an EFR and their possible relationship to the subsequent emission of “blowout” jets. Using a similar numerical setup, Fang et al. (2014) showed that heat conduction leads to an increase of the total mass ejection in the corona during the emission of the “blowout” jets. In previous simulations, Shibata & Uchida (1985) had shown that an unwinding jet could be the result of reconnection between a twisted loop and an open flux tube. The unwinding motion, which has also been reported in observations (e.g., Cirtain et al. 2007; Moore et al. 2013), has been interpreted as the propagation of torsional Alfvén waves

(see also Nishizuka et al. 2008) releasing the stored twist from a twisted magnetic loop into the open ambient field. However, no *direct* evidence of *propagating* Alfvén waves in “blowout” jets has been provided so far, either on numerical or observational studies.

Here, we report on the recurrent emission of “blowout” jets in an EFR with a sea-serpent configuration. We show that the “blowout” jets are untwisted during their ejection and, for the first time, we provide direct evidence of propagating torsional Alfvén waves during the emission of the “blowout” jets.

2. THE MODEL

We solve the 3D time-dependent, resistive and compressible MHD equations in Cartesian geometry, using the Lare3d code (Arber et al. 2001). Explicit (uniform) resistivity of $\eta = 10^{-2}$ is included.

The initial atmosphere consists of horizontal and homogeneous parallel layers in hydrostatic equilibrium. The solar interior is represented by a layer in the range ($-5.4\text{ Mm} \leq z < 0\text{ Mm}$), which is adiabatically stratified. The photosphere/chromosphere layer lies at $0\text{ Mm} \leq z < 1.9\text{ Mm}$. The temperature at the photosphere is 5100 K and it increases up to $\approx 3 \times 10^4\text{ K}$ at the top chromosphere. The transition region is located at heights $1.9\text{ Mm} \leq z \leq 2.7\text{ Mm}$. Above it, there is an isothermal layer ($\mathcal{O}(1)\text{ MK}$), which is mimicking the lower solar corona ($2.7\text{ Mm} < z \leq 57.6\text{ Mm}$). In the solar interior, the magnetic field is a horizontal magnetic flux tube, which is twisted. The flux tube is located at $z_0 = -2.1\text{ Mm}$ and its axis is parallel to the y axis at $x = 0$. The component of the magnetic field along the tube’s axis (axial field) is given by

$$B_y = B_0 \exp(-r^2/R^2), \quad B_\theta = \alpha r B_y, \quad (1)$$

where the tube’s radius is $R = 450\text{ km}$ and r is the radial distance from the tube’s axis ($r^2 = x^2 + (z - z_0)^2$). The twist of the field lines around the tube’s axis is uniform and it is given by $\alpha = 2.2 \times 10^{-3}\text{ km}^{-1}$. With this twist, the tube is marginally stable to the kink instability. Initially, we apply a deficit in density along the tube’s axis, making two segments (i.e., at $x = 0$, $y = \pm 3.2\text{ Mm}$) more buoyant than the rest of the tube:

$$\Delta \rho = [p_t(r)/p(z)] \rho(z) \exp(-y^2/\lambda^2) \sin^2(2\pi y/\omega). \quad (2)$$

Thus, $\Delta \rho$ is the difference between the background density and the density inside the flux tube after we apply the density

deficit along its axis. The pressure within the flux tube is p_t , λ defines half the length of each buoyant part of the tube, and ω is defined as half of the flux tube length. We use $\lambda = 3.6$ Mm, $\omega = 31.5$ Mm, and an initial field strength of $B_0 = 2.4$ kG, which corresponds to plasma $\beta \approx 28$. The corona is filled with an ambient field which is oblique, and is defined by

$$B_c = B_c(z) (0, \cos \theta, \sin \theta), \quad (3)$$

where $\theta = 80^\circ$, and $B_c(z) \approx 3$ G for $z \geq 0.54$ Mm and gradually decreases to 0 for $z < 0.54$ Mm. The numerical domain is $[-31.5, 31.5] \times [-31.5, 31.5] \times [-5.4, 57.6]$ Mm in the direction perpendicular to the tube (x), along the tube's axis (y) and vertical (z), respectively. The numerical grid is a 420^3 box with periodic boundary conditions in the x and y directions. At the top, there is an open boundary allowing plasma to flow out of the grid. There is a non-penetrating conducting wall at the bottom boundary.

3. RESULTS AND DISCUSSION

In the following, we discuss the results of our simulations showing: (1) the emergence of the sea-serpent flux tube at the photosphere, (2) the recurrent onset of the jets, and (3) the propagation of torsional Alfvén waves during the untwisting motion of the blowout jet(s). Figure 1(a) shows the initial (i.e., at $t = 0$) topology of the magnetic field. Figure 1(b) shows the connectivity of the field lines during the first blowout jet emission. The (blue) field lines show (1) the sea-serpent configuration of the emerging field and (2) the twisted field lines at the center of the EFR, which are a result of reconnection between the emerging loops. The (red) field lines are reconnected field lines that join the ambient with the emerging field. Due to the sea-serpent configuration, the downward tension of the uppermost field lines (i.e., the envelope field) of each emerging loop is released more effectively due to reconnection with (both) the ambient field and the field of the neighboring emerging loop. Thus, unavoidably, this interaction affects the onset time and dynamics of the EFR's eruptions, compared to a “single emerging loop-ambient field” reconnection case (e.g., Archontis & Hood 2013).

Figure 2(a) shows the magnetogram during the emergence of the field into the photosphere at $t \simeq 165.7$ minutes. The overplotted arrows represent the direction of the magnetic field vector. Due to the initial density deficit, the emerging magnetic field at the solar surface forms two bipolar regions (hereafter BR1 and BR2). At $t \simeq 237.1$ minutes (Figure 2(b)), the polarities of each BR have moved apart along the y direction. In previous studies (e.g., Archontis & Török 2008), it has been shown that this movement is followed by shearing along the polarity inversion line and the formation of a new flux rope, which might erupt into the corona. This is also found to occur in the present simulation within each BR. Moreover, we find that shearing and reconnection leading to the formation of a flux rope also occurs between the opposite polarities P1 and N2 due to their relative motion.

Before the onset of the eruptions ($t < 200$ minutes), each emerging bipole comes into contact and eventually reconnects with the ambient magnetic field, giving onset to hot reconnection jets. Figures 2(c) and (d) show the temperature and v_z distribution, respectively, at $x = 0$ and $t \simeq 165.7$ minutes. We find that the upward reconnection jets are moving with a velocity of ~ 40 km s $^{-1}$, reaching temperatures of up to 1×10^6 K at the reconnection site, and 8×10^5 K within the jet channels.

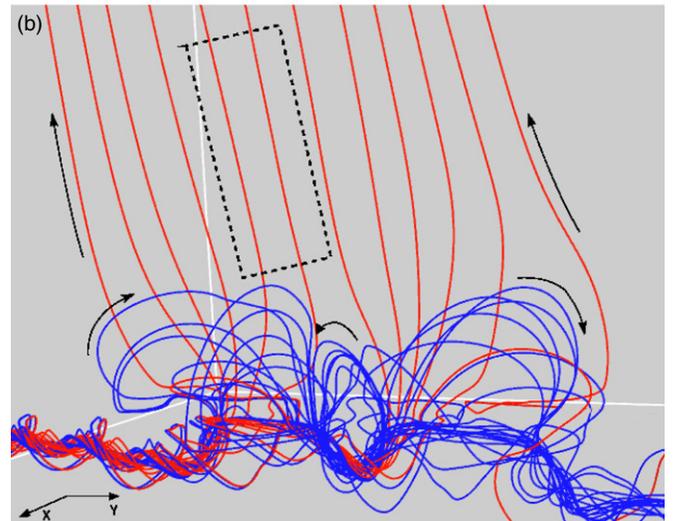
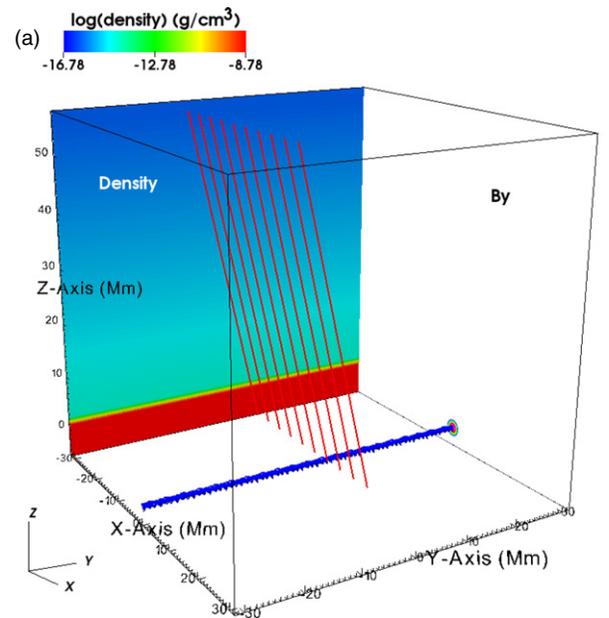


Figure 1. (a) Selected magnetic field lines (red) show the orientation of the oblique magnetic field in the corona. The twisted magnetic flux tube is shown by the blue field lines. The colored vertical y - z plane shows the density distribution of the background atmosphere. Isocontours at the x - z plane represent B_y . (b) A close-up of the field line topology ($t = 232.8$ minutes). Field lines have been traced from the two emerging bipoles (blue) and from the ambient field (red). Arrows show the direction of the magnetic field. The (dashed) rectangle indicates the channel along which the first blowout jet is emitted.

Figures 2(e) and (f) show the first eruptive event within the EFR. The eruption of the cool plasma starts within BR2. Figure 2(e) shows the temperature and Figure 2(f) shows the density distribution at $y \simeq -2$ Mm and $t = 230$ minutes. The rise of the erupting material induces inflow toward the interface, where a current layer has built up between the envelope and the ambient field. This leads to more external reconnection between the two magnetic flux systems and the onset of a hot and fast external reconnection jet. The plasma at the interface is heated up to 3.5×10^6 K. A side effect of the external reconnection is the formation of a hot external arcade with temperature up to 5.5×10^6 K.

At $t \simeq 232.8$ minutes (Figures 2(g) and (h)), the eruption blows out the envelope field and the dense erupting material is emitted along the reconnected field lines of the oblique ambient field. Now the channel of the blowout jet and the external arcade

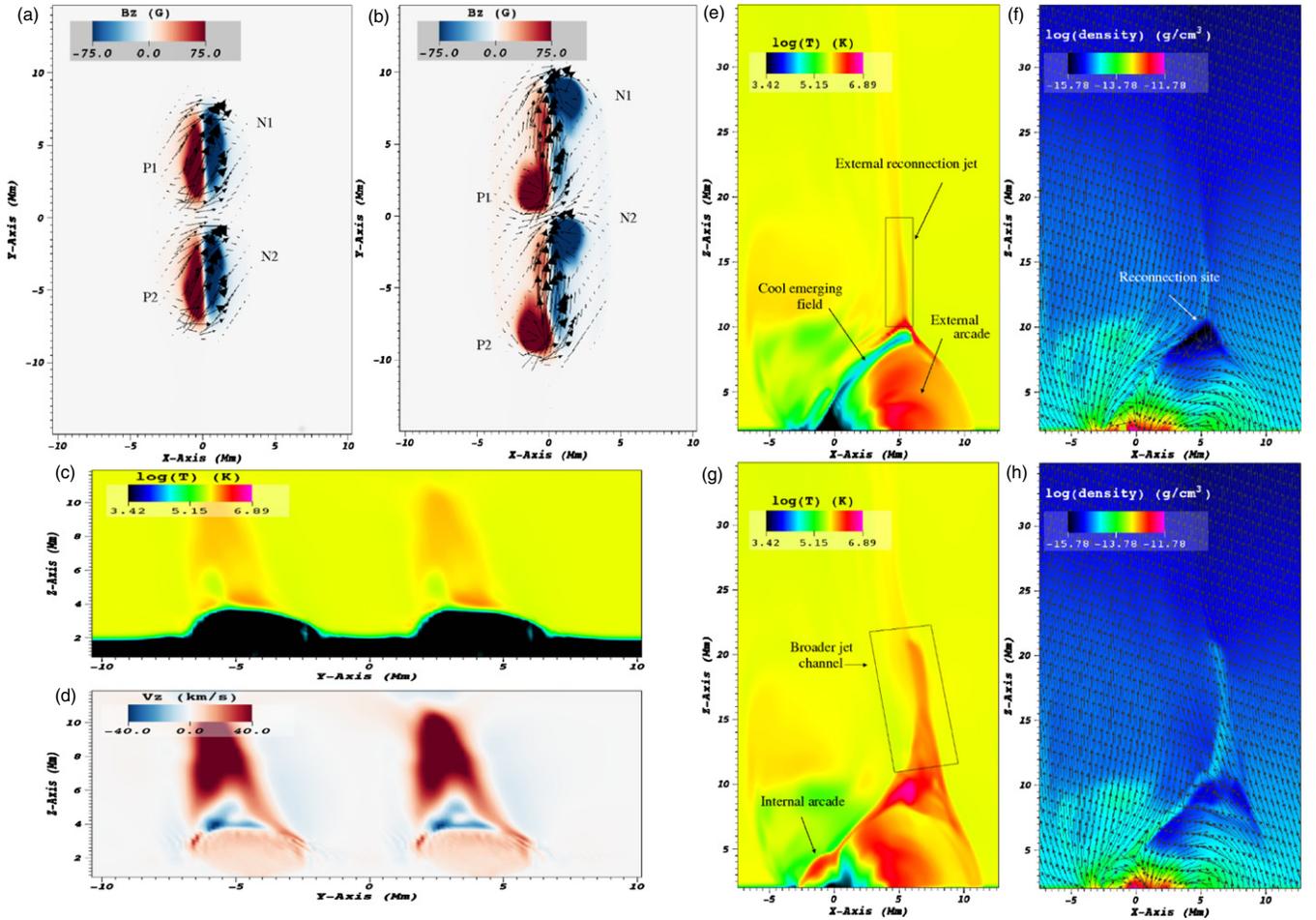


Figure 2. Magnetogram during the emergence at the photosphere, showing the two bipolar regions (BR1 consisting of N1 and P1, and BR2 consisting of N2 and P2) at two times: (a) $t = 165.7$ minutes and (b) $t = 237.1$ minutes. Temperature (c) and v_z (d) distribution during the emission of the reconnection jets at $t = 165.7$ minutes at $x = 0$ Mm. Temperature (e) and density (f) distribution during the eruption preceding the first blowout jet at $t = 230.0$ minutes. Temperature (g) and density (h) distribution during the first “blowout” jet at $t = 232.8$ minutes. Arrows denote the direction of the magnetic field vector (projected onto the plane). The vertical slice in (e)–(h) is at $y = -2$ Mm.

are much wider. Cool (5×10^5 K) and hot (1×10^6 K) plasmas are ejected along the blowout jet’s channel and heating (due to internal reconnection) is produced underneath the erupting plasma, where an internal arcade of temperature up to 4×10^6 K is formed.

The top panel in Figure 3 shows that there are six events during which v_z varies in the range ~ 100 – 250 km s $^{-1}$. Each event consists of several peaks in v_z . The fact that the peaks of minimum and maximum v_z occur at approximately the same time suggests the existence of bidirectional flows from several sites within the EFR. We find that the first peak ($t \approx 165$ minutes) corresponds to the jets caused by reconnection between the emerging and the ambient magnetic field. The second event ($t \approx 185$ – 210 minutes, $v_z \sim 150$ km s $^{-1}$) is the composite effect of an eruption, which starts from the area in between BR1 and BR2, and a reconnection jet that follows the eruption. The jet is initiated due to the restructuring of the nearby magnetic field due to the eruption (in a similar manner to the external reconnection jet in Figure 2(e)). However, on this occasion, the erupting plasma becomes confined by the local ambient field and does not evolve into a profound blowout jet that could reach the higher solar atmosphere. Instead, it is moving sideways from the center of the EFR toward the BR1 where the total pressure is less.

The other four events correspond to the emission of “blowout” jets driven by eruptions that emanate from within the BRs. These jets blast heavy material toward the outer solar atmosphere. The bottom panel in Figure 3 shows the running difference of the integrated plasma density $I = \int \rho^2 dx dy$ for temperatures above 8×10^5 K. It is found that there are four marked events during which dense plasma is brought into the higher atmosphere. The steep vertical slope at the end of each event indicates that most of the ejected plasma is transported upward. The dark regions after the brightening peaks indicate that part of the erupting plasma undergoes gravitational draining. Note that the time period between the onset of the blowout jets is not so “quiet.” There are various, less striking events, which do not dump enough mass and energy into the outer corona. Such events (e.g., at $t \approx 190$ – 210 minutes, $t \approx 310$ minutes, etc.) are small confined eruptions and reconnection coronal jets. By comparing the two panels in Figure 3, we find a good temporal correlation between the four “blowout” jets and the emission of bidirectional flows, which occur after $t \approx 220$ minutes. This is direct evidence that the “blowout” jets are closely associated with reconnection, which is driven by the eruption of cool material from the low atmosphere.

The erupting core of the blowout jets is a twisted flux tube. Figure 4(a) illustrates the magnetic field lines within

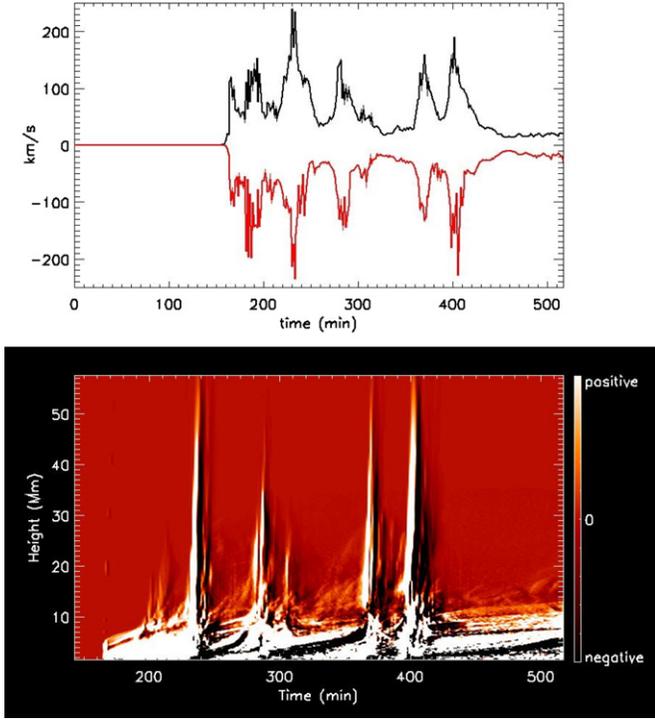


Figure 3. Top: temporal evolution of the maximum (black) and minimum (red) v_z above the photosphere. Bottom: height–time diagram (running difference) of $\int \rho^2 dx dy$, where $T > 8 \times 10^5$ K.

and around the first blowout jet (see also Archontis & Hood 2013). Reconnection of the erupting (twisted) field with the open (non-twisted) ambient field would most likely relax the twist and lead to an untwisting motion of the helical blowout jet. A measurement of the twist is given by j_{\parallel}/B , where j_{\parallel} is the parallel component of the current along the reconnected field lines of the blowout jet. Figure 4(b) shows the temporal evolution of the average $|j_{\parallel}|/B$ along the channel of the blowout jet. We calculate this for heights above $z \approx 9$ Mm, which is the approximate horizontal boundary between the envelope field of the emerging flux and the lower part of the blowout jet. The increase of the twist (up to $t = 235$ minutes) is due to the eruption of the flux rope and the subsequent decrease indicates the untwisting motion of the jet. Figure 4(c) shows the distribution of $|j_{\parallel}|/B$ at the vertical slice with $x = 5.4$ Mm. It is confirmed that the blowout jet possesses considerable twist along its main stream (e.g., in the range $z = 10\text{--}40$ Mm). Note that $|j_{\parallel}|/B$ is also very strong (as expected) in the twisted emerging field underneath the blowout jet (e.g., the area with strong B_x at $z < 7$ Mm).

To show the untwisting motion of the jet, we plot the v_x and B_x distribution (Figures 4(d) and (e), respectively). There is positive v_x (pointing out of the plane) at the right-hand side of the jet and negative v_x (pointing into the plane) at the left-hand side. This corresponds to a magnetized plasma motion where B_x (i.e., the transverse component of the magnetic field) is pointing in the opposite direction to v_x and the associated vorticity (i.e., $\omega = \nabla \times v$) is pointing downward. This is a first indication that the jet undergoes an untwisting motion. More evidence is found

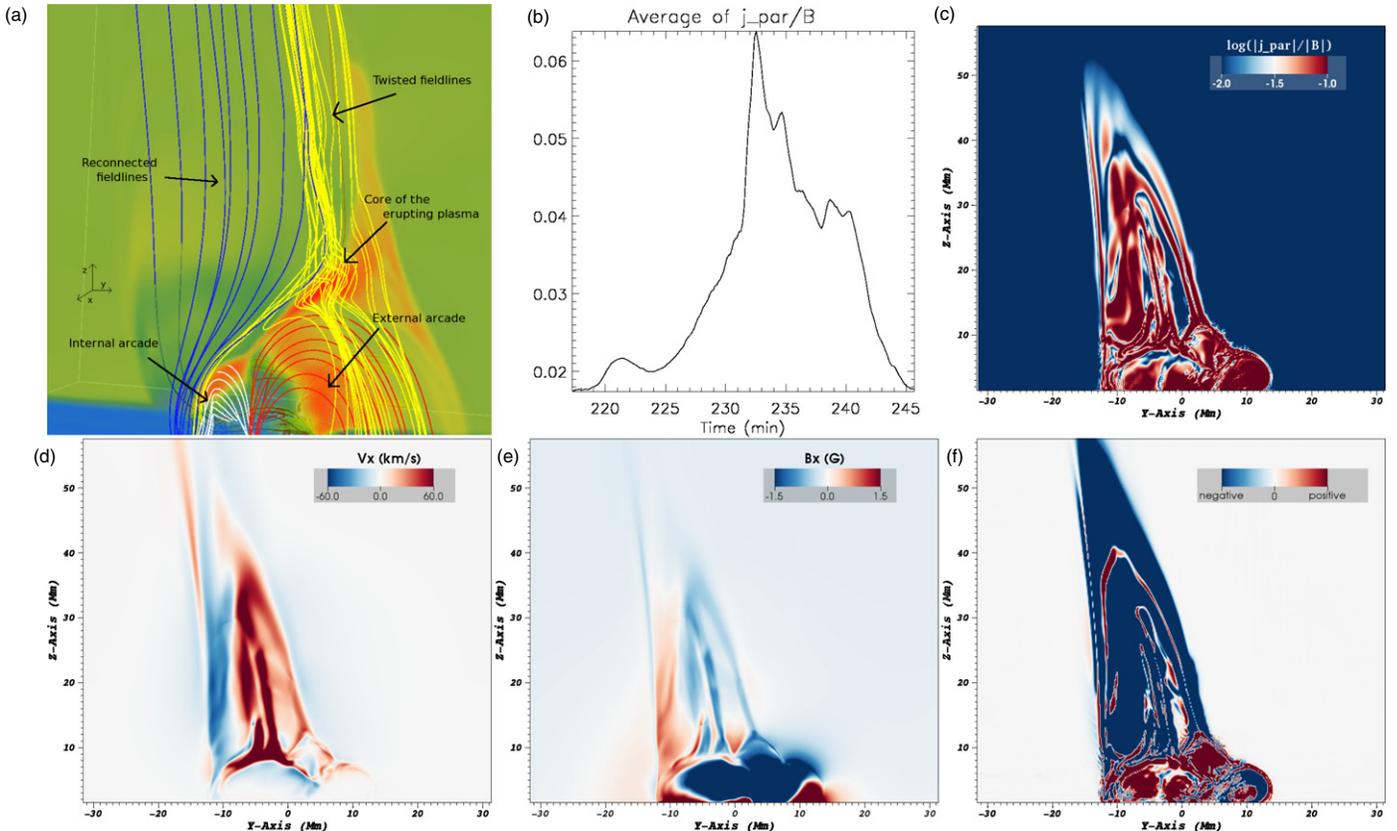


Figure 4. (a) Visualization of the magnetic field lines, showing the helical nature of the blowout jet (yellow), the reconnected field lines that join the emerging with the ambient field (blue), the internal arcade (white), and the external arcade (red). (b) The temporal evolution of the average $j_{\parallel}/|B|$ within the blowout jet. (c) The $|j_{\parallel}|/|B|$ distribution, (d) v_x distribution, and (e) B_x distribution, at $x = 5.4$ Mm, $t = 233.4$ minutes. (f) $j_{\parallel} \cdot \omega_{\parallel}$, where j_{\parallel} (ω_{\parallel}) is the current (vorticity) parallel to the reconnected field lines.

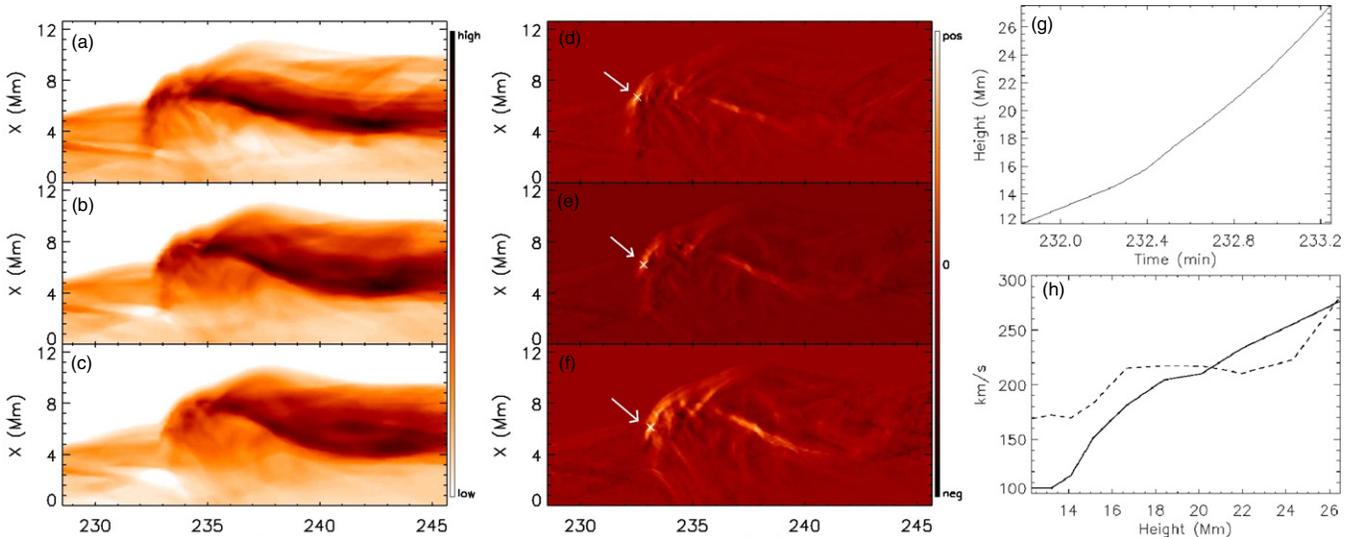


Figure 5. Distance–time diagram of $\int \rho^2 dy$, for $0.6 \text{ mK} < T < 1.2 \text{ mK}$, across the jet at (a) 17 Mm, (b) 21.5 Mm, and (c) 25 Mm. (d)–(f) are the running differences of (a)–(c), respectively. (g) shows the height–time profile of the heavy front of the jet, obtained by tracing the first maximum of the running difference (white cross in (d)–(f)) at various heights. The velocity of this front is shown in panel (h). The dashed line in (h) is the local Alfvén speed in the close vicinity of the front of the wave.

by studying the distribution of current and vorticity within the jet.

Figure 4(f) shows the product of j_{\parallel} with the vorticity ω_{\parallel} . Within the jet channel, this product is mostly negative, indicating that the current and vorticity are pointing in different directions. In fact, we have found (not shown in this figure) that ω_{\parallel} is negative due to the direction of the flow (and it is pointing downward), while j_{\parallel} is positive due to the direction of the twist of the field lines (pointing upward). All the above support the scenario of the untwisting motion of the blowout jet during its emission. It is likely that the process of untwisting of the erupting plasma in our experiment is similar to the sweeping-magnetic-twist mechanism for the acceleration of solar jets (Shibata & Uchida 1985). In the present simulation, it is apparent that the twist stored in the erupting flux is released along the direction of the open, reconnected field lines. Also, the direction of the untwisting motion (e.g., as shown in Figure 4) is consistent with the relative orientation of the two field components (i.e., the erupting and the open field) that reconnect. Thus, it is more likely that the untwisting is related to the propagation of the twist along the reconnected field lines instead of the twist itself.

To study the plasma motion of the jet in relation to the direction of the ambient field, we calculate $I_1 = \int \rho^2 dy$ at various heights, from $z \approx 12 \text{ Mm}$ to $z \approx 27.5 \text{ Mm}$. More precisely, we take a horizontal slit at each height and we plot the above quantity, I_1 , during the evolution of the system. The left panel in Figure 5 shows I_1 at $z = 17 \text{ Mm}$, $z = 21.5 \text{ Mm}$, and $z = 25 \text{ Mm}$. We trace plasma of temperature between $6 \times 10^5 \text{ K}$ and $1.2 \times 10^6 \text{ K}$. We find that first the plasma jet is moving toward the positive x direction, then to the negative x direction, and then back toward the positive x direction. This indicates that the jet undergoes an oscillatory motion during its emission. A similar result, in the context of helical jets, has been reported by Liu et al. (2009). The transverse oscillation velocity amplitude of the jet is $\approx 7 \text{ km s}^{-1}$. Figures 5(a)–(c) shows that the oscillation propagates from low to larger heights. Since the jet is considered to be emitted along the ambient magnetic field, this is an evidence of a wave propagating toward the outer solar atmosphere.

To measure the velocity of the propagating wave, we use the height–time diagrams in panels (a)–(c) and we plot the running difference of each panel, as shown in panels (d)–(f). Then, we find the first (i.e., along time) local maximum in each running difference plot, which corresponds to the first profound increment of plasma density at that height. It is likely that this point is located at the front of the heavy plasma distribution along the jet. The height–time profile of this front is then plotted in Figure 5(g) and by calculating the gradient, we obtain the propagating speed (Figure 5(h)). The latter is comparable to the local Alfvén speed, which suggests that we are witnessing the propagation of a torsional Alfvén wave during the untwisting motion of the blowout jet. We find that the above result is generic: all the blowout jets in our experiment show transverse oscillatory motions and encompass propagating Alfvén waves along their broad outflow streams.

3.1. Discussion

The initial ambient open field in our experiment has a field strength of approximately 3 G. The density in the high corona is approximately $10^{-16} \text{ g cm}^{-3}$. Therefore, our simulation is mimicking the emission of jets in a coronal hole environment. We have calculated that a considerable amount of Poynting energy flux ($\mathcal{O}(10^5) \text{ erg s}^{-1} \text{ cm}^{-2}$) leaves the high corona in every blowout jet emission. Also, the mass deposition in the corona increases by a factor of 2–4 during the blowout ejecta. It is likely that the torsional Alfvén waves transport the emitted flux from the low corona to the outer atmosphere and load the open field with mass. This implies that blowout jets may play a significant role in driving the solar wind.

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REFERENCES

- Alexander, D., & Fletcher, L. 1999, *SoPh*, **190**, 167
 Arber, T., Longbottom, A., Gerrard, C., & Milne, A. 2001, *JCoPh*, **171**, 151

- Archontis, V., & Hood, A. W. 2013, [ApJL](#), **769**, L21
- Archontis, V., & Török, T. 2008, [A&A](#), **492**, L35
- Canfield, R. C., Peardon, K. P., Leka, K. D., et al. 1996, [ApJ](#), **464**, 1016
- Cirtain, J. W., Golub, L., Lundquist, L., et al. 2007, [Sci](#), **318**, 1580
- Fang, F., Fan, Y., & McIntosh, S. W. 2014, [ApJL](#), **789**, L19
- Heyvaerts, J., Priest, E. R., & Rust, D. M. 1977, [ApJ](#), **216**, 123
- Liu, W.L., Berger, T. E., Title, A. M., & Tarbell, T. D. 2009, [ApJL](#), **707**, L37
- Moore, R. L., Cirtain, J. W., Sterling, A. C., & Falconer, D. A. 2010, [ApJ](#), **720**, 757
- Moore, R. L., Sterling, A. C., Falconer, D. A., & Robe, D. 2013, [ApJ](#), **769**, 134
- Moreno-Insertis, F., & Galsgaard, K. 2013, [ApJ](#), **771**, 20
- Nishizuka, N., Shimizu, M., Nakamura, T., et al. 2008, [ApJL](#), **683**, L83
- Schmieder, B., Shibata, K., van Driel-Gesztelyi, L., & Freeland, S. 1995, *SoPh*, **156**, 245
- Shibata, K., Ishido, Y., Acton, L. W., et al. 1992, [PASJ](#), **44**, L173
- Shibata, K., & Uchida, Y. 1985, *PASJ*, **37**, 31
- Shimojo, M., Hashimoto, S., Shibata, K., et al. 1996, *PASJ*, **48**, 123
- Shen, Y., Liu, Y., Su, J., & Deng, Y. 2012, [ApJ](#), **745**, 164
- Sterling, A. C., Harra, L. K., & Moore, R. L. 2010, [ApJ](#), **722**, 1644
- Wang, Y.-M., Sheeley, N. R., Socker, D. G., Jr., et al. 1998, [ApJ](#), **508**, 899