# *SPITZER* OBSERVATIONS OF THE THERMAL EMISSION FROM WASP-43b

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## ABSTRACT

WASP-43b is one of the closest-orbiting hot Jupiters, with a semimajor axis of $a = 0.01526 \pm 0.00018$ AU and a period of only 0.81 days. However, it orbits one of the coolest stars with a hot Jupiter ($T_* = 4520 \pm 120$ K), giving the planet a modest equilibrium temperature of $T_{eq} = 1440 \pm 40$ K, assuming zero Bond albedo and uniform planetary energy redistribution. The eclipse depths and brightness temperatures from our jointly fit model are $0.347\% \pm 0.013\%$ and $1670 \pm 23$ K at 3.6 $\mu$m and $0.382\% \pm 0.015\%$ and $1514 \pm 25$ K at 4.5 $\mu$m. The eclipse timings improved the estimate of the orbital period, $P$, by a factor of three ($P = 0.81347436 \pm 1.4 \times 10^{-7}$ days) and put an upper limit on the eccentricity ($e = 0.010^{+0.010}_{-0.007}$). We use our *Spitzer* eclipse depths along with four previously reported ground-based photometric observations in the near-infrared to constrain the atmospheric properties of WASP-43b. The data rule out a strong thermal inversion in the dayside atmosphere of WASP-43b. Model atmospheres with no thermal inversions and fiducial oxygen-rich compositions are able to explain all the available data. However, a wide range of metallicities and C/O ratios can explain the data. The data suggest low day–night energy redistribution in the planet, consistent with previous studies, with a nominal upper limit of about 35% for the fraction of energy incident on the dayside that is redistributed to the nightside.

*Key words:* eclipses – planets and satellites: atmospheres – planets and satellites: individual (WASP-43b) – techniques: photometric

*Online-only material:* color figures, supplemental data

## 1. INTRODUCTION

Our knowledge of exoplanetary systems is rapidly improving. Recent *Kepler* results (Borucki et al. 2011; Batalha & Kepler Team 2012) have shown a striking increase in detections of the smallest candidates, and the planet candidate lists now show that hot Jupiters are much less common than planets smaller than Neptune. However, nearly all *Kepler* candidates are too small, cold, or distant for atmospheric characterization, except the nearby hot Jupiters. Their host stars, bright enough for radial velocity (RV) measurements, subject these planets to a strong irradiating flux, which governs their atmospheric chemistry and dynamics. Their large sizes and large-scale heights (e.g., Showman & Guillot 2002) give the signal-to-noise ratio (S/N) needed for basic atmospheric characterization.

The most common technique for observing hot Jupiters and characterizing their dayside atmospheres is secondary eclipse photometry (e.g., Fraine et al. 2013; Crossfield et al. 2012; Todorov et al. 2012; Desert et al. 2012; Deming et al. 2011; Beerer et al. 2011; Demory et al. 2007). During secondary eclipse, when the planet passes behind its star, we see a dip in integrated flux proportional to the planet-to-star flux ratio, or usually 0.02%–0.5% in the *Spitzer Space Telescope* infrared wavelengths, where the signal is strongest. This dip is much lower at wavelengths accessible from the ground or from the *Hubble Space Telescope*. Techniques such as phase curve measurement (Knutson et al. 2009, 2012; Lewis et al. 2013; Cowan et al. 2012a, 2012b; Crossfield et al. 2010), transmission spectroscopy (Deming et al. 2013; Gibson et al. 2012; Berta

et al. 2012), and ingress–egress mapping (de Wit et al. 2012; Majeau et al. 2012) can reveal more than a secondary eclipse but are available for only a small number of high-S/N planets.

A secondary eclipse observed in one bandpass places a weak constraint on an exoplanet's temperature near the average altitude of optical depth unity over that bandpass. Multiple wavelengths constrain the planet's dayside spectrum, potentially yielding insight into the atmospheric composition and temperature structure. Different wavelengths probe different atmospheric levels and can be combined into a broadband spectrum for further atmospheric modeling (e.g., Madhusudhan & Seager 2009; Stevenson et al. 2010). Infrared observations are specifically valuable because the most abundant chemical species in planetary atmospheres (aside from $H_2$ and He), such as $H_2O$, $CO$, $CO_2$, and $CH_4$, have significant absorption and emission features at these wavelengths (e.g., Madhusudhan & Seager 2010). Constraints on chemical composition and thermal structure are important for both further atmospheric modeling (e.g., Showman et al. 2009) and studies of the planet's formation. Several recent studies have shown that the atmospheric C/O ratios of giant planets can be significantly different from those of their host stars because of, for example, the formation location of the planet (for exoplanets, see Öberg et al. 2011 and Madhusudhan et al. 2011b, and for Jupiter, see Lodders 2004 and Mousis et al. 2012).

Secondary eclipse observations also provide insight into the exoplanet's orbit. Measuring the time of the secondary eclipse relative to the time of transit can establish an upper limit on orbital eccentricity, $e$, and constrain the argument

of periapsis, $\omega$, independently of RV measurements. Orbital eccentricity is important in dynamical studies and in calculating irradiation levels. Apsidal precession can also be constrained by eclipse timing and can be used to reveal the degree of central concentration of mass in the planetary interior (Ragozzine & Wolf 2009; Campo et al. 2011; López-Morales et al. 2010).

WASP-43b was first detected by the Wide-Angle Search for Planets (WASP) team (Hellier et al. 2011) in 2009 and 2010 from the WASP-South and WASP-North observatories. The WASP team also performed follow-up measurements with the CORALIE spectrograph, the TRAPPIST telescope, and EulerCAM in 2010 December. These observations revealed a planet with a mass of $M_p = 1.78$ Jupiter masses ($M_J$) and a radius of $R_p = 0.93$ Jupiter radii ($R_J$), transiting one of the coldest stars to host a hot Jupiter (type K7V, $T_* = 4400 \pm 200$ K). They found the planet to have an exceptionally short orbital period of 0.81 days and a semimajor axis of only 0.0142 AU, assuming the host star has a mass of $M_* = 0.58 \pm 0.05\,M_\odot$. The planet's orbital eccentricity was constrained by the radial velocity and transit data to $e < 0.04$ at $3\sigma$. Spectroscopic measurements of the star revealed a surface gravity of $\log(g) = 4.5 \pm 0.2$ (cgs) and a projected stellar rotation velocity of $v_* \sin(i) = 4.0 \pm 0.4$ km s$^{-1}$, where $i$ is the inclination of the star's pole to the line of sight. Strong Ca H and K emission indicates that the star is active. The estimated age of the star is $400^{+200}_{-100}$ Myr.

For low-mass stars like WASP-43, there are notable discrepancies (Berger et al. 2006) between interferometrically determined radii and radii calculated in evolutionary models (i.e., Chabrier & Baraffe 1997; Siess et al. 1997). Boyajian et al. (2012) presented high-precision interferometric diameter measurements of 33 late-type K and M stars. They found that evolutionary models overpredict temperatures for stars with temperatures below 5000 K by $\sim$3% and underpredict radii for stars with radii below 0.7 $R_\odot$ by $\sim$5%. Their Table 11 lists an average temperature and radius for each spectral type in the sample, suggesting that WASP-43, with its measured temperature of $4520 \pm 120$ K, is likely a K4 star rather than a K7 as reported by Hellier et al. (2011).

Gillon et al. (2012) analyzed 23 transit light curves, 7 occultations, and 8 new measurements of the star's RV, observed during 2010 and 2011 with TRAPPIST, the Very Large Telescope (VLT), and EulerCAM. They refined eccentricity to $e = 0.0035 \pm 0.0043$ and placed a $3\sigma$ upper limit of 0.0298 using all data simultaneously. They also improved the parameters of the system significantly ($M_p = 2.034 \pm 0.052\,M_J$, $R_p = 1.036 \pm 0.019\,R_J$), refined stellar parameters ($T_{\rm eff} = 4520 \pm 120$ K, $M_* = 0.717 \pm 0.025\,M_\odot$, $R_* = 0.667 \pm 0.011\,R_\odot$), and constrained stellar density ($\rho_* = 2.41 \pm 0.08\,\rho_\odot$). They also confirmed that the observed variability of the transit parameters can be attributed to the variability of the star itself (consistent with Hellier et al. 2011). In addition, they detected the planet's thermal emission at 1.19 $\mu$m and 2.09 $\mu$m, and the atmospheric models of Fortney et al. (2005, 2008) used to infer poor redistribution of heat to the night side and an atmosphere without a thermal inversion.

In this paper, we present two secondary eclipses, observed at 3.6 and 4.5 $\mu$m with the Infrared Array Camera (IRAC; Fazio et al. 2004) on the *Spitzer Space Telescope*, which further constrain the dayside emission of the planet and improve the orbital parameters of the system. We combine our *Spitzer* eclipse depth measurements with previously reported measurements of thermal emission in the near-infrared from Gillon et al. (2012)

**Table 1**
Observation Information

| Channel | Observation Date | Start Time (MJD$_{\rm UTC}$) | Duration (s) | Exposure Time (s) | Number of Frames |
|---------|------------------|------------------------------|--------------|-------------------|------------------|
| Ch1 | 2011 Jul 30 | 2455772.6845 | 21421 | 2 | 10496 |
| Ch2 | 2011 Jul 29 | 2455771.8505 | 21421 | 2 | 10496 |

and Wang et al. (2013) to constrain the atmosphere's energy distribution and thermal profile by using the retrieval method of Madhusudhan & Seager (2009) as subsequently developed.

The following sections present our observations (Section 2); discuss photometric analysis (Section 3); explain specific steps taken to arrive at the fits for each observation and a joint fit (Section 4); give improved constraints on the orbital parameters based on available RV, eclipse, and transit data (Section 5); discuss implications for the planetary emission spectrum and planetary composition (Section 6); state our conclusions (Section 7); and, in the Appendix, supply the full set of system parameters from our own work and previous work. The electronic attachment to this paper includes archival light curve files in FITS ASCII table.

## 2. OBSERVATIONS

We observed two secondary eclipses of WASP-43b with the *Spitzer* IRAC camera in subarray mode (Program ID 70084). A sufficiently long baseline (Figure 1) was monitored before the eclipses, providing good sampling of all *Spitzer* systematics. To minimize intrapixel variability, each target had fixed pointing. We used the Basic Calibrated Data (BCD) from *Spitzer*'s data pipeline, version S.18.18.0. Basic observational information is given in Table 1.

## 3. SECONDARY ECLIPSE ANALYSIS—METHODOLOGY

Exoplanet characterization requires high precision, since the planets' inherently weak signals are weaker than the systematics. In addition, *Spitzer*'s systematics lack full physical characterizations. We have developed a modular pipeline, Photometry for Orbits, Eclipses, and Transits (POET), that implements a wide variety of treatments of systematics and uses Bayesian methods to explore the parameter space and information criteria for model choice. The POET pipeline is documented in our previous papers (Stevenson et al. 2010, 2012; Campo et al. 2011; Nymeyer et al. 2011; Blecic et al. 2013; Cubillos et al. 2013), so we give here just a brief overview of the specific procedures used in this analysis.

The pipeline uses *Spitzer*-supplied BCD frames to produce systematics-corrected light curves and parameter and uncertainty estimates, routinely achieving 85% of the photon S/N limit or better. Initially, POET masks pixels according to *Spitzer's* permanent bad pixel masks, and then it additionally flags bad pixels (energetic particle hits, etc.) by grouping sets of 64 frames and performing a two-iteration, $4\sigma$ rejection at each pixel location. Image centers with 0.01 pixel accuracy come from testing a variety of centering routines (Stevenson et al. 2010, Supplementary Information). Subpixel 5× interpolated aperture photometry (Harrington et al. 2007) produces the light curves. We omit frames with bad pixels in the photometry aperture. The background, subtracted before photometry, is an average of good pixels within an annulus centered on the star in each frame.
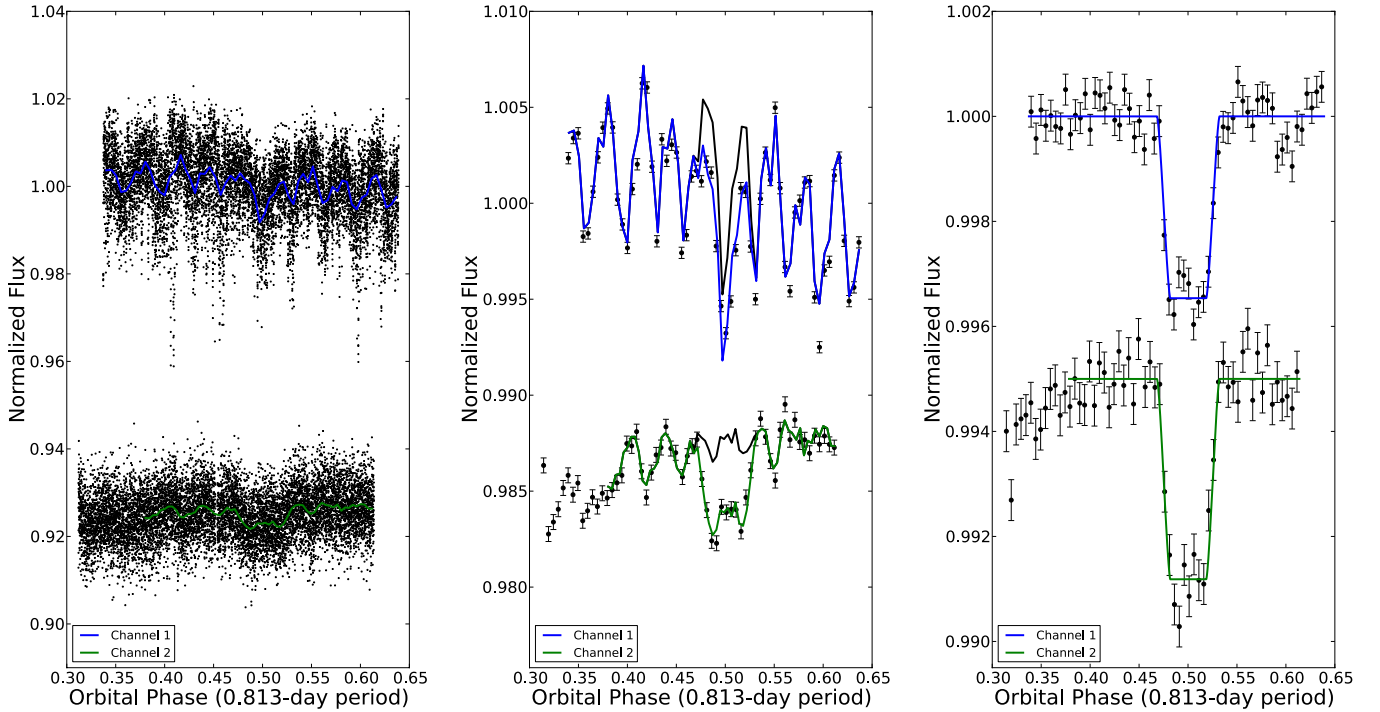
**Figure 1.** Raw (left), binned (center, 60 points per bin, with $1\sigma$ error bars), and systematics-corrected (right) secondary-eclipse light curves of WASP-43b at 3.6 and 4.5 μm. The results are normalized to the system flux and shifted vertically for comparison. Note the different vertical scales used in each panel. The colored lines are best-fit models. The black curves in panel 2 are models without eclipses. As seen in the binned plots of channel 2, a ramp model is not needed to correct for the time-dependent systematic if initial data points affected by pointing drift are clipped (see Section 3).
(A color version of this figure is available in the online journal.)

Detector systematics vary by channel and can have both temporal (detector ramp) and spatial (intrapixel variability) components. At 3.6 and 4.5 μm, intrapixel sensitivity variation is the dominant effect (Charbonneau et al. 2005), so accurate centering at the 0.01 pixel level is critical. We fit this systematic with a Bilinearly Interpolated Subpixel Sensitivity (BLISS) mapping technique, following Stevenson et al. (2012), including the method to optimize the bin sizes and the minimum number of data points per bin.

At 8.0 and 16 μm, there is temporal variability, attributed to charge trapping (Knutson et al. 2009). Weak temporal dependencies can also occur at 3.6 and 4.5 μm (Reach et al. 2005; Charbonneau et al. 2005; Campo et al. 2011; Demory et al. 2011; Blecic et al. 2013), while weak spatial variability has been seen at 5.8 and 8.0 μm (Stevenson et al. 2012; Anderson et al. 2011). Thus, we consider both systematics in all channels when determining the best-fit model.

We fit the model components simultaneously using a Mandel & Agol (2002) eclipse, $E(t)$; the time-dependent detector ramp model, $R(t)$; and the BLISS map, $M(x, y)$:

$$F(x, y, t) = F_s\, R(t)\, M(x, y)\, E(t), \qquad (1)$$

where $F(x, y, t)$ is the aperture photometry flux and $F_s$ is the constant system flux outside of the eclipse.

To choose the best systematics models, we analyze dozens of model combinations and use goodness-of-fit criteria (Campo et al. 2011). For a given channel, we first vary the photometric aperture size and the number of initial data points that we exclude because of instrument settling, and then we test different ramp models and bin sizes for the intrapixel model. To choose the best aperture size and the number of initial points dropped

during instrument settling, we minimize the standard deviation of the normalized residuals (SDNR). Ignoring data points from the beginning of the observation is a common procedure (Knutson et al. 2011) when searching for the best-fitting ramp. We remove the smallest number of points consistent with the minimal SDNR (see each channel analysis for the number of points discarded).

Once we have found the best dataset in this way, we compare different ramp models by applying the Bayesian Information Criterion:

$$\mathrm{BIC} = \chi^2 + k\ln N, \qquad (2)$$

where $N$ is the number of data points. The best model minimizes the chosen criterion. The level of correlation in the photometric residuals is also considered by plotting root-mean-squared (rms) model residuals versus bin size (time interval, Pont et al. 2006; Winn et al. 2008; Campo et al. 2011) and comparing this to the theoretical $1/\sqrt{2N}$ rms scaling (Blecic et al. 2013 explains the factor of 2). Sometimes, we prefer less-correlated models with insignificantly poorer BIC values.

We explore the phase space and estimate errors by using a Markov-chain Monte Carlo (MCMC) routine following the Metropolis–Hastings random-walk algorithm, which uses independent Gaussian proposal distributions for each parameter with widths chosen to give an acceptance rate of 30%–60%. Each MCMC model fit begins with the Levenberg–Marquardt algorithm (least-squares minimization). We use an informative prior (e.g., Gelman 2002) taken from other work on parameters that are more tightly constrained than what our fits can achieve. In this work, those are ingress and egress times, which are not well sampled by our observation. All other parameters have flat priors and are free parameters of the MCMC. For each channel,
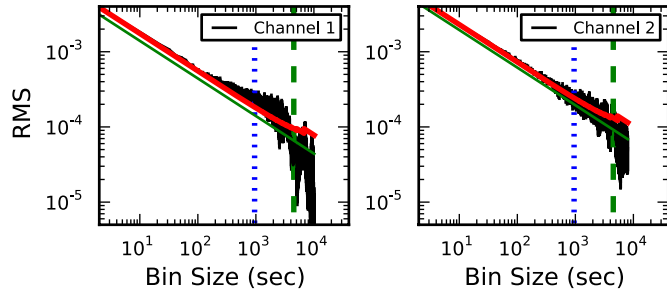
**Figure 2.** Correlations of the residuals for the two secondary eclipses of WASP-43b, following Pont et al. (2006). The black line represents the rms residual flux vs. bin size. The red line shows the predicted standard error scaling for Gaussian noise. The green line shows the theoretical photon noise limit (observed S/N is 80.3% and 85% of the photon-limited S/N for channel 1 and channel 2, respectively; see Section 4.3). The black vertical lines at each bin size depict $1\sigma$ uncertainties on the rms residuals ($rms/\sqrt{2N}$, where N is the number of bins). The dotted vertical blue line indicates the ingress/egress timescale, and the dashed vertical green line indicates the eclipse duration timescale. Large excesses of several $\sigma$ above the red line would indicate correlated noise at that bin size. Inclusion of $1\sigma$ uncertainties shows no noise correlation on the timescales between the ingress/egress duration and the eclipse duration.

(A color version of this figure is available in the online journal.)

they are listed in Section 4.3. For orbital analysis, they are listed in Section 5. We then run enough MCMC iterations to satisfy the Gelman & Rubin (1992) convergence test. After every run, we assess convergence by examining plots of the parameter traces, pairwise correlations, autocorrelations, marginal posteriors, best-fitting model, and systematics-corrected best-fitting model. The final fit is obtained from the simultaneous run of all data sets, sharing parameters such as the eclipse midpoint and duration among some or all datasets.

We report the times of our secondary eclipses in both BJD$_{\rm UTC}$ (Coordinated Universal Time) and BJD$_{\rm TT}$ (BJD$_{\rm TDB}$, Barycentric Dynamical Time), calculated using the Jet Propulsion Laboratory (JPL) Horizons system and following Eastman et al. (2010).

## 4. SECONDARY ECLIPSE ANALYSIS—FIT DETAILS

Light curves for both channels were extracted using every aperture radius from 2.00 to 4.50 pixels, in 0.25 pixel increments. We tested three centering routines, center of light, two-dimensional (2D) Gaussian fit, and least asymmetry (see Supplementary Information of Stevenson et al. 2010 and Lust et al. 2013). A 2D Gaussian fit found the most consistent stellar centers. We estimated the background flux by using an annulus of 7–15 pixels from the center of the star for both channels. For the secondary eclipse ingress and egress time, we used a Bayesian prior ($t_{2-1} = 950.5 \pm 145.5$ s), calculated from unpublished WASP photometric and RV data.

Figure 1 shows our systematics-corrected, best-fit light curve models. Figure 2 presents the scaling of the rms model residuals versus bin size for both channels, which shows no significant time correlation in the residuals.

### 4.1. Channel 1–3.6 μm

The most prominent systematic in this *Spitzer* channel is the intrapixel effect. The best BLISS-map bin size is 0.006 pixels when we exclude bins with less than four measurements. The ramp and eclipse models fit without removing initial data points. The smallest value of BIC reveals that the best ramp model is quadratic; this model is $1.2 \times 10^{30}$ times more probable than
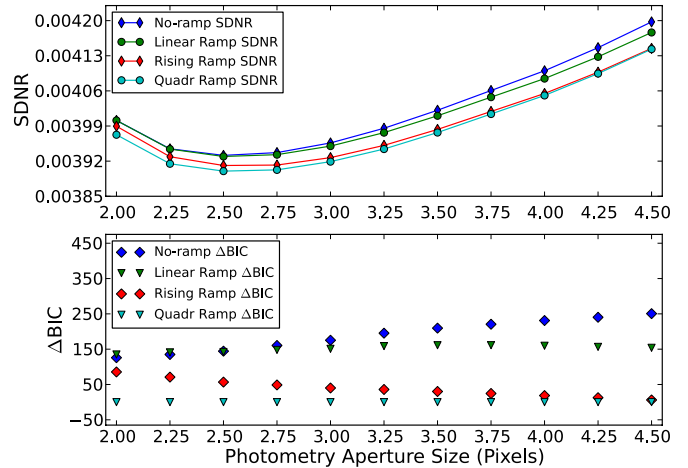


**Figure 3.** Channel 1 comparison between different ramp models. The plots show SDNR vs. aperture size and $\Delta$BIC vs. aperture size. A lower SDNR value indicates a better model fit. The lowest SDNR value marks the best aperture size (2.50 pixels). A lower $\Delta$BIC value at the best aperture size indicates which ramp model is the best (quadratic ramp model, green triangles).

(A color version of this figure is available in the online journal.)

**Table 2**
Channel 1 Ramp Models

| Ramp Model | SDNR | $\Delta$BIC | Eclipse Depth (%) |
|---|---|---|---|
| Quadratic | 0.0039001 | 0.0 | $0.344 \pm 0.013$ |
| Rising | 0.0039113 | 56.8 | $0.292 \pm 0.012$ |
| No-Ramp | 0.0039315 | 144.4 | $0.268 \pm 0.012$ |
| Linear | 0.0039293 | 142.1 | $0.270 \pm 0.012$ |

the next-best (linear) model. Table 2 lists the best ramp models, comparing their SDNR, BIC values, and eclipse depths.

Figure 3 shows a comparison between the best ramp models and their SDNR and BIC values through all aperture sizes, indicating which aperture size is the best and which model has the lowest BIC value.

Photometry generates consistent eclipse depths for all tested apertures, with the lowest SDNR at an aperture radius of 2.50 pixels (Figure 4).

### 4.2. Channel 2–4.5 μm

In this channel, we noticed an upward trend in flux at the beginning of the observation, possibly due to telescope settling, which we do not model. We clipped 2300 initial data points (∼38 minutes of observation), the smallest number of points consistent with the minimal SDNR. The 2.50 pixel aperture radius minimizes SDNR (Figure 5).

To remove intrapixel variability, the BLISS bin size is 0.016 pixels, ignoring bins with less than four points. The lowest BIC value corresponds to the model without a ramp (Table 3), which is 78 times more probable than the linear model.

We tested the dependence of eclipse depth on aperture radius, showing that they are all well within $1\sigma$ of each other, to validate the consistency of our models (Figure 6).

### 4.3. Joint Fit

To improve accuracy, we share the eclipse width (duration), eclipse midpoint phases, and ingress and egress times in a joint fit of both data sets. Table 4 indicates which parameters are free,
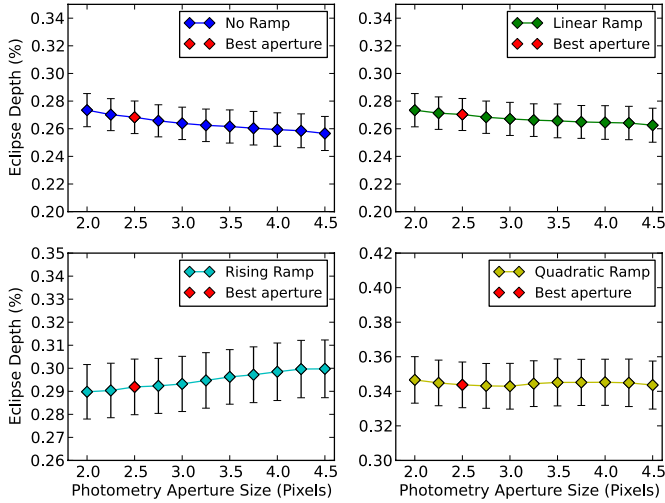
**Figure 4.** Best-fit eclipse depths as a function of photometry aperture size for channel 1. The four best ramp models are plotted. The red point indicates the best aperture size for that channel. The eclipse depth uncertainties are the result of $10^5$ MCMC iterations. The trend shows insignificant dependence of eclipse depth on aperture size (much less than $1\sigma$).

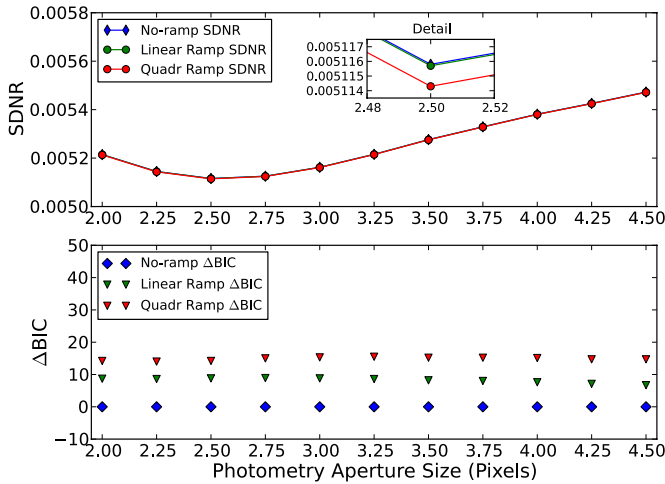(A color version of this figure is available in the online journal.)



**Figure 5.** Channel 2 comparison between different ramp models. Plot shows SDNR vs. aperture size and $\Delta$BIC vs. aperture size. A lower SDNR value indicates a better model fit. The lowest SDNR value marks the best aperture size (2.50 pixels). Lower BIC values at the best aperture size indicate better models (best: no-ramp model, blue diamonds). The inset shows separation in SDNR for different ramp models at the best aperture size.

(A color version of this figure is available in the online journal.)

shared, or have informative priors. The best ramp models and the best aperture sizes from the separate channel analyses are used in the joint fit. To produce the best joint-model fits, we iterated MCMC until the Gelman & Rubin (1992) diagnostics for all parameters dropped below 1%, which happened after $10^5$ iterations.

The best joint-model fit parameters are in Table 4. Files containing the light curves, best model fits, centering data, photometry, etc., are included as electronic supplements to this article. The eclipse midpoint time is further used for the subsequent orbital analysis, and the eclipse depths are used for the atmospheric analysis.
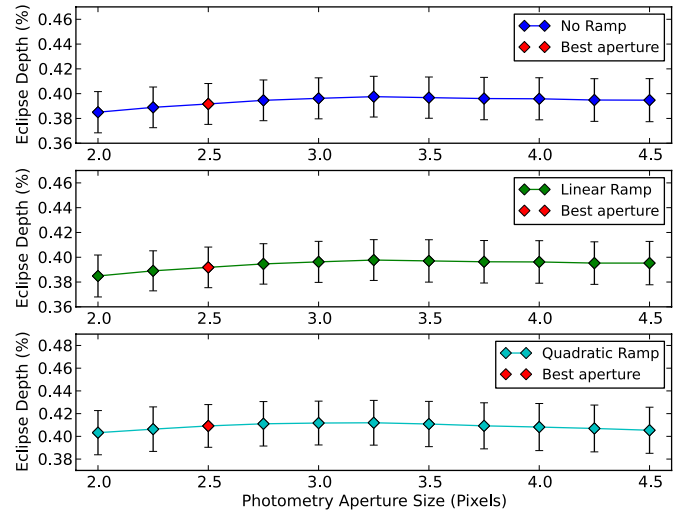


**Figure 6.** Best-fit eclipse depths as a function of photometry aperture size for channel 2. The three best ramp models are plotted. The red point indicates the best aperture size for that channel. The eclipse depth uncertainties are the result of $10^5$ MCMC iterations. The trend shows negligible dependence of eclipse depth on aperture size (much less than $1\sigma$).

(A color version of this figure is available in the online journal.)

**Table 3**
Channel 2 Ramp Models

| Ramp Model | SDNR | $\Delta$BIC | Eclipse Depth (%) |
|---|---|---|---|
| No-Ramp | 0.0051158 | 0 | $0.392 \pm 0.016$ |
| Linear | 0.0051157 | 8.8 | $0.392 \pm 0.016$ |
| Quadratic | 0.0051143 | 14.2 | $0.409 \pm 0.019$ |

## 5. ORBIT

The eclipse midpoint (after a 15.2 s correction for the eclipse transit light-time) has a phase of $0.5001 \pm 0.0004$, so $e \cos \omega = 0.0001 \pm 0.0006$, or a $3\sigma$ upper limit of $|e \cos \omega| < 0.0018$, consistent with a circular orbit.

To improve the orbit solution further, we combined data from our observations with data from a variety of sources (see Table 5). Transit midpoint times were taken from Hellier et al. (2011) and Gillon et al. (2012), and amateur observations were listed in the Exoplanet Transit Database (see Table 5). We used *CORALIE* RV observations published by Hellier et al. (2011) and Gillon et al. (2012). No RV points analyzed were gathered during transit. We subtracted 15.2 s from the eclipse midpoint to correct for light-travel time across the orbit. We corrected all points to TDB if this was not already done (Eastman et al. 2010). We converted the amateur data from HJD to BJD, putting all times in a consistent BJD$_{TDB}$ format. There were 49 transit points, 23 RV points, and 1 effective eclipse observation. We fit all of these data simultaneously, as described by Campo et al. (2011). The free parameters in this fit were $e \sin \omega$; $e \cos \omega$; the period, $P$; the reference transit midpoint time, $T_0$; the RV semi-amplitude, $K$; and the RV offset, $\gamma$. The addition of the amateur transit observations improves the uncertainty of $P$ by a factor of nearly five compared with Gillon et al. (2012), reducing it to 13 ms. The fit finds an eccentricity of $0.010^{+0.010}_{-0.007}$, consistent with a circular orbit and expectations for a close-in planet, where eccentricity should be damped by tidal interactions with the host star (Jackson et al. 2008). Table 6 summarizes the fit results.

**Table 4**
Best-fit Joint Eclipse Light Curve Parameters

| Parameter | Channel 1 | Channel 2 |
|---|---|---|
| Array position ($\bar{x}$, pixel) | 14.99 | 14.8 |
| Array position ($\bar{y}$, pixel) | 15.01 | 15.05 |
| Position consistency[a] ($\delta_x$, pixel) | 0.009 | 0.016 |
| Position consistency[a] ($\delta_y$, pixel) | 0.012 | 0.014 |
| Aperture size (pixel) | 2.5 | 2.5 |
| Sky annulus inner radius (pixel) | 7.0 | 7.0 |
| Sky annulus outer radius (pixel) | 15.0 | 15.0 |
| System flux $F_s$ ($\mu$Jy)[b] | 64399.0 ± 5.0 | 37911.0 ± 2.0 |
| Eclipse depth (%)[b] | 0.347 ± 0.013 | 0.382 ± 0.015 |
| Brightness temperature (K) | 1670 ± 23 | 1514 ± 25 |
| Eclipse midpoint (orbits)[b] | 0.4986 ± 0.0004 | 0.4986 ± 0.0004 |
| Eclipse midpoint (BJD$_{UTC}$ −2,450,000) | 5773.3172 ± 0.0003 | 5772.5037 ± 0.0003 |
| Eclipse midpoint (BJD$_{TDB}$ −2,450,000) | 5773.3179 ± 0.0003 | 5772.5045 ± 0.0003 |
| Eclipse duration ($t_{4-1}$, hr)[b] | 1.25 ± 0.02 | 1.25 ± 0.02 |
| Ingress/egress time ($t_{2-1}$, hr)[b] | 0.268 ± 0.018 | 0.268 ± 0.018 |
| Ramp name | quadramp | no-ramp |
| Ramp, quadratic Term[b] | −0.0827 ± 0.0069 | ⋯ |
| Ramp, linear term[b] | −0.0002 ± 0.0005 | ⋯ |
| Intrapixel method | BLISS | BLISS |
| BLISS bin size in $x$ (pixel) | 0.006 | 0.016 |
| BLISS bin size in $y$ (pixel) | 0.006 | 0.016 |
| Minimum number of points per bin | 4 | 4 |
| Total frames | 10496 | 10496 |
| Frames used | 10124 | 8004 |
| Rejected frames (%) | 0.44 | 1.12 |
| Free parameters | 7 | 2 |
| Number of data points in fit | 10124 | 8004 |
| BIC | 18207.3 | 18207.3 |
| SDNR | 0.0039007 | 0.0051167 |
| Photon-limited S/N (%) | 80.3 | 85.0 |

**Notes.**
[a] rms frame-to-frame position difference.
[b] Free parameter in MCMC fit. All priors are flat except $t_{2-1}$, which uses a Gaussian prior of 950.5 ± 145.5 s, calculated from unpublished WASP photometric and RV data. Parameters with identical values and uncertainties are fit jointly.

Because $e \sin \omega$ is a much larger component of the eccentricity than $e \cos \omega$, it is possible that much of this eccentricity signal comes from the effect of the planet raising a tidal bulge on its host star. Arras et al. (2012) predict that the RV semi-amplitude of this effect is 8.9 ms$^{-1}$. Since our model shows that $eK = 5^{+6}_{-3}$ ms$^{-1}$, it is possible that the majority or entirety of the eccentricity signal is due to the tidal bulge interaction, and the true eccentricity is closer to the upper limit derived from the secondary eclipse.

We found that, for a linear ephemeris fit to just the transit timing data, there is considerable scatter in O-C (observed time minus calculated time). The root-mean-square of the stated transit-time uncertainties is 51 s, while the standard deviation of the residuals is 124 s. WASP-43b is close enough to its host star that tidal decay is a significant factor in its evolution, so we attempted to estimate the decay rate by adding a quadratic term to our ephemeris model, following Adams et al. (2010). Our model for the transit ephemeris is now:

$$T_N = T_0 + PN + \delta P \frac{N(N-1)}{2}, \quad (3)$$

where $N$ is the number of orbits elapsed since the epoch $T_0$, $P$ is the orbital period at $T_0$, and $\delta P = \dot{P} P$, where $\dot{P}$ is the short-term rate of change in the orbital period. Fitting this model to the transit data, we find that $\delta P = (-2.5 \pm 0.9) \times 10^{-9}$ days

orbit$^{-2}$, or $\dot{P} = -0.095 \pm 0.036$ s yr$^{-1}$. This is illustrated in Figure 7. This is a nondetection, though the best-fit value is comparable to the value of $-0.060 \pm 0.015$ s yr$^{-1}$ found by Adams et al. (2010) for OGLE-TR-113b.

While the suggestion of a quadratic inspiral existed in the data gathered prior to BJD 2456035, the addition of 17 new amateur observations from BJD 2456250 to BJD 2456440 (11 of which were of sufficiently high data quality for inclusion) put the linear fit within the credible region of the quadratic fit.

Table 7 summarizes all three transit ephemeris models. The first is the linear ephemeris from the fit discussed above. We also fit linear and quadratic models to just the transit data to study any trend in orbital period. These are the second and third fits in Table 7. Their BIC and chi $\chi^2$ are comparable to each other but not to the first fit, since the data sets are not identical.

The BIC values still favor the quadratic (decaying) ephemeris (BIC = 450) over the linear ephemeris (BIC = 452.8) by a probability ratio of $e^{\Delta BIC/2} = 4$. Two lines of reasoning favor the linear ephemeris, however, so we consider the linear ephemeris to be more likely.

First, the inspiral time predicted by the quadratic is extremely short compared to the planet's lifetime. We would need to believe that we are seeing the planet at a very brief and special time in its history to accept the conclusion of inspiral.

**Table 5**
Transit Timing Data

| Mid-transit Time (BJD$_{TDB}$) | Uncertainty | Source[a] | Quality Rating |
|---|---|---|---|
| 2456440.77250 | 0.00037 | Phil Evans | 1 |
| 2456410.67704 | 0.00097 | Robert Majewski | 5 |
| 2456407.42697 | 0.00097 | Enrique Díez Alonso | 3 |
| 2456407.41972 | 0.00102 | Ullrich Dittler | 3 |
| 2456403.34716 | 0.0017 | Jens Jacobsen | 5 |
| 2456401.72885 | 0.00199 | Alex Chassy | 5 |
| 2456387.89785 | 0.00046 | Phil Evans | 2 |
| 2456375.70739 | 0.00158 | Parijat Singh | 4 |
| 2456368.37413 | 0.00076 | Adam Büchner | 3 |
| 2456335.83452 | 0.00113 | Phil Evans | 3 |
| 2456328.51426 | 0.00039 | Juan Lozano de Haro | 2 |
| 2456328.51296 | 0.00181 | Daniel Staab | 4 |
| 2456326.88711 | 0.00072 | Phil Evans | 3 |
| 2456313.87042 | 0.00039 | P. Kehusmaa & C. Harlingten | 2 |
| 2456288.65334 | 0.00087 | Jordi Lopesino | 3 |
| 2456283.77139 | 0.00068 | A. Chapman & N. D. Díaz | 3 |
| 2456250.42005 | 0.00067 | L. Zhang, Q. Pi & A. Zhou | 5 |
| 2456035.66489 | 0.0005 | George Hall | 2 |
| 2456015.3273 | 0.00057 | Martin Zíbar | 2 |
| 2456006.38071 | 0.00101 | František Lomoz | 3 |
| 2456006.3781 | 0.00109 | František Lomoz | 3 |
| 2456001.49859 | 0.00156 | Alfonso Carreño | 3 |
| 2456001.49662 | 0.00035 | Gustavo Muler Schteinman | 1 |
| 2456001.49531 | 0.00019 | Fernand Emering | 1 |
| 2455997.43508 | 0.00086 | René Roy | 3 |
| 2455997.43105 | 0.00051 | Faustino Garcia | 3 |
| 2455997.43008 | 0.0004 | Nicolas Esseiva | 2 |
| 2455997.42981 | 0.00068 | Juanjo Gonzalez | 3 |
| 2455984.41939 | 0.00064 | Ferran Grau Horta | 3 |
| 2455984.41548 | 0.00126 | František Lomoz | 3 |
| 2455984.4149 | 0.00047 | Fabio Martinelli | 2 |
| 2455984.41472 | 0.00071 | František Lomoz | 3 |
| 2455979.534 | 0.00044 | Nicolas Esseiva | 2 |
| 2455979.5335 | 0.0004 | Juanjo Gonzalez | 2 |
| 2455957.57296 | 0.00122 | František Lomoz | 4 |
| 2455944.55468 | 0.00106 | Roy René | 3 |
| 2455940.48744 | 0.0005 | Anthony Ayiomamitis | 2 |
| 2455939.67475 | 0.00052 | Ramon Naves | 2 |
| 2455933.16473 | 0.00025 | Peter Starr | 1 |
| 2455686.68399 | 0.0008 | Stan Shadick | 3 |
| 2455682.61364 | 0.00039 | Tanya Dax, Stacy Irwin | 5 |

**Notes.** [a] The TRansiting ExoplanetS and CAndidates group (TRESCA, http://var2.astro.cz/EN/tresca/index.php) supply their data to the Exoplanet Transit Database (ETD, http://var2.astro.cz/ETD/), which performs the uniform transit analysis described by Poddaný et al. (2010). The ETD web site provided the numbers in this table, which were converted from HJD (UTC) to BJD (TDB).

**Table 6**
Eccentric Orbital Model

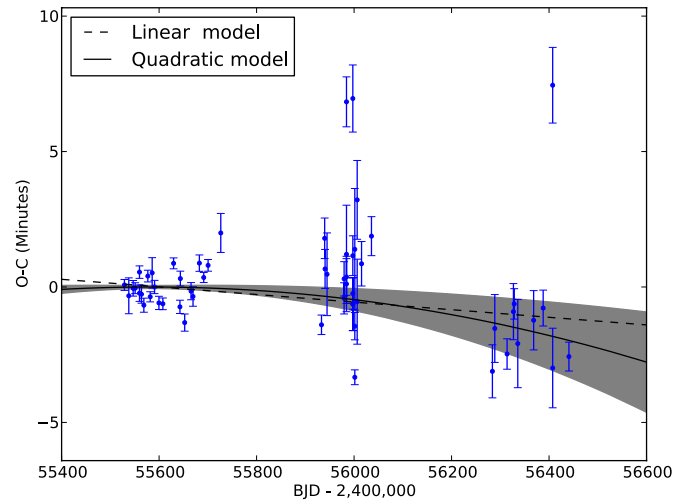| Parameter | Value |
|---|---|
| $e \sin \omega$[a] | $-0.010 \pm 0.011$ |
| $e \cos \omega$[a] | $-0.0003 \pm 0.0006$ |
| $e$ | $0.010^{+0.010}_{-0.007}$ |
| $\omega$ (°) | $-88^{+5}_{-9}$ |
| $P$ (days)[a] | $P = 0.81347436 \pm 1.4 \times 10^{-7}$ |
| $T_0$ (BJD$_{TDB}$)[a] | $2455528.86857 \pm 0.00005$ |
| $K$ (m s$^{-1}$)[a] | $549 \pm 6$ |
| $\gamma$ (m s$^{-1}$)[a] | $-3595 \pm 4$ |
| $\chi^2$ | 458 |

**Note.** [a] Free parameter in MCMC fit.



**Figure 7.** O-C diagram for transit observations of WASP-43b with respect to the linear terms in the best-fit quadratic ephemeris. The quadratic ephemeris is shown with the $1\sigma$ prediction uncertainty in grey. The best-fit linear ephemeris is also shown as a dashed line. Only points with an Exoplanet Transit Database quality rating of 3 or better are shown here and used in this analysis.
(A color version of this figure is available in the online journal.)

Second, the rms scatter about the quadratic ephemeris is 121 s, much larger than the 51 s typical transit time uncertainty and not much different from the linear result. The reduced $\chi^2$ for the quadratic fit is ∼8, so much of the residual scatter is unexplained by either model. Possible explanations include stellar activity, transit timing variations (TTVs), or problems in data processing or reporting. Mistakes in the time corrections for heterogeneous transit data are unlikely because the Exoplanet Transit Database indicates that all amateur observations were submitted in UTC, while the professional data were unambiguous in their use of TDB. While it is possible that uncertainties for certain sets of transit data points may have been underestimated, the data come from many amateur and professional sources, and all would have had to make such errors. Since WASP-43 is an early K-type star, a likely explanation for the scatter is stellar activity.

Effects like TTVs or tidal infall could still contribute to the scatter, so we present related calculations below, mainly as motivation and background for future studies.

From the measured period change of $\dot{P} = -0.095 \pm 0.036$ s yr$^{-1}$, we adopt a $3\sigma$ upper limit, $|\dot{P}| < 0.129$ s yr$^{-1}$. For WASP-43b, this translates to a maximum change in the semi-major axis of $|\dot{a}| < 1.9 \times 10^{-8}$ AU yr$^{-1}$. The three-sigma upper limit on the period decay also suggests an infall timescale of at least $5 \times 10^5$ yr. Levrard et al. (2009) give a relation for tidal decay, which for synchronous planetary rotation and negligible eccentricity and obliquity reduces to:

$$\frac{1}{a}\frac{da}{dt} = \frac{6}{Q'_\star}\frac{M_p}{M_\star}\left(\frac{R_\star}{a}\right)^5\left(\omega_\star - \frac{2\pi}{P}\right), \qquad (4)$$

where $Q'_\star$ is the ratio of the stellar tidal quality factor to the second-order stellar tidal Love number, $k_2$, and $\omega_\star$ is the stellar rotation rate. The upper limit on the quadratic term implies $Q'_\star > 12{,}000$. This is much lower than the values of $10^5$–$10^{10}$ normally assumed and thus cannot rule out any plausible values. A small value of $Q'_\star$ was also found by Adams et al. (2010).

While the quadratic fit failed to produce a useful upper limit on tidal decay, observations with a longer time baseline may yet

**Table 7**
Ephemeris Solutions

| Parameter | All Data (Transits, Eclipses, RV) Linear Ephemeris | Transits Only Linear Ephemeris | Transits Only Quadratic Ephemeris |
|---|---|---|---|
| $T_0$ (BJD$_{TDB}$) | $2455528.86857 \pm 0.00005$ | $2455581.74439 \pm 0.00004$ | $2455581.74437 \pm 0.00004$ |
| $P$ (days orbit$^{-1}$) | $0.81347436 \pm 1.4 \times 10^{-7}$ | $0.81347450 \pm 1.5 \times 10^{-7}$ | $0.81347530 \pm 3.8 \times 10^{-7}$ |
| $\delta P$ (days orbit$^{-2}$)[a] | $\cdots$ | $\cdots$ | $(-2.5 \pm 0.9) \times 10^{-9}$ |
| $\dot{P}$ (s yr$^{-1}$)[b] | $\cdots$ | $\cdots$ | $-0.095 \pm 0.036$ |
| $\chi^2$ | $\cdots$ | $444.7$ | $437.9$ |
| BIC | $\cdots$ | $452.8$ | $450.0$ |

**Notes.**
[a] $\delta P = \dot{P} P$.
[b] Derived parameter.

find secular changes or TTVs. Until then, the linear ephemeris presented here is the most reliable predictor of future transit times.

## 6. ATMOSPHERE

We modeled the dayside atmosphere of WASP-43b by using the atmospheric modeling and retrieval method of Madhusudhan & Seager (2009, 2010). The model computes line-by-line radiative transfer in a one-dimensional, plane-parallel atmosphere, with constraints of local thermodynamic equilibrium, hydrostatic equilibrium, and global energy balance. The pressure–temperature profile and components of the molecular composition are free parameters of the model, allowing exploration of models with and without thermal inversions and those with oxygen-rich as well as carbon-rich compositions (Madhusudhan 2012).

The model includes all the primary sources of opacity expected in hydrogen-dominated giant-planet atmospheres in the temperature regimes of hot Jupiters, such as WASP-43b. The opacity sources include line-by-line absorption due to $H_2O$, CO, $CH_4$, $CO_2$, and $NH_3$ and collision-induced absorption (CIA) due to $H_2$-$H_2$. We also include hydrocarbons besides $CH_4$, such as HCN and $C_2H_2$, which may be abundant in carbon-rich atmospheres (Madhusudhan et al. 2011b; Kopparapu et al. 2012; Madhusudhan 2012). Since in highly irradiated oxygen-rich atmospheres TiO and VO may be abundant (Fortney et al. 2008), we also include line-by-line absorption due to TiO and VO in regions of the atmosphere where the temperatures exceed the corresponding condensation temperatures. Our molecular line data are from Freedman et al. (2008), R. S. Freedman (2009, private communication), Rothman et al. (2005), Karkoschka & Tomasko (2010), E. Karkoschka (2011, private communication), and Harris et al. (2008). We obtain the $H_2$-$H_2$ CIA opacities from Borysow et al. (1997) and Borysow (2002). The volume mixing ratios of all the molecules are free parameters in the model.

We constrain the thermal structure and composition of WASP-43b by combining our *Spitzer* photometric observations at 3.6 $\mu$m and 4.5 $\mu$m combined with previously reported ground-based narrow-band photometric data from Gillon et al. (2012), obtained using VLT/HAWK-I at 1.19 $\mu$m and 2.09 $\mu$m, and broadband photometric data from Wang et al. (2013) obtained using CFHT/WIRCAM in the (1.6 $\mu$m) and $K_s$ (2.1 $\mu$m) bands. The data also place a joint constraint on the day–night energy redistribution and the Bond albedo by requiring global energy balance, i.e., that the integrated emergent power from
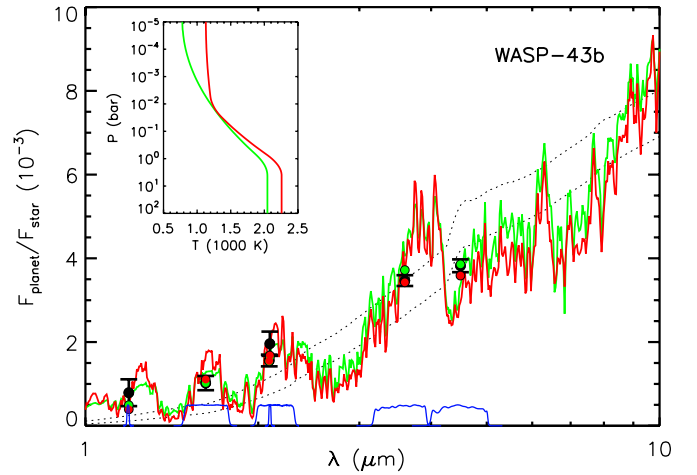


**Figure 8.** Observations and model spectra for dayside thermal emission from WASP-43b. The black filled circles with error bars show our data in *Spitzer* IRAC channels 1 (3.6 $\mu$m) and 2 (4.5 $\mu$m) and previously published ground-based near-infrared data in narrow-band photometry at 1.19 $\mu$m and 2.09 $\mu$m (Gillon et al. 2012) and in broadband photometry at 1.6 $\mu$m and 2.1 $\mu$m (Wang et al. 2013). The solid curves show the model spectra in the main panel, and the corresponding temperature–pressure profiles, with no thermal inversions, in the inset. The green and red curves correspond to models with compositions of nearly solar and 10 × solar metallicity, respectively. Both models fit the data almost equally well. The dashed curves show blackbody spectra corresponding to planetary brightness temperatures of 1670 K and 1514 K, the observed brightness temperatures in the *Spitzer* IRAC channels 1 and 2, respectively.

(A color version of this figure is available in the online journal.)

the planet does not exceed the incident irradiation. Given that the number of model parameters is $\geqslant 10$ (depending on the C/O ratio) and the number of available data points is 6, our goal is to find the regions of model space favored by the data, rather than to determine a unique fit. We explore the model parameter space by using an MCMC routine (for details, see Madhusudhan & Seager 2009, 2010; Madhusudhan et al. 2011a).

The data rule out a strong thermal inversion in the dayside atmosphere of WASP-43b. The data and two model spectra of atmospheres without thermal inversions are shown in Figure 8. The ground-based and *Spitzer* data provide complementary constraints on the atmospheric properties. The ground-based photometric bandpasses, in narrow bands at 1.19 $\mu$m and 2.09 $\mu$m (Gillon et al. 2012) and in broad bands at 1.6 $\mu$m and 2.1 $\mu$m (Wang et al. 2013), span spectral regions of low molecular opacity and hence probe the deep layers of the atmosphere at pressures of $P \sim 1$ bar, beyond which the atmosphere
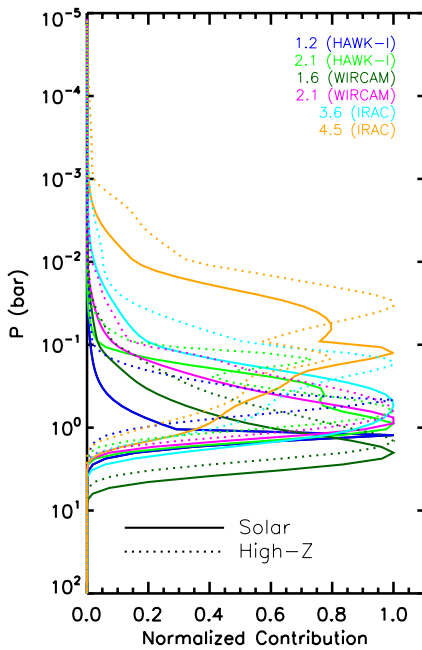
**Figure 9.** Contribution functions for the atmospheric models. The solid (dashed) curves show the contribution functions for the model in Figure 8 with solar (10 × solar, "High-*Z*") composition. The contribution functions are shown in all the bandpasses corresponding to the data, denoted by the central wavelengths in $\mu$m and instruments in parentheses, as shown in the legend.

(A color version of this figure is available in the online journal.)

is optically thick because of collision-induced opacity (the contribution functions are given in Figure 9). Consequently, the brightness temperatures from such ground-based data constrain the isothermal temperature structure of the deep atmosphere (Madhusudhan 2012). On the other hand, the two *Spitzer* data sets show lower brightness temperatures at 3.6 $\mu$m and 4.5 $\mu$m relative to the ground-based data, which is possible only if the temperature structure is decreasing outward in the atmosphere, causing molecular absorption in the *Spitzer* bands. The presence of a strong thermal inversion, on the contrary, would have caused molecular emission leading to higher brightness temperatures in the *Spitzer* bands relative to the ground-based bands. Consequently, the sum total of *Spitzer* and ground-based data rule out a strong thermal inversion in WASP-43b's dayside photosphere.

The molecular composition is less well constrained by the data. Several physically plausible combinations of molecules can explain the absorption in the two *Spitzer* bands (e.g., Madhusudhan & Seager 2010; Madhusudhan 2012). Figure 8 shows two oxygen-rich models in chemical equilibrium, with C/O ratios of 0.5 (solar value) but with different metallicities (nearly solar and ten times solar) and thermal profiles, both of which explain the data almost equally well. In both cases, $H_2O$ absorption in the 3.6 $\mu$m band, and $H_2O$, CO, and $CO_2$ absorption in the 4.5 $\mu$m band explain the *Spitzer* data. The comparable fits demonstrate the degeneracy between the molecular mixing ratios (via the metallicity) and the temperature gradient. Given the current photometric data, the solar metallicity model with a steep temperature profile (green curve) produces almost as good a fit as the higher metallicity model with a shallower temperature profile (red curve). On the other hand, carbon-rich models with C/O $\geqslant$ 1 (e.g., Madhusudhan 2012), with absorption due to $CH_4$, CO, $C_2H_2$, and HCN, could also explain the data. As such, the current data are insufficient to discriminate between

O-rich and C-rich compositions. Thus, new observations are required to obtain more stringent constraints on the chemical composition of WASP-43b. Observations using the *HST* Wide Field Camera 3 in the 1.1–1.8 $\mu$m bandpass can help constrain the $H_2O$ abundance in the atmosphere. As shown in Figure 8, an O-rich composition predicts strong absorption due to $H_2O$ in the WFC3 bandpass, which would be absent in a carbon-rich atmosphere (Madhusudhan et al. 2011a). Similarly, observations in other molecular bands, such as the CO band at 2.3 $\mu$m, can provide constraints on the corresponding molecular mixing ratios.

Our observations provide nominal constraints on the day–night energy redistribution fraction ($f_r$; Madhusudhan & Seager 2009) in WASP-43b. The models shown in Figure 8 have $f_r = 16\%$–20%, assuming zero Bond albedo. Our population of model fits to the combined *Spitzer* and ground-based data allow for up to ∼35% day–night energy redistribution in the planet. Among the acceptable models, those with higher $f_r$ values require cooler lower atmospheres on the dayside and hence predict lower fluxes in the ground-based channels. While such models produce an acceptable fit to all four data points overall, they predict systematically lower fluxes in the ground-based channels, some fitting the ground-based data points only at the $\sigma$ lower error bars. Considering the ground-based points alone, without the *Spitzer* data, would imply a significantly higher continuum flux and correspondingly a significantly lower day–night redistribution in the planet than ∼35%, consistent with the findings of Gillon et al. (2012) and Wang et al. (2013).

The lack of a strong thermal inversion in WASP-43b is not surprising. At an equilibrium temperature of ∼1400 K, the dayside atmosphere of WASP-43b is not expected to host gaseous TiO and VO, which have been proposed to cause thermal inversions (Spiegel et al. 2009; Hubeny et al. 2003; Fortney et al. 2008), though hitherto unknown molecules that could also potentially cause such inversions cannot be ruled out (Zahnle et al. 2009). The lack of a thermal inversion is also consistent with the hypothesis of Knutson et al. (2010), since the host star WASP-43 is known to be active (Hellier et al. 2011).

## 7. CONCLUSIONS

Exoplanet secondary eclipses provide us with a unique way to observe the dayside spectrum of an irradiated planetary atmosphere, where the opacities of the mixture of atmospheric trace molecules determine the thermal structure of the planetary atmosphere.

WASP-43b has a 0.81 day period, making it one of the shortest-period transiting planets. It has a small semimajor axis (0.01526 ± 0.00018 AU, Gillon et al. 2012). WASP-43 is a low-mass star ($M_* = 0.717 \pm 0.025\,M_\odot$, Gillon et al. 2012) and is also one of the coldest of all stars hosting hot Jupiters. The close proximity of the planet probably induces large tidal bulges on the planet's surface (Ragozzine & Wolf 2009). The planet's projected lifetime is also unusually short for such a late-type host star, owing to tidal in-spiral. The estimated lifetime for this planet is perhaps 10 Myr–1 Gyr (Hellier et al. 2011).

In this paper we report two *Spitzer* secondary eclipse observations, using the IRAC 3.6 and 4.5 $\mu$m channels. The S/N of 26 in channel 1 and 24 in channel 2 allowed a nonambiguous analysis. The final eclipse depths from our joint-fit models are 0.347% ± 0.013% and 0.382% ± 0.015%, in channels 1 and

**Table 8**
System Parameters of WASP-43

| Parameter | Value | Reference |
|---|---|---|
| Eclipse Parameters | | |
| Eclipse midpoint ($BJD_{TDB}$) (2011 Jul 30) | $2455773.3179 \pm 0.0003$ | (1) |
| Eclipse midpoint ($BJD_{TDB}$) (2011 Jul 29) | $2455772.5045 \pm 0.0003$ | (1) |
| Eclipse duration $t_{4-1}$ (hr) | $1.25 \pm 0.02$ | (1) |
| Eclipse depth *Spitzer* IRAC, 3.6 $\mu$m (%) | $0.347 \pm 0.013$ | (1) |
| Eclipse depth *Spitzer* IRAC, 4.5 $\mu$m (%) | $0.382 \pm 0.015$ | (1) |
| Eclipse depth VLT HAWK-I, 2.095 $\mu$m (ppm) | $1560 \pm 140$ | (2) |
| Eclipse depth VLT HAWK-I, 1.186 $\mu$m (ppm) | $790 \pm 320$ | (2) |
| Ingress/egress time $t_{2-1}$ (hr) | $0.264 \pm 0.018$ | (1) |
| Orbital parameters | | |
| Orbital period, $P$ (days) | $P = 0.81347436 \pm 1.4 \times 10^{-7}$ | (3) |
| Semimajor axis, $a$ (AU) | $0.01526 \pm 0.00018$ | (2) |
| Transit time ($BJD_{TDB}$) | $2455726.54336 \pm 0.00012$ | (2) |
| Orbital eccentricity, $e$ | $0.010^{+0.010}_{-0.007}$ | (3) |
| Argument of pericenter, $\omega$ (deg) | $-88^{+5}_{-9}$ | (3) |
| Velocity semiamplitude, $K$ (m s$^{-1}$) | $549 \pm 6$ | (3) |
| Center-of-mass velocity $\gamma$ (m s$^{-1}$) | $-3595 \pm 4$ | (3) |
| Stellar parameters | | |
| Spectral type | see Section 1 | (2) |
| Mass, $M_*$ ($M_\odot$) | $0.717 \pm 0.025$ | (2) |
| Radius, $R_*$ ($R_\odot$) | $0.667^{+0.011}_{-0.010}$ | (2) |
| Mean density, $\rho_*$ ($\rho_\odot$) | $2.410^{+0.079}_{-0.075}$ | (2) |
| Effective temperature, $T_{eff}$ (K) | $4520 \pm 120$ | (2) |
| Surface gravity, log $g_*$ (cgs) | $4.645^{+0.011}_{-0.010}$ | (2) |
| Projected rotation rate, $v_* \sin(i)$ (kms$^{-1}$) | $4.0 \pm 0.4$ | (2) |
| Metallicity [Fe/H] (dex) | $-0.01 \pm 0.12$ | (2) |
| Distance (pc) | $80 \pm 20$ | (4) |
| Planetary parameters | | |
| Mass, $M_p$ ($M_J$) | $2.034^{+0.052}_{-0.051}$ | (2) |
| Radius, $R_p$ ($R_J$) | $1.036 \pm 0.019$ | (2) |
| Surface gravity, log $g_p$ (cgs) | $3.672^{+0.013}_{-0.012}$ | (2) |
| Mean density, $\rho_p$ (g cm$^{-3}$) | $1.377^{+0.063}_{-0.059}$ | (2) |
| Equilibrium temperature (A = 0), $T_{eq}$ (K) | $1440^{+40}_{-39}$ | (2) |

**Notes. References.** (1) This work (parameters, derived using joint fit; see Section 4.3); (2) Gillon et al. (2012); (3) This work (see Section 5); (4) Hellier et al. (2011).

2, respectively. The corresponding brightness temperatures are $1670 \pm 23$ K and $1514 \pm 25$ K.

Our secondary eclipse timings, along with the available RV data and transit photometry from the literature and amateur observations, provide better constraints on the orbital parameters. WASP-43b's orbital period is improved by a factor of three ($P = 0.81347436 \pm 1.4 \times 10^{-7}$ days). The timing of our secondary eclipse observations is consistent with and suggestive of a circular orbit.

We combined our *Spitzer* eclipse depths with ground-based data in the near-infrared from Gillon et al. (2012) and Wang et al. (2013) to constrain the atmospheric properties of WASP-43b. The data rule out a strong thermal inversion in the dayside atmosphere. This is particularly evident because the brightness temperatures in both the *Spitzer* channels are lower than those observed in the ground-based channels, suggesting temperatures decreasing outward. The data do not suggest very efficient day–night energy redistribution in the planet, consistent with previous studies, though models with up to ~35% redistribution can explain the data reasonably well. Current data are insufficient to provide stringent constraints on the chemical composition.

WASP-43b is a promising planet for a variety of future observations. Its high eclipse S/N makes it a prime candidate for dayside mapping using eclipse ingress and egress data (e.g., de Wit et al. 2012; Majeau et al. 2012). Observations in the HST WFC3 bandpass could break the degeneracy between O-rich and C-rich atmospheric models. Finally, the possibility of measuring orbital decay in the future is exciting because of the unique constraints this could place on stellar interior parameters.

## APPENDIX

## SYSTEM PARAMETERS

Table 8 lists WASP-43 system parameters derived from our analysis and the literature.

## REFERENCES

Adams, E. R., López-Morales, M., Elliot, J. L., Seager, S., & Osip, D. J. 2010, ApJ, 721, 1829
Anderson, D. R., Smith, A. M. S., Lanotte, A. A., et al. 2011, MNRAS, 416, 2108
Arras, P., Burkart, J., Quataert, E., & Weinberg, N. N. 2012, MNRAS, 422, 1761
Batalha, N. M., & Kepler Team. 2012, BAAS, 220, 306.01
Beerer, I. M., Knutson, H. A., Burrows, A., et al. 2011, ApJ, 727, 23
Berger, D. H., Gies, D. R., McAlister, H. A., et al. 2006, ApJ, 644, 475
Berta, Z. K., Charbonneau, D., Désert, J.-M., et al. 2012, ApJ, 747, 35
Blecic, J., Harrington, J., Madhusudhan, N., et al. 2013, ApJ, 779, 5
Borucki, W. J., Koch, D. G., Basri, G., et al. 2011, ApJ, 736, 19
Borysow, A. 2002, A&A, 390, 779
Borysow, A., Jorgensen, U. G., & Zheng, C. 1997, A&A, 324, 185
Boyajian, T. S., von Braun, K., van Belle, G., et al. 2012, ApJ, 757, 112
Campo, C. J., Harrington, J., Hardy, R. A., et al. 2011, ApJ, 727, 125
Chabrier, G., & Baraffe, I. 1997, A&A, 327, 1039
Charbonneau, D., Allen, L. E., Megeath, S. T., et al. 2005, ApJ, 626, 523
Cowan, N. B., Machalek, P., Croll, B., et al. 2012a, ApJ, 747, 82
Cowan, N. B., Voigt, A., & Abbot, D. S. 2012b, ApJ, 757, 80
Crossfield, I. J. M., Barman, T., Hansen, B. M. S., Tanaka, I., & Kodama, T. 2012, ApJ, 760, 140
Crossfield, I. J. M., Hansen, B. M. S., Harrington, J., et al. 2010, ApJ, 723, 1436
Cubillos, P., Harrington, J., Madhusudhan, N., et al. 2013, ApJ, 768, 42
de Wit, J., Gillon, M., Demory, B.-O., & Seager, S. 2012, A&A, 548, A128
Deming, D., Knutson, H., Agol, E., et al. 2011, ApJ, 726, 95
Deming, D., Wilkins, A., McCullough, P., et al. 2013, ApJ, 774, 95
Demory, B.-O., Gillon, M., Barman, T., et al. 2007, A&A, 475, 1125
Demory, B.-O., Gillon, M., Deming, D., et al. 2011, A&A, 533, A114
Desert, J.-M., Charbonneau, D., Fressin, F., & Torres, G. 2012, BAAS, 219, 414.02
Eastman, J., Siverd, R., & Gaudi, B. S. 2010, PASP, 122, 935
Fazio, G. G., Hora, J. L., Allen, L. E., et al. 2004, ApJS, 154, 10
Fortney, J. J., Lodders, K., Marley, M. S., & Freedman, R. S. 2008, ApJ, 678, 1419
Fortney, J. J., Marley, M. S., Lodders, K., Saumon, D., & Freedman, R. 2005, ApJL, 627, L69

Fraine, J. D., Deming, D., Gillon, M., et al. 2013, ApJ, 765, 127
Freedman, R. S., Marley, M. S., & Lodders, K. 2008, ApJS, 174, 504
Gelman, A. 2002, in Encyclopedia of Environmetrics, Vol. 3, ed. A. H. E.-S. W. W. Piegorsch (Chichester, NY: Wiley), 1634
Gelman, A., & Rubin, D. 1992, StaSc, 7, 457
Gibson, N. P., Aigrain, S., Pont, F., et al. 2012, MNRAS, 422, 753
Gillon, M., Triaud, A. H. M. J., Fortney, J. J., et al. 2012, A&A, 542, A4
Harrington, J., Luszcz, S., Seager, S., Deming, D., & Richardson, L. J. 2007, Natur, 447, 691
Harris, G. J., Larner, F. C., Tennyson, J., et al. 2008, MNRAS, 390, 143
Hellier, C., Anderson, D. R., Collier Cameron, A., et al. 2011, A&A, 535, L7
Hubeny, I., Burrows, A., & Sudarsky, D. 2003, ApJ, 594, 1011
Jackson, B., Greenberg, R., & Barnes, R. 2008, ApJ, 678, 1396
Karkoschka, E., & Tomasko, M. G. 2010, Icar, 205, 674
Knutson, H. A., Charbonneau, D., Cowan, N. B., et al. 2009, ApJ, 703, 769
Knutson, H. A., Howard, A. W., & Isaacson, H. 2010, ApJ, 720, 1569
Knutson, H. A., Lewis, N., Fortney, J. J., et al. 2012, ApJ, 754, 22
Knutson, H. A., Madhusudhan, N., Cowan, N. B., et al. 2011, ApJ, 735, 27
Kopparapu, R. k., Kasting, J. F., & Zahnle, K. J. 2012, ApJ, 745, 77
Levrard, B., Winisdoerffer, C., & Chabrier, G. 2009, ApJL, 692, L9
Lewis, N. K., Knutson, H. A., Showman, A. P., et al. 2013, ApJ, 766, 95
Lodders, K. 2004, ApJ, 611, 587
López-Morales, M., Coughlin, J. L., Sing, D. K., et al. 2010, ApJL, 716, L36
Lust, N. B., Britt, D. T., Harrington, J., et al. 2013, PASP, submitted
Madhusudhan, N. 2012, ApJ, 758, 36
Madhusudhan, N., Harrington, J., Stevenson, K. B., et al. 2011a, Natur, 469, 64
Madhusudhan, N., Mousis, O., Johnson, T. V., & Lunine, J. I. 2011b, ApJ, 743, 191
Madhusudhan, N., & Seager, S. 2009, ApJ, 707, 24
Madhusudhan, N., & Seager, S. 2010, ApJ, 725, 261
Majeau, C., Agol, E., & Cowan, N. B. 2012, ApJL, 747, L20
Mandel, K., & Agol, E. 2002, ApJL, 580, L171
Mousis, O., Lunine, J. I., Madhusudhan, N., & Johnson, T. V. 2012, in European Planetary Science Congress 2012, EPSC2012-366
Nymeyer, S., Harrington, J., Campo, C., et al. 2011, ApJ, 742, 35
Öberg, K. I., Murray-Clay, R., & Bergin, E. A. 2011, ApJL, 743, L16
Poddaný, S., Brát, L., & Pejcha, O. 2010, NewA, 15, 297
Pont, F., Zucker, S., & Queloz, D. 2006, MNRAS, 373, 231
Ragozzine, D., & Wolf, A. S. 2009, ApJ, 698, 1778
Reach, W. T., Megeath, S. T., Cohen, M., et al. 2005, PASP, 117, 978
Rothman, L. S., Jacquemart, D., Barbe, A., et al. 2005, JQSRT, 96, 139
Showman, A. P., Fortney, J. J., Lian, Y., et al. 2009, ApJ, 699, 564
Showman, A. P., & Guillot, T. 2002, A&A, 385, 166
Siess, L., Forestini, M., & Dougados, C. 1997, A&A, 324, 556
Spiegel, D. S., Silverio, K., & Burrows, A. 2009, ApJ, 699, 1487
Stevenson, K. B., Harrington, J., Fortney, J. J., et al. 2012, ApJ, 754, 136
Stevenson, K. B., Harrington, J., Nymeyer, S., et al. 2010, Natur, 464, 1161
Todorov, K. O., Deming, D., Knutson, H. A., et al. 2012, ApJ, 746, 111
Wang, W., van Boekel, R., Madhusudhan, N., et al. 2013, ApJ, 770, 70
Winn, J. N., Holman, M. J., Torres, G., et al. 2008, ApJ, 683, 1076
Zahnle, K., Marley, M. S., Freedman, R. S., Lodders, K., & Fortney, J. J. 2009, ApJL, 701, L20