

GALAXY AND MASS ASSEMBLY (GAMA): MID-INFRARED PROPERTIES AND EMPIRICAL RELATIONS FROM *WISE*

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ABSTRACT

The Galaxy And Mass Assembly (GAMA) survey furnishes a deep redshift catalog that, when combined with the *Wide-field Infrared Survey Explorer* (*WISE*), allows us to explore for the first time the mid-infrared properties of >110,000 galaxies over 120 deg^2 to $z \simeq 0.5$. In this paper we detail the procedure for producing the matched GAMA-*WISE* catalog for the G12 and G15 fields, in particular characterizing and measuring resolved sources; the complete catalogs for all three GAMA equatorial fields will be made available through the GAMA public releases. The wealth of multiwavelength photometry and optical spectroscopy allows us to explore empirical relations between optically determined stellar mass (derived from synthetic stellar population models) and $3.4 \mu\text{m}$ and $4.6 \mu\text{m}$ *WISE* measurements. Similarly dust-corrected $H\alpha$ -derived star formation rates can be compared to $12 \mu\text{m}$ and $22 \mu\text{m}$ luminosities to quantify correlations that can be applied to large samples to $z < 0.5$. To illustrate the applications of these relations, we use the $12 \mu\text{m}$ star formation prescription to investigate the behavior of specific star formation within the GAMA-*WISE* sample and underscore the ability of *WISE* to detect star-forming systems at $z \sim 0.5$. Within galaxy groups (determined by a sophisticated friends-of-friends scheme), results suggest that galaxies with a neighbor within $100 h^{-1} \text{ kpc}$ have, on average, lower specific star formation rates than typical GAMA galaxies with the same stellar mass.

Key words: catalogs – galaxies: star formation – galaxies: stellar content – infrared: galaxies – surveys

Online-only material: color figures

1. INTRODUCTION

Spanning almost six decades and counting, large area surveys have revolutionized our view of the structure and evolution of the universe, for example, the Palomar Observatory Sky Survey (POSS II; Reid et al. 1991), the IIIa-J SRC Southern Sky Survey, the *Infrared Astronomy Satellite* (*IRAS*) Sky Survey Atlas (Wheelock et al. 1994), 2dF Galaxy Redshift Survey (2dFGRS; Colless et al. 2001), the Sloan Digital Sky Survey (SDSS; Abazajian et al. 2003), the H I Parkes All Sky Survey (HIPASS; Meyer et al. 2004), the 6dF Galaxy Survey (6dFGS; Jones et al. 2004), the 2 Micron All Sky Survey (2MASS; Skrutskie et al. 2006), and the UKIRT Infrared Deep Sky Survey (UKIDSS; Lawrence et al. 2007). Uncovering the properties of large galaxy populations has proved essential

to our understanding of how galaxies develop and transform. Increasing sensitivity and scale in the time domain, as with Skymapper and the planned Large Synoptic Survey Telescope holds the promise of enormous future scientific returns. In the near future, leading-edge radio telescopes and planned strategic surveys will add a key ingredient to the mix—H I and continuum measurements will trace neutral gas reservoirs and activity to unprecedented depth, sky coverage and resolution with WALLABY and EMU on Australian Square Kilometer Array Pathfinder (ASKAP), Apertif-WNSHS, LADUMA (Holwerda & Blyth 2010) on MeerKAT, and the JVLA.

In order to better understand the key physics at work in star-forming galaxies, and crucially the efficacy of different physical mechanisms, one requires the combination of extended area and sufficient depth in order to capture a cosmologically and evolutionarily significant volume. The Galaxy and Mass Assembly (GAMA; Driver et al. 2009, 2011) survey provides exactly this

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laboratory. At its heart, GAMA is an optical spectroscopic survey of up to $\sim 300,000$ galaxies (Hopkins et al. 2013) obtained at the Anglo-Australian Telescope situated at Siding Spring Observatory. Three equatorial fields (G09, G12, and G15) covering 180 deg^2 sample large-scale structure to a redshift of $z \simeq 0.5$, with overall median redshift of $z \simeq 0.3$, and two southern fields (G02 and G23). Multiwavelength ancillary data from ultraviolet to far-infrared wavelengths provides comprehensive spectral energy distribution (SED) wavelength real estate. At mid-infrared wavelengths it falls to *WISE*, the *Wide-field Infrared Survey Explorer*, to span the near-infrared transition from stellar to dust emission.

Launched in 2009 December, *WISE* surveyed the entire sky in four mid-infrared bands: $3.4 \mu\text{m}$, $4.6 \mu\text{m}$, $12 \mu\text{m}$, and $22 \mu\text{m}$ (Wright et al. 2010). In the local Universe, the $3.4 \mu\text{m}$ (W1) and $4.6 \mu\text{m}$ (W2) bands chiefly trace the continuum emission from evolved stars (see, for example, Meidt et al. 2012). The W1 band is the most sensitive to stellar light, typically reaching L^* depths to $z \simeq 0.5$. The W2 band is additionally sensitive to hot dust; hence, this makes the $3.4 \mu\text{m} - 4.6 \mu\text{m}$ color very sensitive to galaxies dominated globally by active galactic nucleus (AGN) emission (see, e.g., Jarrett et al. 2011; Stern et al. 2012). The $12 \mu\text{m}$ (W3) band is broad (see Wright et al. 2010, for relative system response curves) and can trace both the $9.7 \mu\text{m}$ silicate absorption feature, as well as the $11.3 \mu\text{m}$ polycyclic aromatic hydrocarbon (PAH) and [Ne II] emission lines. Lastly, the W4 band traces the warm dust continuum at $22 \mu\text{m}$, sensitive to reprocessed radiation from star formation and AGN activity. *WISE* is thus optimally suited to study the diverse emission mechanisms of galaxies.

WISE is confusion-noise-limited—the structure of the background (e.g., distant galaxies and scattered light) increases noise by contributing flux—and under these conditions point spread function (PSF) profile-fitted photometry performs robustly (Marsh & Jarrett 2012). For unresolved sources, this is the adopted method of obtaining photometry and accordingly the *WISE* Source Catalog is optimized and calibrated for point sources. The angular resolution of *WISE* frames, used for determining the profile photometry measurement (keyword *w*mpro* in the *WISE* All-Sky Data Release), is $6''.1$, $6''.4$, $6''.5$, and $12''.0$ for W1, W2, W3, and W4, respectively. The *WISE* public-release data products also includes an image Atlas of tiled mosaics; however, the angular resolution is slightly degraded in the Atlas images because they are convolved with a matched filter to optimize point-source detection. The crucial point here is that the *WISE* catalogs and images are not well-suited to characterization of resolved sources. Resolved sources will either be detected and measured as conglomerations of several point sources, or have their flux grossly underestimated by the PSF profile. As a consequence, it is left to the community to properly measure sources that are resolved by *WISE* (i.e., nearby galaxies). In this work we have produced new, better-optimized images to extract resolved galaxy measurements from the GAMA fields to complement the point-source measurements obtained from the public-release catalogs.

A detailed *WISE* study of several nearby ($< 60 \text{ Mpc}$) galaxies (Jarrett et al. 2013) has highlighted how *WISE* data can be used in multiwavelength studies of star formation and interstellar medium (ISM) conditions, as well as tracing global properties such as stellar mass and star formation. The work of Donoso et al. (2012) showed the power of combining *WISE* with a large-area survey (SDSS); they investigated the effect of star formation- and AGN-activity on the *WISE* properties of

$> 95,000$ galaxies, including calibrations using the $12 \mu\text{m}$ and $22 \mu\text{m}$ luminosities. These empirical calibrations have also been investigated in $22 \mu\text{m}$ selected galaxies in SDSS galaxies with $z < 0.3$ by Lee et al. (2013). *WISE* colors have proved to be an excellent AGN selection tool (Jarrett et al. 2011; Stern et al. 2012; Assef et al. 2013), and as a diagnostic for determining the accretion modes amongst radio-loud AGNs (Gurkan et al. 2013). Yan et al. (2013) combine *WISE* and SDSS to provide a phenomenological characterization for *WISE* extragalactic sources.

In this paper we harness the power of the GAMA survey, and its value-added data products, cross matching GAMA galaxies in two completed GAMA fields (the G12 and G15 regions) to their *WISE* counterparts. The G09 region is not considered in this analysis, but will be included as part of the GAMA-*WISE* data release.

GAMA counterparts at low redshift will be predominantly resolved, particularly in the W1 and W2 bands. Unlike previous studies, we include robust measurements of resolved sources in these fields and distinguish between resolved and unresolved systems to determine the photometry most appropriate for a given source. Since GAMA is highly spectroscopically complete to $r_{\text{AB}} < 19.8$ within the two GAMA regions, we are able to push *WISE* to higher redshifts than previously possible for wide-field surveys. We explore empirical relationships for stellar mass and star formation rates (SFRs), the color distribution of *WISE* sources in GAMA and their behavior in dense environments. The paper is arranged as follows: in Section 2 we describe the data used in this analysis, Section 3 discusses how the GAMA-*WISE* matched catalog is constructed, Section 4 contains results derived from the combined surveys and Section 5 illustrates scientific applications of the catalog. Conclusions are presented in Section 6.

The cosmology adopted throughout this paper is $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $h = H_0/100$, $\Omega_M = 0.3$, and $\Omega_\Lambda = 0.7$. All magnitudes are in the Vega system (*WISE* calibration described in Jarrett et al. 2011) unless indicated explicitly by an AB subscript.

2. DATA

2.1. GAMA

Our sample is drawn from the G12 and G15 equatorial fields of the GAMA II survey (Driver et al. 2009, 2011) combining high spectroscopic completeness ($\simeq 97\%$) to a limiting magnitude of $r_{\text{AB}} = 19.8$, with a wealth of ancillary photometric data. Details of target selection for the survey are outlined in Baldry et al. (2010) and optimal tiling for the survey in Robotham et al. (2010). Spectra were obtained primarily with the 2dF instrument mounted on the 3.9 m Anglo-Australian Telescope, and additionally from the Sloan Digital Sky Survey (SDSS DR7; Abazajian et al. 2009). Reduction and analysis of the spectra is discussed in detail in Hopkins et al. (2013) and J. Liske et al. (in preparation).

Photometric data for galaxies within the G12 and G15 volumes (ApMatchedCatv05) is drawn from SDSS imaging (u, g, r, i, z) as outlined in Hill et al. (2011) and VISTA VIKING (ZYJHK) as detailed in S. P. Driver et al. (in preparation). Photometry is corrected for Galactic foreground dust extinction using Schlegel et al. (1998). Stellar mass estimates (StellarMassesv15) are derived from matched-aperture photometry using synthetic stellar population models from Bruzual &

Table 1
WISE Cross-match Statistics for Each Band

G12 (Total: 60645 Sources)				
Band	Detections	S/N > 2 ^a	S/N > 5	Upper Limits ^b
W1	100%	100%	99.9%	0%
W2	100%	96%	77%	4%
W3	99%	47%	17%	52%
W4	98%	19%	2%	79%
G15 (Total: 58199 Sources)				
Band	Detections	S/N > 2	S/N > 5	Upper Limits ^b
W1	100%	100%	99.9%	0.01%
W2	100%	98%	88%	2%
W3	99%	58%	29%	42%
W4	98%	19%	3%	78%

Notes.

^a Signal-to-noise ratio as measured by *w*mp_{ro}*.

^b Detections with 2σ upper limits.

Charlot (2003) as detailed in Taylor et al. (2011). We take advantage of the completed observations in the equatorial fields and updated redshifts (I. Baldry et al., in preparation) and stellar masses. Throughout our analysis we only use sources with reliable redshifts, i.e., $nQ \geq 3$ (Driver et al. 2011).

SFRs can be derived from the H α equivalent width, combined with the *r*-band absolute magnitude, and are available for the GAMA phase-I survey ($r_{AB} < 19.4$ in G15 and $r_{AB} < 19.8$ in G12) as determined in Gunawardhana et al. (2011, 2013). Corrections are applied for the underlying Balmer stellar absorption, dust obscuration, and fiber aperture effects. To maintain consistency we only use SFRs from Gunawardhana et al. (2013) derived for galaxies with redshifts matching in both GAMA I and GAMA II ($\simeq 10\%$ were excluded because of this criterion).

2.2. WISE Image Construction

The WISE mission provides “Atlas” Images via its data repository, which are co-added and interpolated from the multiple single exposure frames taken during the survey (Cutri et al. 2012). These $1^{\circ}56 \times 1^{\circ}56$ mosaics have a $1/375$ pixel scale, but the resampling and co-addition method is optimized for point source detection; Atlas images have beam sizes of $8''.1$, $8''.8$, $11''.0$, and $17''.5$ for the W1, W2, W3, and W4 bands, respectively. To better preserve the native resolution of the single frames we employ a mosaic construction that uses the “drizzle” resampling technique. Variable-Pixel Linear Reconstruction or drizzling co-addition algorithm using a tophat point response function can be used to improve the spatial resolution compared to the nominal Atlas Images (Jarrett et al. 2012). Drizzled image mosaics, 1.3×1.3 degrees in size, were created using the software package ICORE (Masci 2013) achieving a resolution of $5''.9$, $6''.5$, $7''.0$, and $12''.4$ in the $3.4 \mu\text{m}$, $4.6 \mu\text{m}$, $12 \mu\text{m}$, and $22 \mu\text{m}$, respectively. Within ICORE, background level offset-matching, flagging and outlier rejection, and co-addition using overlap area-weighted interpolation ensures optimal background stability. The number of images that are combined is dependent on the depth of coverage and additional orbits, a function of the field’s location relative to the ecliptic. For the GAMA fields the coverage is typically 12 and 24 frames for the G12 and G15 fields, respectively, where G15 benefits from additional scans in that region arising from multiple epoch WISE orbits.

Table 2
Parameters from the WISE All-Sky Catalog

<i>designation</i>	Unique WISE source designation
<i>ra</i> (deg)	Right ascension (J2000)
<i>dec</i> (deg)	Declination (J2000)
<i>sig_{ra}</i> (arcsec)	Uncertainty in R.A.
<i>sig_{dec}</i> (arcsec)	Uncertainty in Decl.
<i>w*mp_{ro}</i> (mag)	Instrumental profile-fit photometry magnitude
<i>w*sig_{mp_{ro}}</i> (mag)	Instrumental profile-fit photometry uncertainty
<i>w*r_{chi}2</i>	Instrumental profile-fit photometry reduced χ^2
<i>nb</i>	Number of blend components used in each fit
<i>na</i>	Active deblend flag (=1 if actively deblended)
<i>xsc_{prox}</i> (arcsec)	Distance between source center and XSC ^a galaxy
<i>w*r_{semi}</i>	(Scaled) semi-major axis of galaxy from XSC ^a
<i>w*g_{mag}</i>	Elliptical aperture mag of extracted galaxy

Notes.

* = 1, 2, 3, 4.

^a 2MASS Extended Source Catalog (Jarrett et al. 2000).

2.3. Cross Match to GAMA G12 and G15

Galaxies were extracted from the G12 and G15 GAMA II catalogs of observed sources (i.e., TilingCatv41 with redshifts as in SpecAllv21 and $z > 0.001$)—69,693 and 64,822 for G12 and G15, respectively, and cross-matched to the WISE All-Sky Catalog using the NASA/IPAC Infrared Science Archive and a $3''$ cone search radius. The number of unique objects in each field were 60,645 GAMA matches for G12 and 58,199 for G15, i.e., a 87.0% and 89.8% match rate for G12 and G15, respectively, with a higher match rate in G15 due to the increased WISE depth (from additional scans) and therefore sensitivity. Further details are provided in Table 1 and a region of G15 shown in Figure 1 to illustrate the difference between the optical and mid-infrared sky.

For this analysis, we rely on several parameters output by the WISE All-Sky Catalog; these are listed in Table 2 with a brief explanation.

3. SOURCES RESOLVED BY WISE

A key feature of a catalog containing WISE photometry is determining which of the galaxies are resolved and then ensuring their fluxes are measured robustly. A broad indication is given by the reduced χ^2 (*w*r_{chi}2*) of the profile-fit solution, where high values of the *w*r_{chi}2* parameter indicate that the wpro algorithm measurement is underestimating the flux of the source. Sources with *w*r_{chi}2* > 2 are often resolved, with *w*r_{chi}2* > 5 usually indicating well-resolved systems.

Unfortunately *w*r_{chi}* is not by itself a robust measure of “resolvedness,” particularly where sources have a low signal-to-noise ratio (S/N). The “reduced” metric tends to unity as noise begins to dominate the measurement. Moreover, the *w*mp_{ro}* fitting process can be fooled for sources with relatively bright cores that are just being resolved by WISE—this happens with 2MASS compact extended sources which have low *w*r_{chi}2*, but the *w*mp_{ro}* and isophotal photometry can be systematically offset.

Even though 2MASS has superior ($2\times$) spatial resolution compared to WISE, with its greater sensitivity WISE is able to resolve many nearby, relatively small galaxies. Therefore, all 2MASS Extended Source Catalog (XSC; Jarrett et al. 2000) sources should be tested to determine if they are resolved by WISE (see Section 3.2). For galaxies not in the 2MASS XSC, values of *w*r_{chi}2* ≥ 2 should be used as an initial selection for

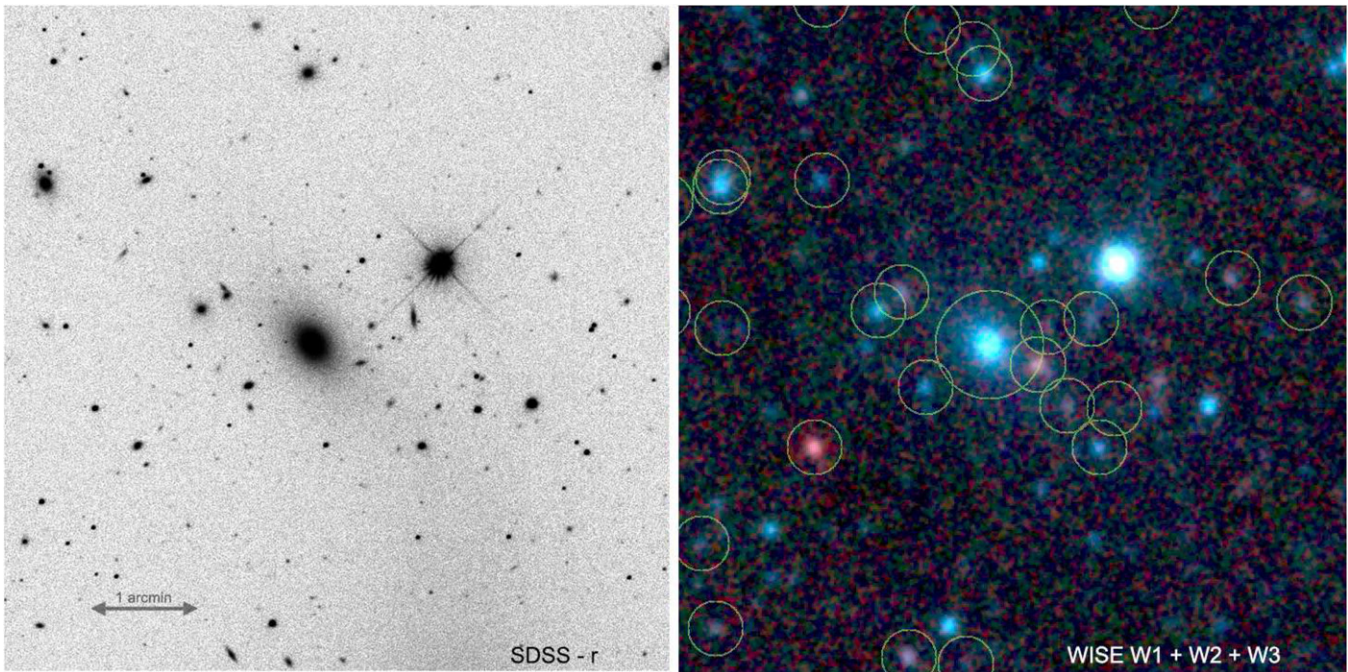


Figure 1. Representative $6' \times 6'$ cutout from the G15 field with the SDSS r -band image (left) and the three-color ($3.4 \mu\text{m}$, $4.6 \mu\text{m}$, $12 \mu\text{m}$) *WISE* image of the same region (right). *WISE*-GAMA galaxies are circled, demonstrating a mix of resolved and point-like sources.

(A color version of this figure is available in the online journal.)

resolved objects. Caution is required, however, since blended objects will also satisfy this criterion and can be a source of false positives, notably when the stellar confusion is significant.

3.1. Source Characterization of Potentially Resolved Sources

Sources are measured using custom software adapting tools and algorithms developed for the 2MASS XSC (Jarrett et al. 2000) and *WISE* photometry pipelines (Jarrett et al. 2011, 2013; Cutri et al. 2012). The process is semi-automated in that photometry measurements are automated, but each result is assessed by visual inspection with intervention where necessary.

The first step is to remove point sources by PSF subtraction which preserves the structure of the background. If necessary surrounding contaminating sources are masked (e.g., bright stars). The local background is estimated from pixel value (trimmed average) distribution that lies within an elliptical annulus located just outside of an “active region” which represents the image area that contains measurable light from the galaxy. The active region is not initially known, but it determined through successive iteration of the characterization process until convergence is reached.

The galaxy is modeled using an ellipsoid built from azimuthally averaging the (background-subtracted) surface brightness; any masked pixels are recovered using a weighted combination of the local background (to the masked source) and the galaxy model. The best-fit axis ratio and ellipticity are determined using the 3σ isophote and the galaxy shape is defined by this isophote and is assumed to be fixed at all radii. The primary “isophotal” photometry (W_{iso}) is then measured from the 1σ (of the background rms) elliptical isophote, capturing over 90% of the total light (see Jarrett et al. 2013). Other measurements include a double Sérsic fit, to the inner galaxy region (i.e., the bulge) and the outer region (i.e., the disk), thus allowing estimation of the total integrated flux that extends beyond the 1σ isophote. Since *WISE* is confusion-noise limited, we track the

photometry using a curve-of-growth table. Where a large mismatch occurs between the isophotal radius and the radius where the change in flux is less than 2% (the “convergence point”), the size of the active region is automatically decreased. This usually occurs where background levels are elevated due to a neighboring bright source which contaminates the background level. The process is iterated to adjust the active region until the source measurements have converged.

Once all the sources in a field are measured and modeled accordingly, the process is rerun to effectively perform galaxy–galaxy deblending. Where a measurement exists for a galaxy when measuring an adjacent source, the galaxy model is subtracted in addition to neighboring point sources. This mitigates contamination from nearby resolved sources.

3.2. The Resolved Parameter R_{fuzzy}

The *WISE* bands at $3.4 \mu\text{m}$ and $4.6 \mu\text{m}$ achieve the best spatial resolution, at $6''.1$ and $6''.4$, respectively, and capture extended light arising from evolved stars in the bulge and disk regions. Since *wIrchi2* cannot discriminate well between faint, resolved, and unresolved systems, we use the 2nd order intensity-weighted moment to describe the shape (major and minor axis) of the object and measure the intensity-weighted moment of the object on either side of the main axis.

If the object is symmetric and resolved on both sides of the major axis (as with a nominal resolved source), then each moment will have a resolved signature (i.e., the moment is large compared to a point source). If the object is, for example, a star blended with a galaxy, then their moments will be different (one half will have a moment that is point-like, the other resolved). Thus, taking the minimum moment value between the two halves, determines whether the source is resolved. Fuzzy objects will have a large moment regardless of which major axis half that is measured; stars or double stars will have a minimum

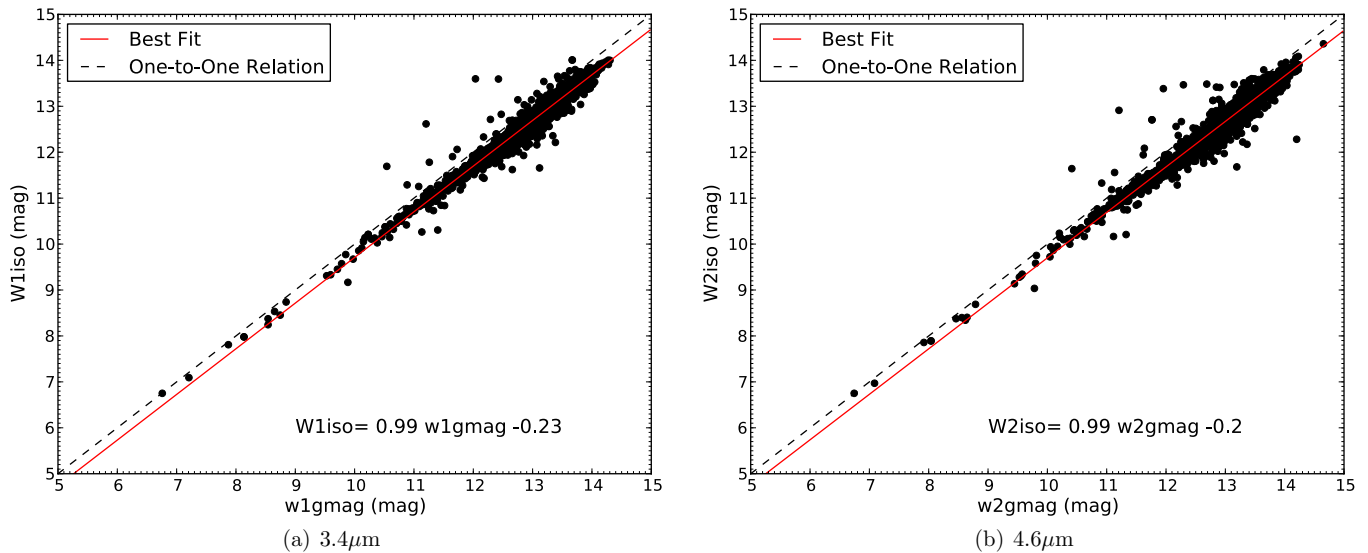


Figure 2. Comparisons of 1σ isophotal photometry (W^*iso) and the 2MASS XSC-derived scaled aperture (w^*gmag) photometry for resolved sources in the G12 and G15 fields. See Figure 20 for versions plotted as a function of w^*gmag .

(A color version of this figure is available in the online journal.)

moment that appears to be unresolved. The performance of the Rfuzzy parameter is illustrated in Section A.1 of the Appendix.

3.3. Catalog of Resolved Sources

Isolating the resolved sources using the Rfuzzy parameter (in W1) gives 1390 and 1368 sources in G12 and G15, respectively, i.e., 2%–3% of the total *WISE* matched sources in these fields. It is important to bear in mind that despite the resolution of the W1 and W2 bands, *WISE* is far more sensitive compared to, for example, 2MASS and will detect the extended light profiles from nearby galaxies. The typical W1 1σ isophotal radius is more than a factor of $\simeq 2$ in scale compared to the equivalent 2MASS K_s -band isophotal radius. An illustration of this is shown in Figure 2 where the 2MASS XSC-derived scaled aperture photometry (w^*gmag) are compared to the isophotal photometry of resolved sources in the GAMA G12 and G15 fields. The 2MASS XSC-derived magnitudes underestimate the inferred flux due to the increased sensitivity of *WISE*. Note that no star subtraction or deblending has been attempted when measuring the w^*gmags and contamination may be present; as a consequence, these measurements should only be used as a last resort (employing the offset shown in Figure 2 to obtain the correct galaxy flux). The behavior of resolved sources is further explored in Section A.2 of the Appendix.

3.4. GAMA-WISE Catalog Photometry

For point sources, the primary photometry are the profile fit measurements (w^*mpro), and for well-resolved sources the isophotal photometry (Section 3.1). When the source is not well resolved, or the S/N is low, the following steps are used to choose the best photometry:

1. w^*mpro photometry is used when the S/N is low, as measured by the isophotal aperture process for a given band: S/N of 10, 10, 15, and 15 for W1, W2, W3, and W4, respectively.
2. If the Rfuzzy parameter is false (measured in the W1 band), classify source as unresolved and use w^*mpro photometry for bands W1 and W2.

3. If $w3rchi2 < 2$, classify source as unresolved and revert to the $w3mpro$ photometry.
4. If $w4rchi2 < 2$, classify source as unresolved and revert to the $w4mpro$ photometry.
5. W1 is the most sensitive *WISE* band. To obtain an accurate W1–W2 color for galaxies resolved in W1 and W2, the following steps are employed. Firstly, an accurate W1–W2 color is determined using a matched aperture derived from the smaller of the two isophotal radii (which is usually W2). Then, from the isophotal magnitude of the W1 band, the corresponding W2 magnitude is determined, thereby reflecting the sensitivity of the W1 band. This is also done for the W4 band using W3 in the same way as W1, if sources are resolved at these wavelengths, which in GAMA is rare (only 16 sources in G12 and G15). We note that the W1 and W3 bands cannot be matched in this way since the flux at these wavelengths is produced by different physical processes (evolved stars versus PAH features).

Aperture corrections are applied to isophotal measurements as detailed in the Explanatory Supplement to the *WISE* All-Sky Data Release Products.²⁰

In Figure 3 we plot the S/N for the G12 and G15 fields; due to the additional $\simeq 2\times$ frame coverage the G15 field has higher S/N detections compared to G12 at a given W1 magnitude. The magnitude sensitivity limit in the W1 band for the G12 field is 16.6 (10σ), 17.3 (5σ) and for G15, 17.0 (10σ) and 17.7 (5σ). The combined fields have a 10σ magnitude sensitivity limit of 16.6 or $71 \mu\text{Jy}$. The redshift distribution of sources in the two fields (excluding upper limits) is shown in Figure 4. Since the G15 field is more sensitive, the peak of the distribution is shifted to a slightly higher redshift compared to G12. The two fields show very similar redshift distributions for $z < 1$ (Figure 4(b)), although the G12 matches have relatively more high- z sources.

3.5. Blending

Point sources are both passively and actively deblended by the *WISE* pipeline and the *WISE* All-Sky catalog provides the *na*

²⁰ http://WISE2.ipac.caltech.edu/docs/release/allsky/expsup/sec4_4c.html#apcor

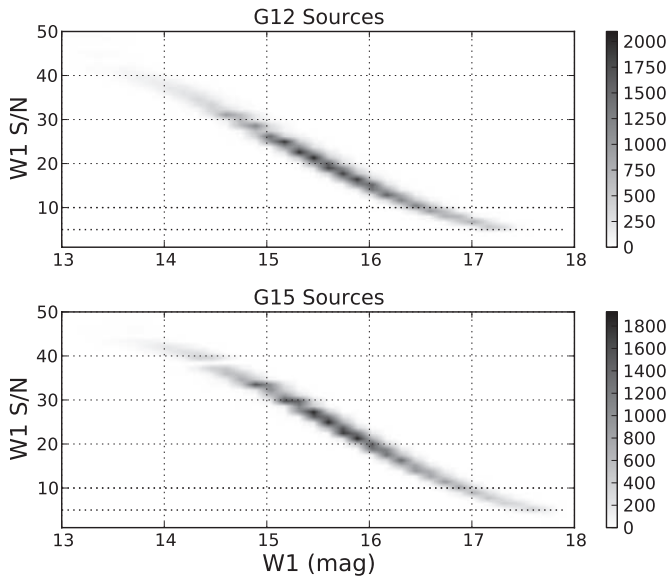


Figure 3. *WISE* W1 band ($3.4 \mu\text{m}$) signal-to-noise in the W1 band as a function of magnitude for G12 and G15 sources (resolved and unresolved sources). Note that G15 has better overall coverage and thus achieves greater sensitivity in luminosity.

and *nb* blending flags where a value of “0” indicates no active deblending has been performed, and “1” if sources have been deblended; deblending introduces an additional uncertainty on the measurement provided (*nb* is the number of components that were deblended).

However, an additional flag is needed to indicate where *WISE* cannot distinguish between GAMA sources (one *WISE* source for multiple GAMA sources) and where contamination is expected to be high since more than one GAMA source lies within the *WISE* beam. As such, we use an additional blending flag as determined by the proximity of neighboring GAMA-*WISE* sources. Sources within $5''$ are viewed as a catastrophic blend due to the size of the *WISE* beam and multiple GAMA sources will probably have the same *WISE* source matched to it and/or have a highly uncertain match. A source within $15''$

indicates a potential blend or contamination and is also flagged. For the analysis in this study we remove all flagged sources from our sample ($\approx 6\%$ of G12 and G15) to ensure the cleanest *WISE* photometry, but may bias the sample against the densest environments.

3.6. Rest-frame Colors

Rest-frame colors are determined using the GAMA redshift by building an SED combining the optical, near-infrared, and mid-infrared flux densities and fitting to an empirical template library, consisting of 126 galaxy templates, of local well-studied and morphologically diverse galaxies (e.g., the *Spitzer* Infrared Nearby Galaxy Survey; Kennicutt et al. 2003) from Brown et al. (2013). The templates are constructed from optical and *Spitzer* spectroscopy, with matched aperture photometry from *GALEX* (Martin et al. 2005), *Swift* (UVOT; Roming et al. 2005), SDSS, 2MASS, *Spitzer* (Werner et al. 2004), and *WISE*, combined with MAGPHYS (da Cunha et al. 2008). The *WISE* relative spectral response curves are from Jarrett et al. (2011) and are available as part of the *WISE* Explanatory Supplement.

The best template fit is determined by minimizing the function that describes the difference between the flux density in the measured band, and the model-template flux density. Each resulting best-fit is assigned a normalized score derived from the reduced χ^2 minimization, with the most consistent fits having values of < 2 and relatively poor fits with scores of > 3 . In the sections that follow we use only those sources with a score ≤ 2 , although the fitting accuracy is minimally important for nearby galaxies (i.e., with low rest-frame color corrections).

4. RESULTS

4.1. *WISE* Colors of the G12 and G15 Fields

As shown by Jarrett et al. (2011), the $3.4 \mu\text{m}$, $4.5 \mu\text{m}$, and $12 \mu\text{m}$ bands of *WISE* can be combined in a mid-infrared color-color diagram (Figure 5); the shortest bands are sensitive to the evolved stellar population and hot dust, therefore their color indicates increased activity (AGN or starburst). The $12 \mu\text{m}$ band is dominated by the $11.3 \mu\text{m}$ PAH, as well as the dust continuum, sensitive to star formation. This color-color

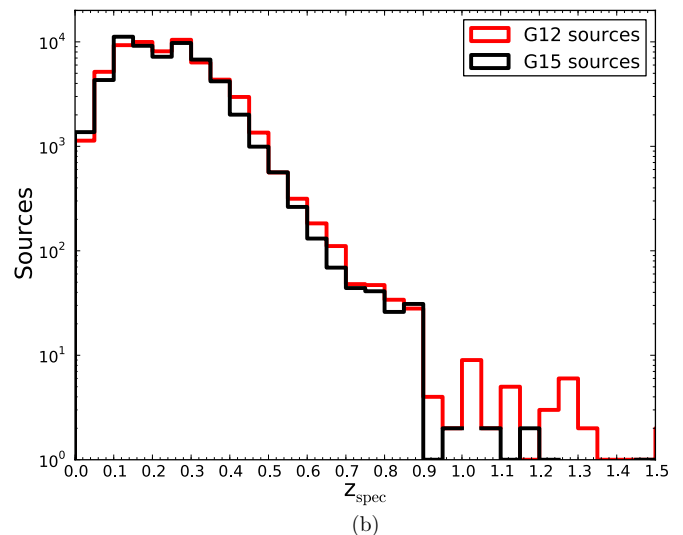
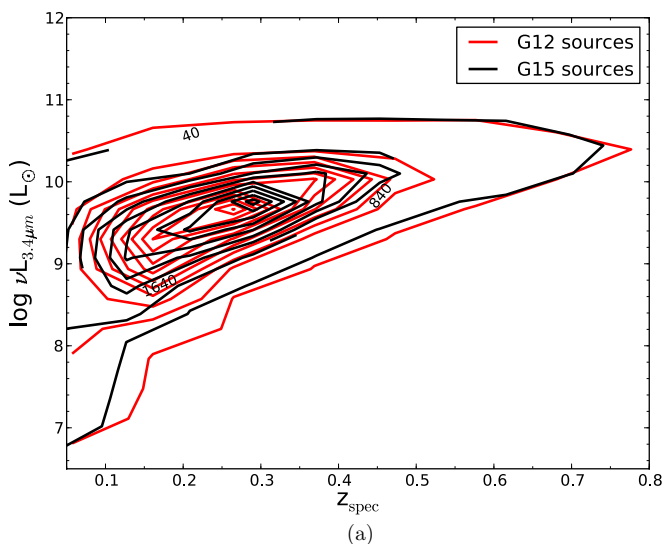


Figure 4. (a) The $3.4 \mu\text{m}$ luminosity density (νL_ν) distribution of sources in the two GAMA fields as a function of redshift. The number distribution as a function of redshift is shown in panel (b).

(A color version of this figure is available in the online journal.)

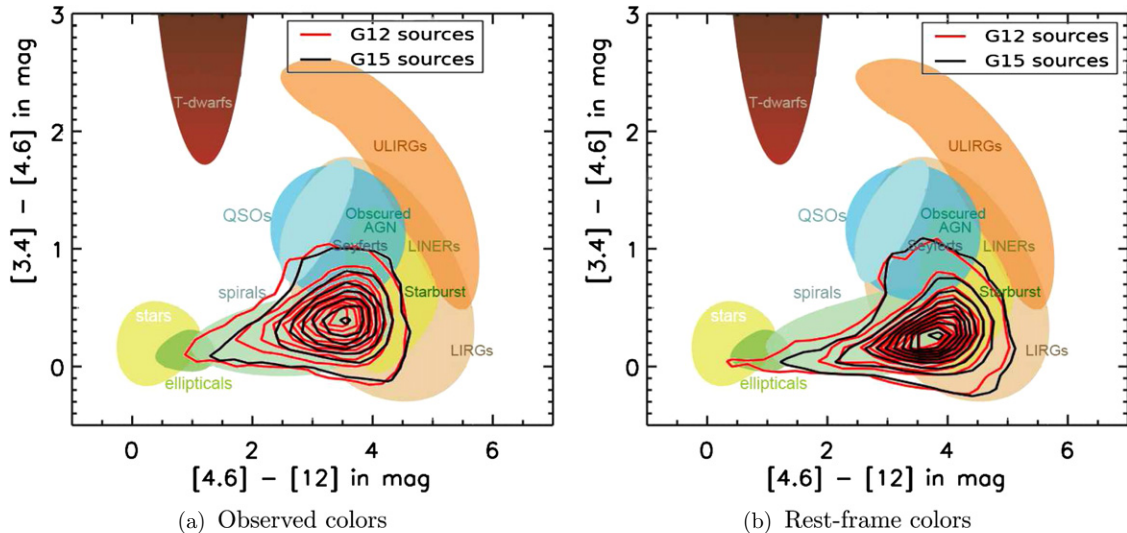


Figure 5. *WISE* colors of sources in G12 and G15 plotted on the color–color diagram of Jarrett et al. (2011) with observed colors shown in panel (a) and rest-frame (or “*k*-corrected”) colors in panel (b).

(A color version of this figure is available in the online journal.)

diagnostic is therefore useful to separate galaxy populations, particularly old stellar population-dominated, star-forming, and systems dominated by AGN-activity (Jarrett et al. 2011; Stern et al. 2012). In Figure 5(a) the observed colors of the G12 and G15 are shown, and the *k*-corrected version, Figure 5(b), shows the distribution of the rest-frame colors. Both fields have similar distributions, biased toward star-forming systems. This is not surprising given the optical selection of the *GAMA* sample. Heavily obscured galaxies, notably ultra-luminous infrared galaxies (ULIRGs) are absent due to the insensitivity (selective extinction) of SDSS and *GAMA* optical catalogs. Systems globally dominated by AGN-activity also appear sparse within the sample; this is consistent with the findings of Gunawardhana et al. (2013) where spectroscopically identified AGN make up <20% of the *GAMA* I emission-line catalog (see also Lara-López et al. 2013).

4.2. Aggregate Stellar Mass Estimation

The $3.4\ \mu\text{m}$ band of *WISE* is dominated by the light from old stars and can be used as an effective measure of stellar mass (Jarrett et al. 2013). To explore this further, we calculate the “in-band” luminosity for W1, i.e., the luminosity of the source as measured relative to the Sun in the W1 band, and use the stellar masses for *GAMA* as determined by Taylor et al. (2011) to determine a mass-to-light (*M/L*) ratio. These stellar masses are best constrained for $z < 0.15$ (Taylor et al. 2011) and we apply this cut to our sample. As shown by Jarrett et al. (2013), the stellar mass-to-light has a linear trend with *WISE* W1–W2 and W2–W3 color, reflecting systematic *M/L* differences between passive and star-forming systems.

In order to empirically calibrate a relation we require high S/N measurements and additionally use only the rest-frame W1–W2 color since the detection rate is much higher in W2 compared to W3. We apply a S/N cut of 13.5 in W1 and W2 and remove known AGNs (from the *GAMA* I spectroscopy measurements of Gunawardhana et al. 2013) and also systems with *WISE* colors consistent with AGN activity dominating their global colors ($W1-W2 \geq 0.8$) as discussed by Stern et al. (2012).

The first population we investigate are resolved, low- z galaxies, shown in Figure 6(a). We illustrate the danger of a

least-squares minimization fit compared to a bivariate-Gaussian, maximum-likelihood (or “best” fit) line (Taylor et al. 2014; Hogg et al. 2010). Throughout the paper we plot both the least-squares and maximum-likelihood lines to show the large discrepancies that may exist; we provide relations for the maximum-likelihood (or best fit) only. We find very good agreement with the $M_{\text{stellar}}/L_{W1}$ relation of Jarrett et al. (2013) based on their relatively small sample of 17 galaxies and with stellar masses effectively derived from $2\ \mu\text{m}$ K_s band photometry. Our relation for resolved low- z sources is

$$\log_{10} M_{\text{stellar}}/L_{W1} = -2.54(W_{3.4\ \mu\text{m}} - W_{4.6\ \mu\text{m}}) - 0.17, \quad (1)$$

with

$$L_{W1} (L_{\odot}) = 10^{-0.4(M - M_{\text{Sun}})},$$

where M is the absolute magnitude of the source in W1, $M_{\text{Sun}} = 3.24$, and $W_{3.4\ \mu\text{m}} - W_{4.6\ \mu\text{m}}$ reflects the rest-frame color of the source.

For comparison to the resolved sample, in Figure 6(b) we show all sources that meet our S/N selection criterion and redshift cut, which shows the distribution to be shifted to “warmer” W1–W2 colors, i.e., signifying systems with increased activity such as star-formation or low-power AGNs. In addition, the contours suggest that the distribution is made up of two populations, most probably due to passive galaxies having a larger mass-to-light ratio than disk-dominated and dwarf galaxies. We investigate this further in Figure 7 where we make a color-distinction based on the W3 measurement. By selecting galaxies with $W2-W3 \geq 1.5$ (see Figure 5) we choose systems that are most likely dominated by star formation. These selected types show a clear trend with mass-to-light (Figure 7(a)), but lie parallel and offset to the relation of Jarrett et al. (2013). But now including the passive galaxies, Figure 7(b), we see a clear clustering around a fixed mass-to-light ratio of $\simeq 0.7$.

We conclude that the best-fit for our entire sample is

$$\log_{10} M_{\text{stellar}}/L_{W1} = -1.96(W_{3.4\ \mu\text{m}} - W_{4.6\ \mu\text{m}}) - 0.03, \quad (2)$$

and for star-forming (lower mass-to-light) systems only,

$$\log_{10} M_{\text{stellar}}/L_{W1} = -1.93(W_{3.4\ \mu\text{m}} - W_{4.6\ \mu\text{m}}) - 0.04. \quad (3)$$

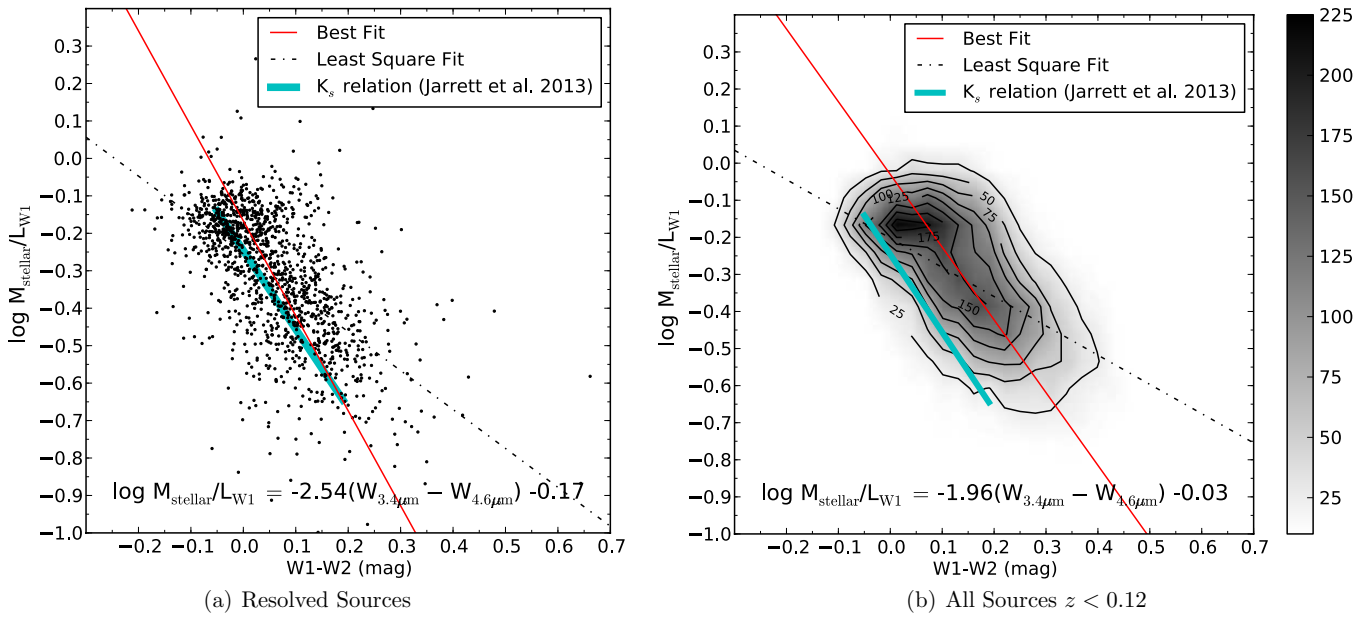


Figure 6. $3.4 \mu\text{m}$ mass-to-light ratio plotted as a function of $W1-W2$ (Vega) color for (a) resolved sources and (b) all sources in G12 and G15. The larger volume sample shows a shift to more active galaxies with “warmer” $W1-W2$ colors.

(A color version of this figure is available in the online journal.)

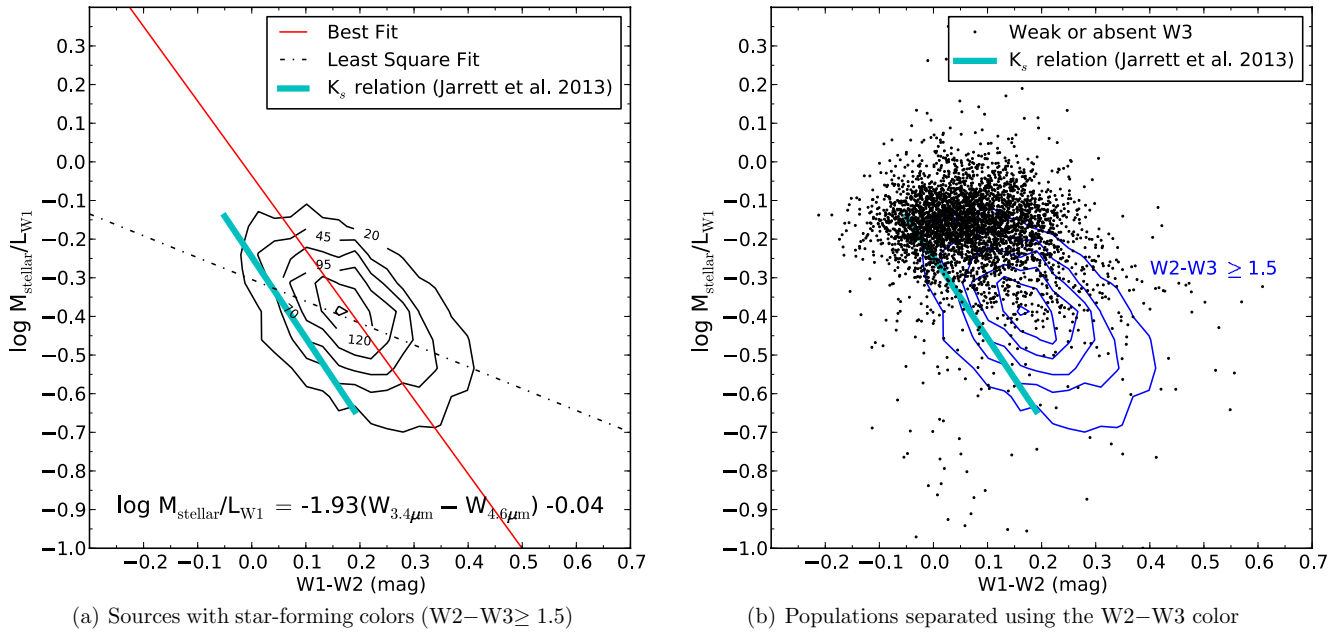


Figure 7. $W2-W3$ color allows us to separate star-forming systems from ones that are predominantly passive.

(A color version of this figure is available in the online journal.)

This likely explains the shift observed in the relation for resolved sources and the entire field sample. The resolved sources are relatively nearby and will be a mix of passive and star-forming systems. However, at higher redshifts we add relatively more star-forming galaxies (higher infrared luminosity) causing a larger spread in $W1-W2$ color. Notably galaxies with higher SFRs will have a larger $W1-W2$ color due to more hot dust emission giving a brighter $W2$ measurement. Also, at higher redshifts we will preferentially detect galaxies with higher SFRs (see next section). It should be noted that AGN activity would have a similar effect (as shown by the upward trend

in Figure 5(b)) and some contamination from nuclear activity is inevitable.

In Figure 8 we plot the residuals of the GAMA stellar masses (used to calibrate the *WISE* relation in Equation (2)) and the *WISE*-derived values themselves. The mass estimates agree within a factor of 1.2 at $10^{10} M_{\odot}$ (with a standard deviation of 0.5), within a factor of 0.98 at $5 \times 10^{10} M_{\odot}$ (with a standard deviation of 0.4) and a factor of 0.96 at $10^{11} M_{\odot}$ (with a standard deviation of 0.4). For stellar masses $>10^{10} M_{\odot}$ the *WISE*-derived masses appear overall lower than the GAMA stellar masses. This is probably unsurprising given that the

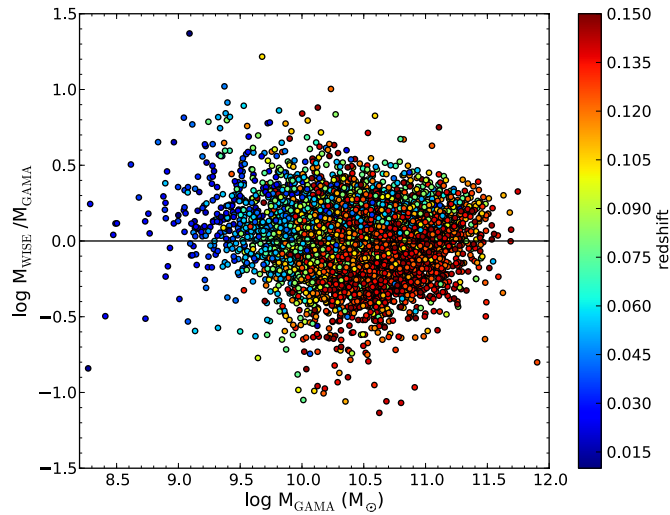


Figure 8. Difference between the *WISE*-derived (using Equation (2)) and GAMA stellar mass estimates as a function of GAMA stellar masses (Taylor et al. 2011), color-coded by redshift. The mass estimates agree within a factor of 1.2 at $10^{10} M_{\odot}$ (with a standard deviation of 0.5), within a factor of 0.98 at $5 \times 10^{10} M_{\odot}$ (with a standard deviation of 0.4) and a factor of 0.96 at $10^{11} M_{\odot}$ (with a standard deviation of 0.4).

(A color version of this figure is available in the online journal.)

sample is dominated by star-forming galaxies (see Figure 5) and the *WISE* $3.4 \mu\text{m}$ band is sensitive to the light from evolved stars in passively evolving galaxies.

4.3. Star Formation Rate Comparisons

4.3.1. $22 \mu\text{m}$ Warm Dust Continuum

IRAS, *ISO*, and *Spitzer* revolutionized our understanding of dust emission as a tracer of star formation. In the mid-infrared, the *Spitzer* MIPS $24 \mu\text{m}$ band measures the warm dust continuum excited by hot, young stars and is therefore sensitive to recent star formation, as well as AGN-activity. Numerous studies have investigated its stability as a measure of SFR (see, e.g., Alonso-Herrero et al. 2006; Calzetti et al. 2007; Rieke et al. 2009) and we are able to transfer this understanding to *WISE* and its $22 \mu\text{m}$ band.

We determine the luminosity density (νL_{ν}) for the $22 \mu\text{m}$ band and normalize by the total solar luminosity ($L_{\odot} = 3.839 \times 10^{26} \text{ W}$). The W4 band is the least sensitive *WISE* band, but we impose a S/N cut of 7 to ensure high quality photometry. Sources flagged as AGN based on optical spectroscopy diagnostics are removed from the sample. Cross-matching these sources with SFRs available for GAMA I, using sources with SFR between 0.1 and $100 M_{\odot} \text{ yr}^{-1}$ (see Gunawardhana et al. 2013), yields the best-fit relation shown in Figure 9. The optical SFRs are determined from $\text{H}\alpha$ equivalent widths applying an extinction-correction based on the Balmer decrement (full details can be found in Gunawardhana et al. 2013) and are sampled to $z < 0.35$ (i.e., beyond which $\text{H}\alpha$ is redshifted out of the observed spectral range). Galaxies where the $\text{H}\alpha$ line is contaminated by atmospheric O_2 (*A* band) absorption, in the redshift range $0.155 < z < 0.170$ are not included in the sample as recommended in Gunawardhana et al. (2013). The points in Figure 9 are color-coded according to Balmer decrement and on average higher infrared luminosities correspond to higher Balmer decrements (i.e., increased dust obscuration). However, without detailed knowledge of the dust geometry this is merely a crude comparison (see also Wijesinghe et al. 2011a, 2011b),

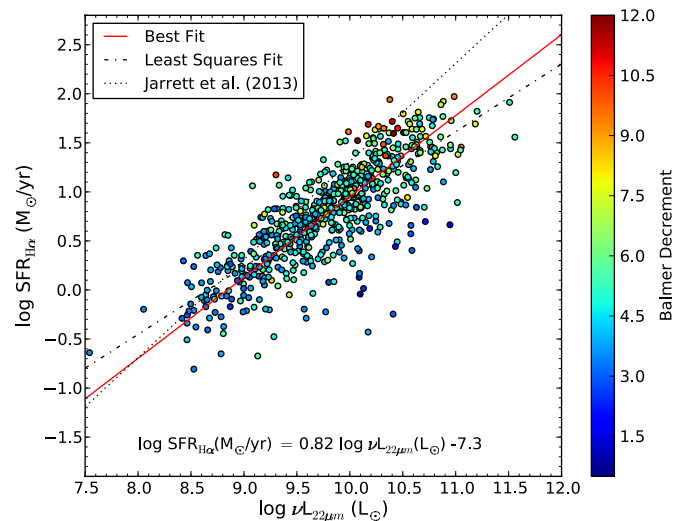


Figure 9. $\text{H}\alpha$ -derived star formation rates as a function of $\nu L_{22 \mu\text{m}}$ luminosity color-coded by Balmer decrement ($\text{H}\alpha/\text{H}\beta$) from Gunawardhana et al. (2013).

(A color version of this figure is available in the online journal.)

and as we discuss below, biases due to extinction are probably in play at higher redshifts. Balmer decrement has also been shown to be correlated with stellar mass, stellar mass surface density, and metallicity (see for example, Boselli et al. 2013).

Figure 9 also shows the linear fit of Jarrett et al. (2013), where the *WISE* $22 \mu\text{m}$ based SFR relation was calibrated using *Spitzer* $24 \mu\text{m}$ photometry and the relation of Rieke et al. (2009). This fit is somewhat steeper, likely explained by the distinction made by Rieke et al. (2009) based on total infrared luminosity of the source and the relatively small number of sources in the Jarrett et al. (2013) sample (but which also included Local Group dwarf galaxies). The best fit to the GAMA sample distribution is

$$\log_{10} \text{SFR}_{\text{H}\alpha} (M_{\odot} \text{ yr}^{-1}) = 0.82 \log_{10} \nu L_{22 \mu\text{m}} (L_{\odot}) - 7.3. \quad (4)$$

4.3.2. $12 \mu\text{m}$ ISM Tracer

Infrared-luminous sources will be detected by *WISE* $22 \mu\text{m}$ out to moderate redshifts ($z \sim 2$ to 3), but most typical galaxies would not. Hence there is a strong bias for W4 detections to be luminous galaxies (LIRGs and ULIRGs), which are highly obscured at optical wavelengths. The *WISE* $12 \mu\text{m}$ band, on the other hand, has greater sensitivity by comparison (see Table 1) and also probes the ISM, sensitive to a larger (representative) sampling of galaxies, and thus making it potentially the primary star formation indicator for *WISE*. The dominant feature within this W3 band is the $11.3 \mu\text{m}$ PAH, and to a lesser extent the $[\text{Ne II}]$ emission line. This PAH is large, neutral, and excited by ultraviolet radiation from young stars, as well as radiation from older, evolved stars (see, for example, Kaneda et al. 2008). $[\text{Ne II}]$ is associated with emission from H II regions. Passive disks generally lack the $7.7 \mu\text{m}$ PAH tracing current star formation, but still show prominent $11.3 \mu\text{m}$ PAH features (see, for example, Cluver et al. 2013). As first demonstrated with *Spitzer* measurements (see for example Houck et al. 2007; Farrah et al. 2007), PAHs can be used to estimate star formation, although there is larger scatter (and potential biases) relative to the superior $24 \mu\text{m}$ mid-infrared or $70 \mu\text{m}$ far-infrared tracers. The *WISE* $12 \mu\text{m}$ band is thus a powerful, yet also more problematic, tracer of recent star formation.

We proceed as with W4, but with a S/N cut of 10 in the W3 ($12 \mu\text{m}$ band), similarly color-coding to reflect Balmer

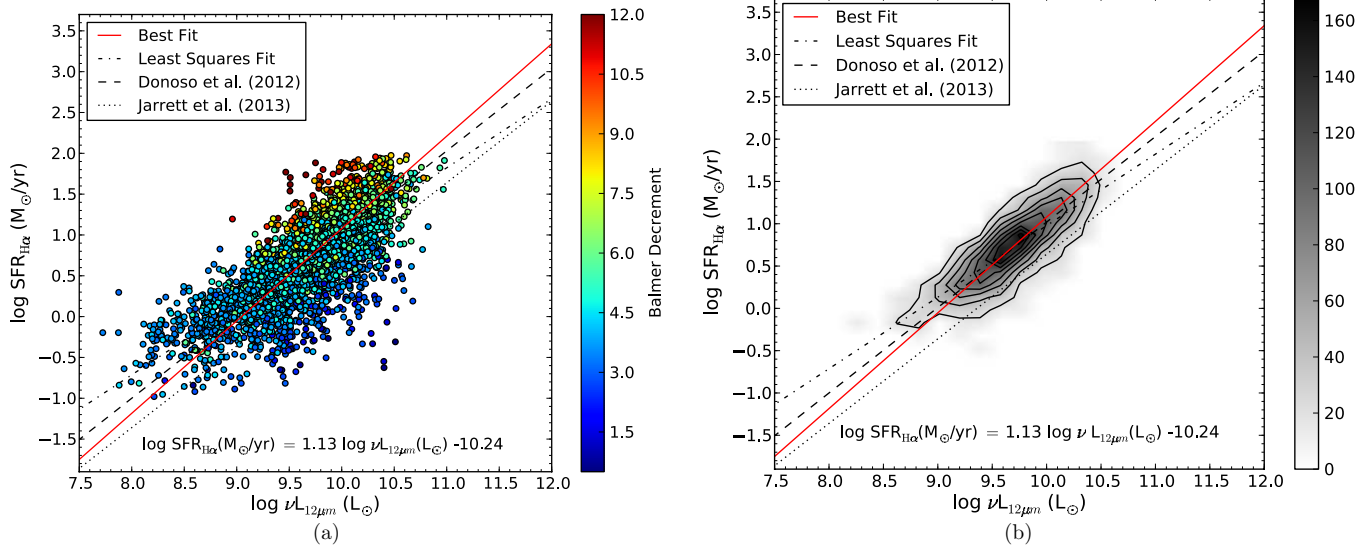


Figure 10. $H\alpha$ -derived star formation rates as a function of $\nu L_{12\mu\text{m}}$ luminosity color-coded by Balmer decrement in panels (a) and (b) shown as a contour and density plot.

(A color version of this figure is available in the online journal.)

decrement (Figure 10(a)). A contour plot of the distribution is shown in Figure 10(b) and the trend appears very tight for $\text{SFR} > 5 M_{\odot} \text{yr}^{-1}$. At low SFR and νL_{ν} , however, the distribution appears to flatten and probably accounts for the differences in slope between our relation and the least squares fit of Donoso et al. (2012). The relation of Jarrett et al. (2013) lies offset below our best fit, most likely as a result of relatively low SFR within their small sample ($\text{SFR} < 5 M_{\odot} \text{yr}^{-1}$) of nearby galaxies. The best-fit relation for the $12\mu\text{m}$ band is

$$\log_{10} \text{SFR}_{H\alpha} (M_{\odot} \text{yr}^{-1}) = 1.13 \log_{10} \nu L_{12\mu\text{m}} (L_{\odot}) - 10.24. \quad (5)$$

This flattening of the distribution in Figure 10 at low SFR is most likely a feature of the $11.3\mu\text{m}$ PAH tracer (since we do not observe it at $22\mu\text{m}$) and is also observed by Lee et al. (2013) in their study.

Given that the $12\mu\text{m}$ band emission appears too low for the given $H\alpha$ SFR, a possible explanation is that the relative abundance of PAH molecules to big grain emission is diminished, due to low metallicity in these galaxies. If this is the case, we do not expect to see a similar effect at $22\mu\text{m}$, since this band is dominated by big grains in equilibrium with the strong radiation fields inside star formation regions. Indeed, the $22\mu\text{m}$ sample of Lee et al. (2013) does not show such a flattening.

An alternative explanation would be that in a low SFR regime, the preponderance of the diffuse medium increases with respect to the star-formation component of the ISM (galaxies are more quiescent). Bear in mind that the $24\mu\text{m}$ warm dust emission is powered by the UV radiation fields from the massive stars inside the star-forming regions, while the $12\mu\text{m}$ emission is mainly powered by the diffuse interstellar radiation fields (Popescu et al. 2011). Thus, the PAH and small grains emission seen in the $12\mu\text{m}$ band increases with increasing SFR, but this increase is mediated by the propagation of the UV photons in the diffuse ISM, while the $24\mu\text{m}$ emission directly traces star formation regions, as does the Balmer line corrected $H\alpha$ emission from which the SFRs are derived.

To investigate this further we show individual points for contiguous redshift slices overplotted on the full sample in Figure 11; this serves as an indication of how the distribution

is built up. The flattening at low SFR appears to be a feature of nearby sources alone (Figure 11(a)) supporting the hypothesis that we are observing the PAH features and small grain emission of relatively quiescent galaxies, dominated by the diffuse ISM. The flattening in the slope is not detected at higher redshifts ($z > 0.05$) as these galaxies have more star formation activity (i.e., more clumpy). In Figures 11(b)–(d) we see how as we move to higher redshifts, we sample systems with higher $\nu L_{12\mu\text{m}}$ and, on average, higher SFRs. The relatively low number of sources in Figure 11(d) reflects the atmospheric contamination affecting sources at redshifts of $z \simeq 0.16$.

The higher infrared luminosity in these sources, however, also indicate greater obscuration in the optical bands; this explains the sharp drop in $H\alpha$ luminosities relative to W3, apparent in Figures 10(c) and (d) with $\text{SFR}(H\alpha) < 3 M_{\odot} \text{yr}^{-1}$. Corrections using the Balmer decrement become ineffective when the extinctions are high, ($A_V \gg 1$), which is why infrared tracers of obscured star formation and UV/optical tracers of unobscured star formation are best combined to estimate the total SFR.

5. SCIENCE WITH THE GAMA-WISE CATALOG

The GAMA-WISE catalog and empirically derived relations can be used any number of ways to probe the behavior of large populations of galaxies. To illustrate, we highlight below how WISE can be used to study the specific star formation of galaxies.

5.1. Specific Star Formation

We can investigate the specific star formation rate (sSFR) of the entire sample by using $L_{12\mu\text{m}}$ -derived SFRs (Equation (5)) and the GAMA stellar masses (Taylor et al. 2011). In Figure 12 we show specific star formation as a function of stellar mass, color-coded by redshift. Two clear trends are seen: (1) lower mass galaxies are actively building their disks while massive galaxies have expended their gas reservoirs rendering mostly passive evolution (the “SFR– M_* ” relation) and (2) with increasing redshift, a shift to higher SFR for a given stellar mass as we capture these infrared-luminous systems. The behavior of star formation and specific star formation, as a function of stellar

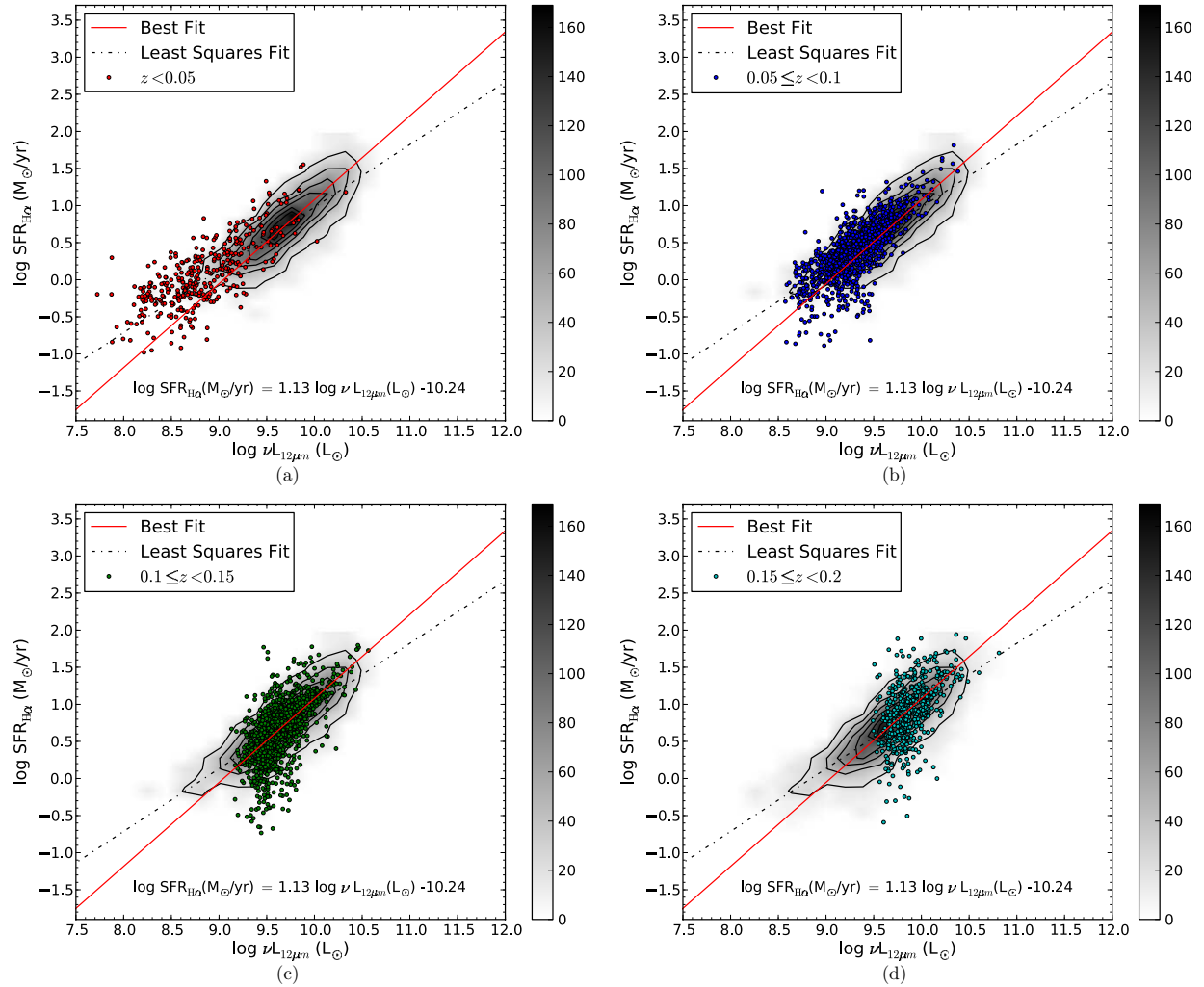


Figure 11. $H\alpha$ -derived star formation rates as a function of $\nu L_{12\mu\text{m}}$ luminosity for redshift slices of 0.05 between $0 < z < 0.2$ overlotted on the full distribution from Figure 9.

(A color version of this figure is available in the online journal.)

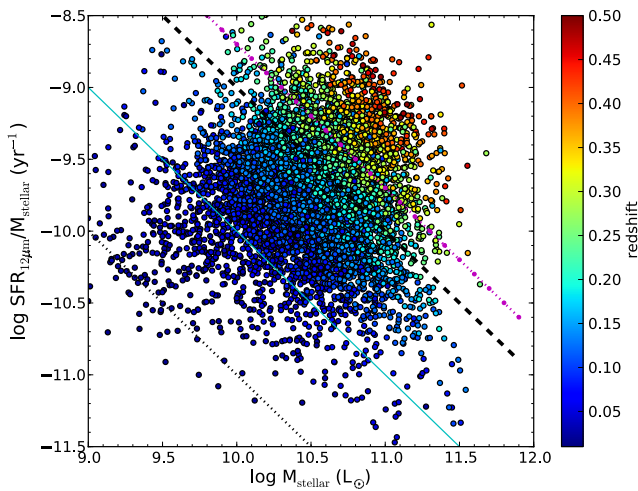


Figure 12. Specific star formation with redshift. The star formation derived from the $\nu L_{12\mu\text{m}}$ luminosity, for the entire sample $\nu L_{12\mu\text{m}}$ as a function of stellar mass, giving the sSFR, color-coded by redshift. Lines of constant SFR (0.1, 1, 10, and $20 M_{\odot} \text{yr}^{-1}$) are shown as dotted, solid, dashed, and dash-dot lines, respectively.

(A color version of this figure is available in the online journal.)

mass and redshift, within the GAMA sample is explored in detail in Bauer et al. (2013) and Lara-López et al. (2013). The observed increase in mean sSFR of star-forming galaxies with increasing redshift is well-established (see for example Noeske et al. 2007; Elbaz et al. 2007; Rodighiero et al. 2010) and we illustrate here that the GAMA-*WISE* sample is sufficiently large and diverse to explore galaxy evolution between the local universe and $z < 0.5$.

From the relation of Bouché et al. (2010), the “main-sequence” of galaxies with stellar mass of $10^{11} M_{\odot}$ at $z \simeq 0.5$ have typical SFRs of $\simeq 20 M_{\odot} \text{yr}^{-1}$. Converting to luminosity density by way of Equation (5) gives $\log_{10} L_{12\mu\text{m}} \simeq 10.2 L_{\odot}$. As illustrated in Figure 13, *WISE* can detect these systems that are within the GAMA-*WISE* sample. A luminosity density of $\log_{10} L_{12\mu\text{m}} \simeq 10.2 L_{\odot}$ corresponds to a flux density of 0.24 mJy or 12.8 mag; using the S/N detection statistics of the G15 sources, *WISE* will detect this source with $S/N \simeq 20$. The magnitude sensitivity limit in the W3 band for the G12 field is 13.2 (10σ), 14.0 (5σ) and for G15, 13.8 (10σ) and 14.5 (5σ).

5.2. Specific Star Formation of Galaxies in Pairs

The $12\mu\text{m}$ derived SFRs, computed from Equation (5), can also be used to investigate sSFR trends within populations of

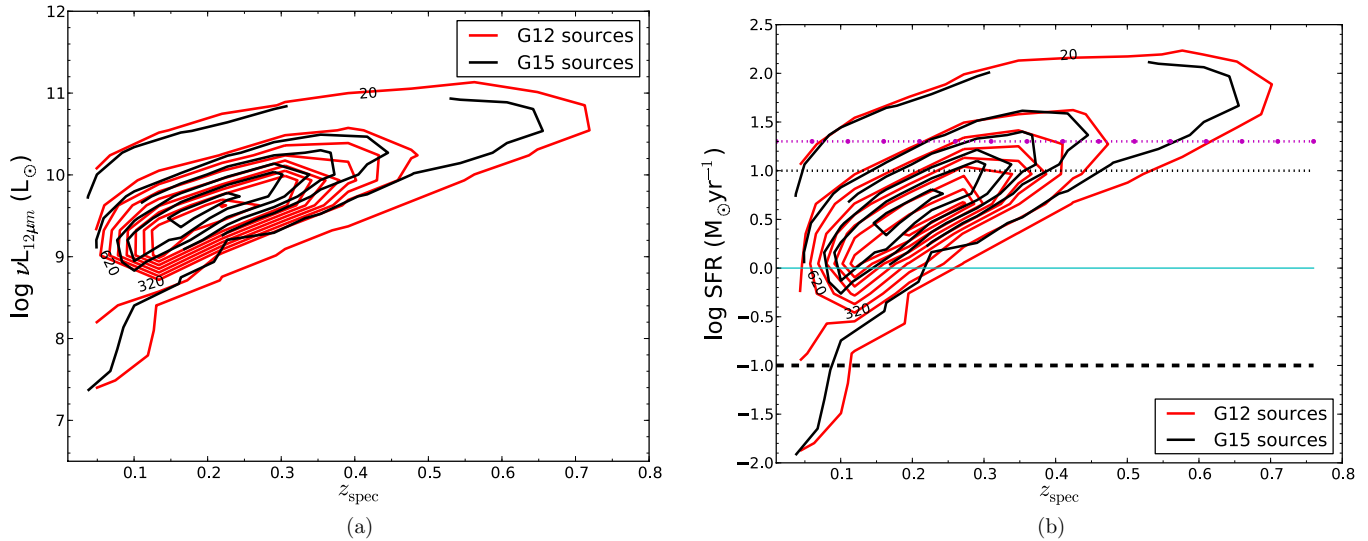


Figure 13. Redshift distribution of sources in the two GAMA fields relative to (a) the $12\ \mu\text{m}$ luminosity density (νL_ν) and in (b) converting luminosity density to SFR using Equation (5). Lines of constant SFR (0.1 , 1 , 10 , and $20\ M_\odot\ \text{yr}^{-1}$) are shown as dotted, solid, dashed, and dash-dot lines, respectively.

(A color version of this figure is available in the online journal.)

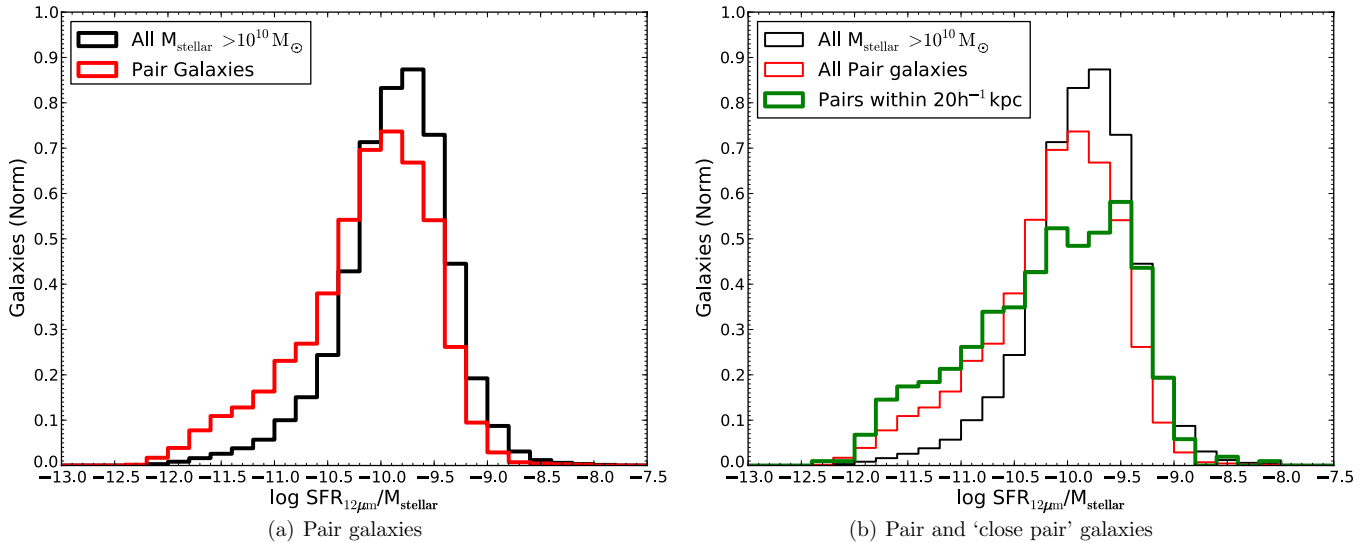


Figure 14. Specific star formation rate normalized distribution of galaxies with stellar masses $> 10^{10}\ M_\odot$ and the subsets that reside within a pair and a close pair (separation $< 20\ h^{-1}\ \text{kpc}$).

(A color version of this figure is available in the online journal.)

GAMA galaxies. For example, we use the updated GAMA II Galaxy Group Catalogue ($G^3\text{Cv6}$), as detailed in Robotham et al. (2011), to isolate galaxies that reside within a pair. In $G^3\text{C}$ a galaxy pair is defined as two galaxies within $100\ h^{-1}\ \text{kpc}$, in physical separation, and $1000\ \text{km}\ \text{s}^{-1}$ in velocity separation. Figure 14 is a normalized histogram of all galaxies in the sample with stellar mass $> 10^{10}\ M_\odot$ and the subset that reside in a pair (17,475 galaxies) and a close pair (i.e., separation $< 20\ h^{-1}\ \text{kpc}$). We note that potential blends (as outlined in Section 3.5) are removed.

The striking feature of Figure 14(a) is that galaxies that reside within a pair appear to have, on average, lower sSFR than a typical GAMA galaxy of the same mass, suggesting that instead of star formation being enhanced in these systems, it is suppressed. Although initially counter-intuitive, and contrary to studies of local galaxy pairs compared to field control samples (see, for example, Ellison et al. 2013), it should be noted

that $\sim 40\%$ of GAMA galaxies reside within a group, with numerous mechanisms at play. Within the context, therefore, of the environment of a typical galaxy this highlights the complexity of interacting and merging systems and is discussed further in a detailed study of merging and interacting galaxies within GAMA (A. S. G. Robotham et al., in preparation). We note that a lack of SFR enhancement has emerged from studies of galaxy pairs probing higher redshift samples (see, for example, Xu et al. 2012), with one suggested explanation that higher gas fractions at higher redshift reduce the efficiency of torque-driven gas infall.

Further we include in Figure 14(b) a histogram of galaxies within $20\ h^{-1}\ \text{kpc}$ of their neighbor, as compared to the distribution from Figure 14(a). It appears that this subset shows a broader range of sSFRs compared to the larger sample of galaxies in pairs, and may even be bimodal, consisting of suppressed systems and enhanced systems. Adding this

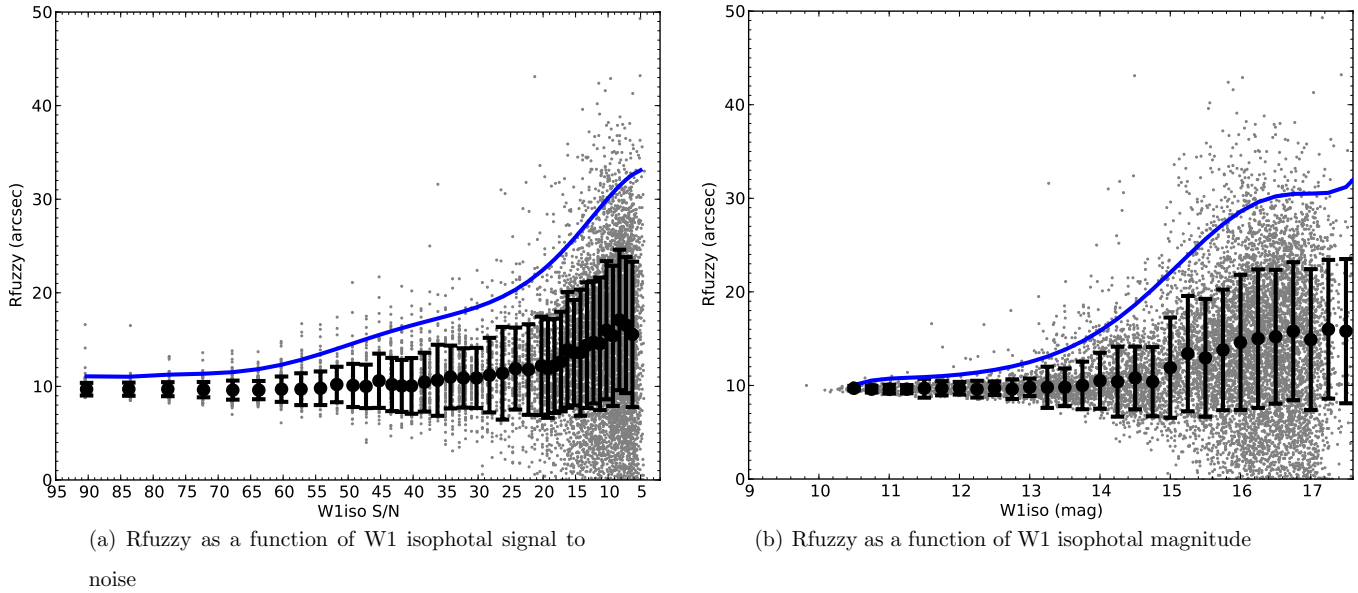


Figure 15. $R_{\text{fuzzy}}^{\text{limit}}$ is calibrated using sources in G12 that are expected to be unresolved, i.e., $w1\text{rchi}2 < 2$.
(A color version of this figure is available in the online journal.)

kinematic trigger to galaxy evolution is clearly a complex process, but crucial to understand how galaxies evolve in the group environment. Finally, we note that blending limitations of the *WISE* data within pairs that have a smaller separation make this parameter space uncertain, but optical tracers will be exploited here in forthcoming GAMA papers.

6. CONCLUSIONS

We have detailed the procedure for creating a source-matched catalog between galaxies in the GAMA G12 and G15 fields, and *WISE* photometry. In particular we have outlined how best to extract photometry for low S/N, unresolved and resolved sources in the *WISE* All-Sky Catalog. Complete GAMA-*WISE* catalogs for the G09, G12, and G15 fields will be made available through the GAMA Public Releases (see <http://www.gama-survey.org/>).

Using the *WISE* measurements and matched GAMA galaxies in the G12 and G15 fields, we have investigated the following:

1. The *WISE* color distribution for the sample shows most systems are globally dominated by star formation, with few passive and AGN-dominated systems; this is consistent with the GAMA sample selection and spectroscopic sensitivity to higher redshift, star-forming galaxies.
2. Empirical relations of stellar mass as a function of $3.4\ \mu\text{m} - 4.6\ \mu\text{m}$ color for resolved sources, our entire sample and star-formation-dominated galaxies only. We provide relations that can be applied to large samples for $z < 0.5$.
3. SFR relations can be derived using $22\ \mu\text{m}$ and $12\ \mu\text{m}$ luminosities. The $12\ \mu\text{m}$ derived SFR relation (Equation (5)) is a complex tracer of the ISM and should, however, be used with caution. We show that the distribution of galaxies in the $12\ \mu\text{m}$ luminosity and $H\alpha$ SFR plane as a function of redshift are affected by selection effects and most probably dust geometry.
4. Specific star formation (using $12\ \mu\text{m}$ derived SFRs) for the sample illustrates how the GAMA-*WISE* catalog detects only the most massive, highest star-forming systems at the highest redshifts. *WISE* is, however, able to detect star-

forming main-sequence systems, of stellar mass $\sim 10^{11}\ M_{\odot}$, to $z \simeq 0.5$ with $S/N > 10$.

5. Controlling for stellar mass, galaxies with an associated neighbor appear to experience, on average, a shift to lower specific star formation. Extracting pairs with relatively small separations ($< 20\ h^{-1}\ \text{kpc}$) suggests a broader behavior consistent with populations experiencing either star formation suppression or enhancement.

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GAMA is a joint European-Australasian project based around a spectroscopic campaign using the Anglo-Australian Telescope. The GAMA input catalog is based on data taken from the Sloan Digital Sky Survey and the UKIRT Infrared Deep Sky Survey. Complementary imaging of the GAMA regions is being obtained by a number of independent survey programs including *GALEX* MIS, VST KiDS, VISTA VIKING, *WISE*, *Herschel*-ATLAS, GMRT, and ASKAP providing UV to radio coverage. GAMA is funded by the STFC (UK), the ARC (Australia), the AAO, and the participating institutions. The GAMA Web site is <http://www.gama-survey.org/>.

APPENDIX

ANALYSIS OF *WISE* PHOTOMETRY FOR NOMINALLY RESOLVED SOURCES

In order to understand the photometry of resolved and partially resolved sources in *WISE* we use several diagnostics.

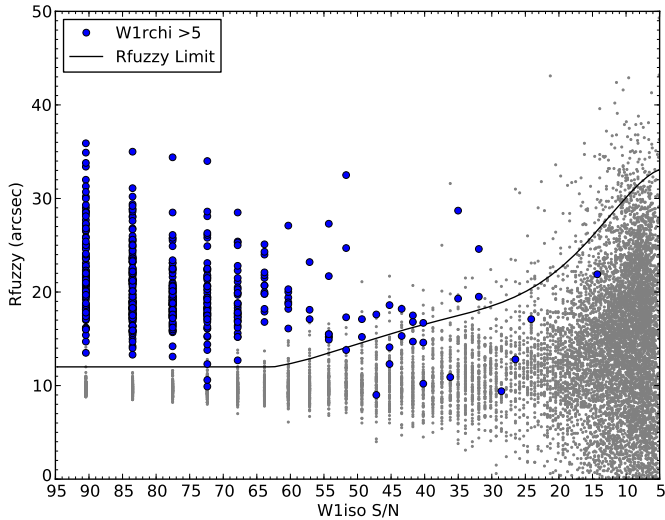


Figure 16. Performance of Rfuzzy is shown for sources in G15 where the majority are expected to be resolved, i.e., $w1rchi2 > 5$.

(A color version of this figure is available in the online journal.)

Since *WISE* is calibrated in the Vega system, images and all-sky catalogs, all comparisons are done as such.

A.1. Rfuzzy Defined

The Rfuzzy parameter is used to determine whether a source is resolved in the $3.4 \mu\text{m}$ band of *WISE* (see Section 3.2). The Rfuzzy parameter is measured in the following way:

1. The source is rotated to determine the optimal 2nd moment parameters (R_{major} , R_{minor} , position angle of major axis).
2. The source is divided along the minor axis into two halves: positive x , and negative x , where the central x position is the nominal position of the source.
3. Compute the 2nd order intensity-weighted moment for each half and derive the R_{major} for each half: R_{major_a} , R_{major_b} .
4. $R_{\text{fuzzy}} = \text{minimum}(R_{\text{major}_a}, R_{\text{major}_b})$.

If $R_{\text{fuzzy}} < R_{\text{fuzzy}_{\text{limit}}}$, then the source is unresolved. The $R_{\text{fuzzy}_{\text{limit}}}$ is determined by measuring sources that are expected to be unresolved. For example, Figure 15 shows the Rfuzzy values for *WISE* sources in the GAMA G12 region with $w1rchi2 < 2$ which preferentially selects unresolved sources. The

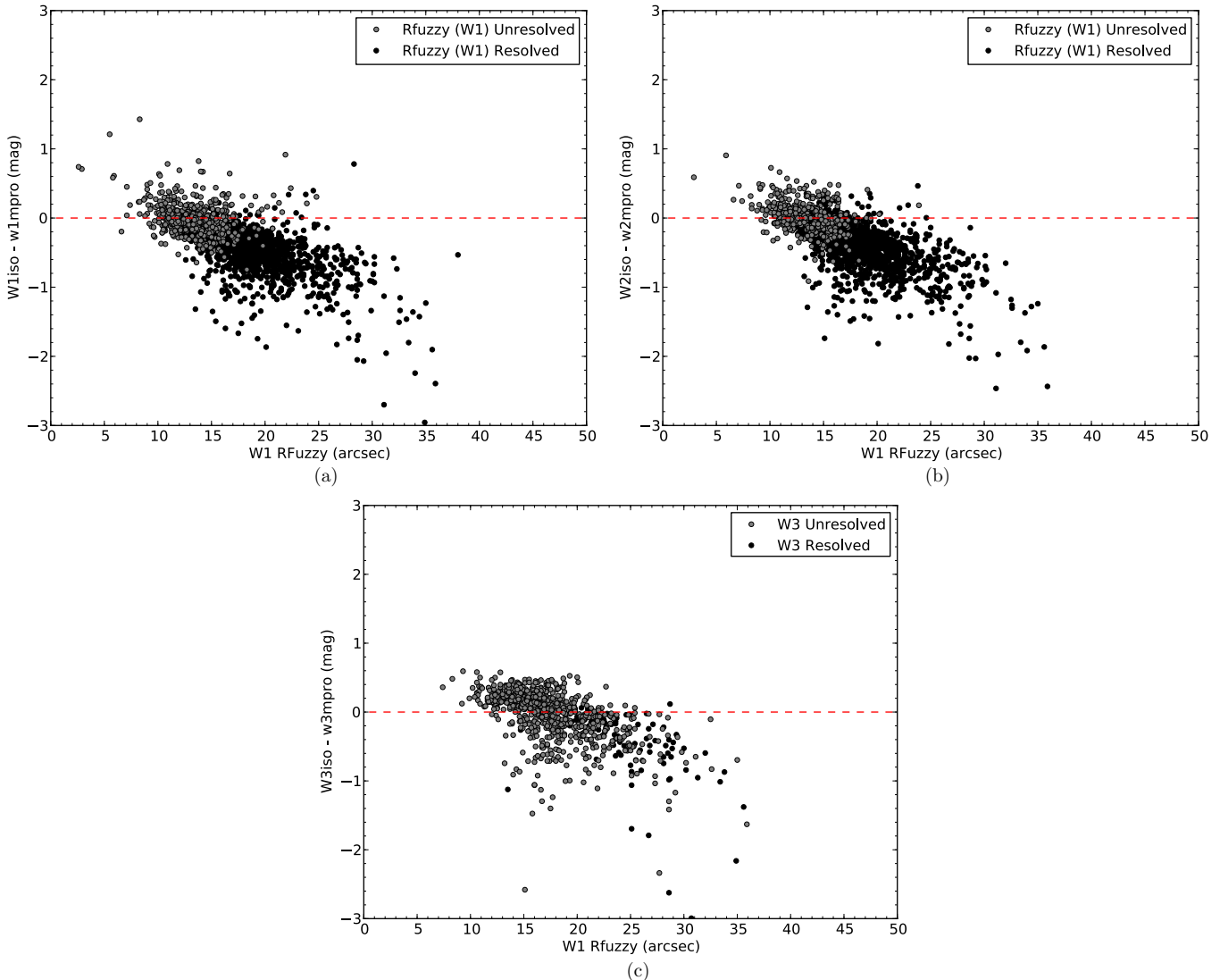


Figure 17. Difference between a 1σ isophotal measurement (Section 3.1) and a profile-fit measurement, as a function of the Rfuzzy parameter (with low signal-to-noise sources removed) for 2590 sources in G15.

(A color version of this figure is available in the online journal.)

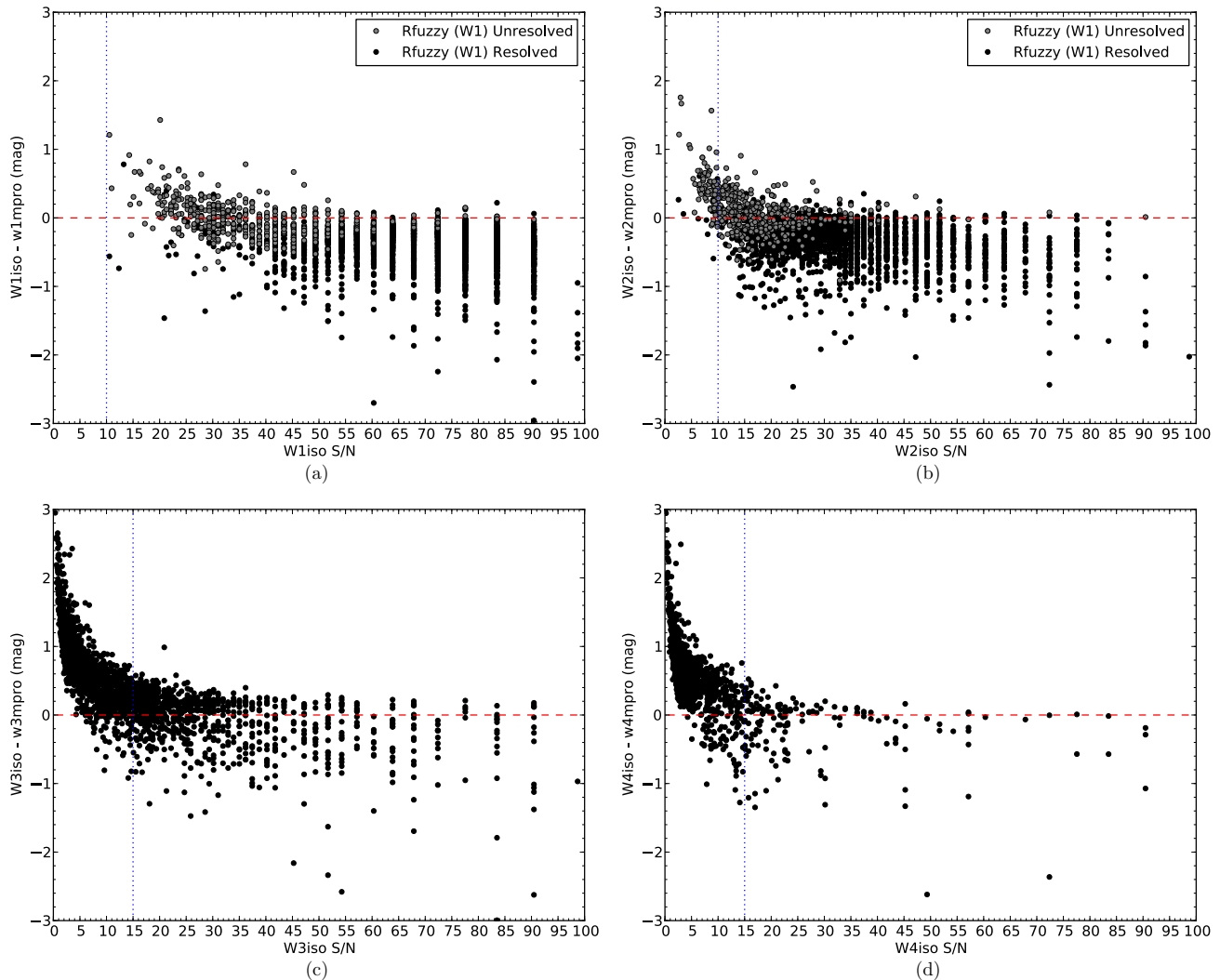


Figure 18. Difference between an isophotal measurement and a profile-fit measurement as a function of signal to noise (S/N) in the isophotal measurement in each band. The vertical dotted lines indicate the S/N limits where the isophotal photometry becomes less reliable and $w\star mpro$ photometry is recommended. (A color version of this figure is available in the online journal.)

typical value for R_{fuzzy} is $\simeq 10\%$ for high S/N point sources (e.g., stars), but the distribution can be used to identify the limit where resolvedness can be determined. A power series fit to the 2σ mean of the distribution yields the $R_{fuzzy,limit}$ curve—points that lie above this curve are classified as resolved.

R_{fuzzy} can be defined in terms of isophotal S/N or isophotal magnitude. Given the differing exposure coverage (number of *WISE* frames that samples a given patch of sky) for the entire sky, the sensitivity limits vary depending on location on the sky. For this reason, a magnitude-dependent $R_{fuzzy,limit}$ curve will be sensitive to the *WISE* coverage and would need to be derived for each region. However, a S/N function is robust against this and therefore the curve shown in Figure 15(a) is used.

As a test of the performance of the R_{fuzzy} parameter, Figure 16 plots R_{fuzzy} for a sample from GAMA G12 expected to be dominated by resolved systems (based on $w1rchi2$). The majority of sources with $w1rchi2 > 5$ lie above the curve defined by the point sources in Figure 15(a).

In Figure 17 we test 2590 sources in G15 chosen using the prescription of Section 3, i.e., most of which are expected to be resolved in W1. This shows that the largest offset between the $W1iso$ and $w1mpro$ photometry measurements occurs for

sources deemed resolved by the R_{fuzzy} parameter. This behavior is consistent with sources whose flux is underestimated by profile fitting, i.e., resolved. Unresolved sources show the expected scatter around 0. In Figure 17(c) we see that few sources are resolved in the W3 band.

A.2. Behavior of Resolved Photometry

In order to establish the reliability of isophotal photometry, we use our test sample (2590 potentially resolved sources from the G15 GAMA field). In Figure 18 we compare the difference between $W\star iso$ and $w\star mpro$ photometry as a function of S/N ratio. At low S/N the isophotal measurement becomes unreliable—in some cases the fluxes are inflated, notably for W3 and W4—due to contamination from background sources and $w\star mpro$ provide the most robust measurement.

In Figure 19 we again use the difference between $W\star iso$ and $w\star mpro$ photometry to illustrate that the sensitivity of the W1 and W2 bands prevents the $w1rchi2$ and $w2rchi2$ values from acting as a reliable discriminator of resolvedness. For W3 and W4, however, it can be used and the limits derived from the plots are indicated.

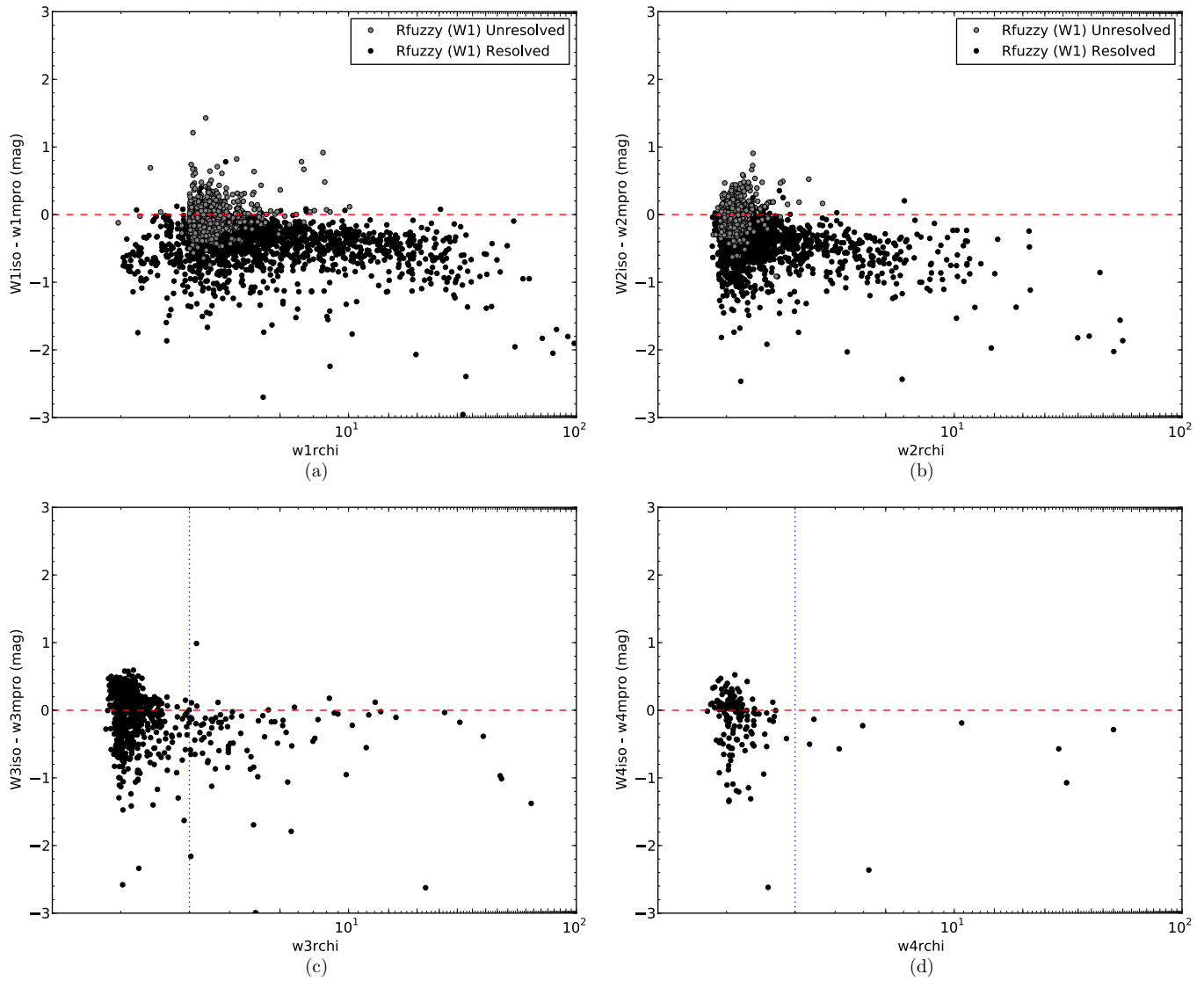


Figure 19. Difference between an isophotal measurement and a profile-fitted measurement as a function of $w\star chi^2$ with low signal to noise sources removed. The vertical lines indicate the limits for resolved sources in the W3 and W4 bands. (A color version of this figure is available in the online journal.)

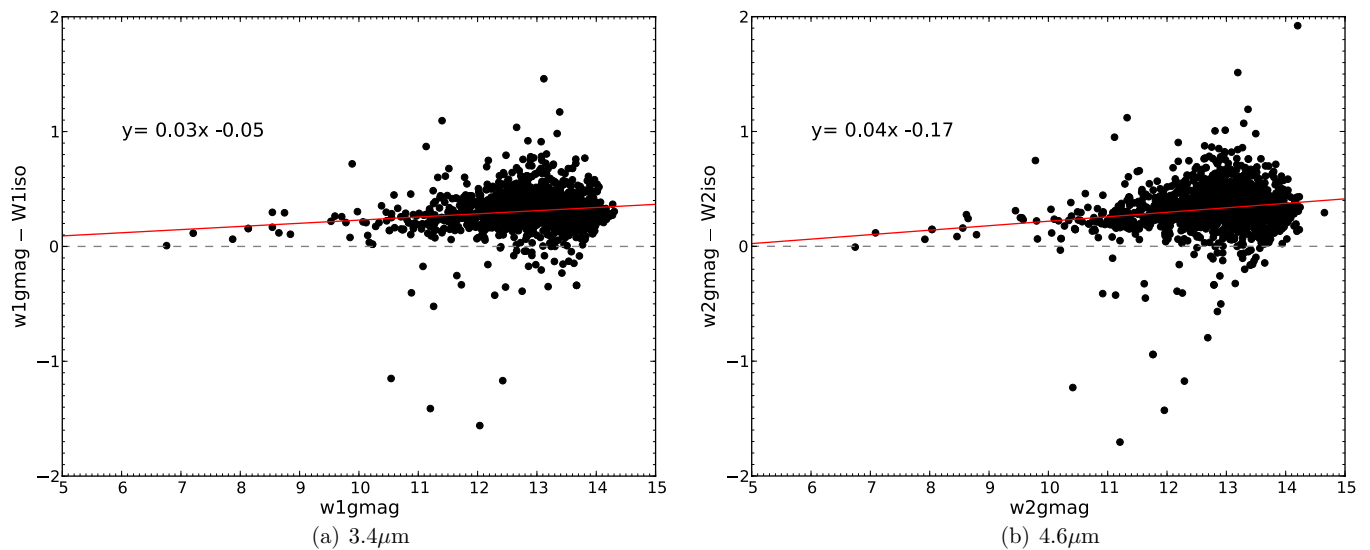


Figure 20. Comparison of $w\star mag$ and $W\star iso$ photometry as a function of $w\star mag$. (A color version of this figure is available in the online journal.)

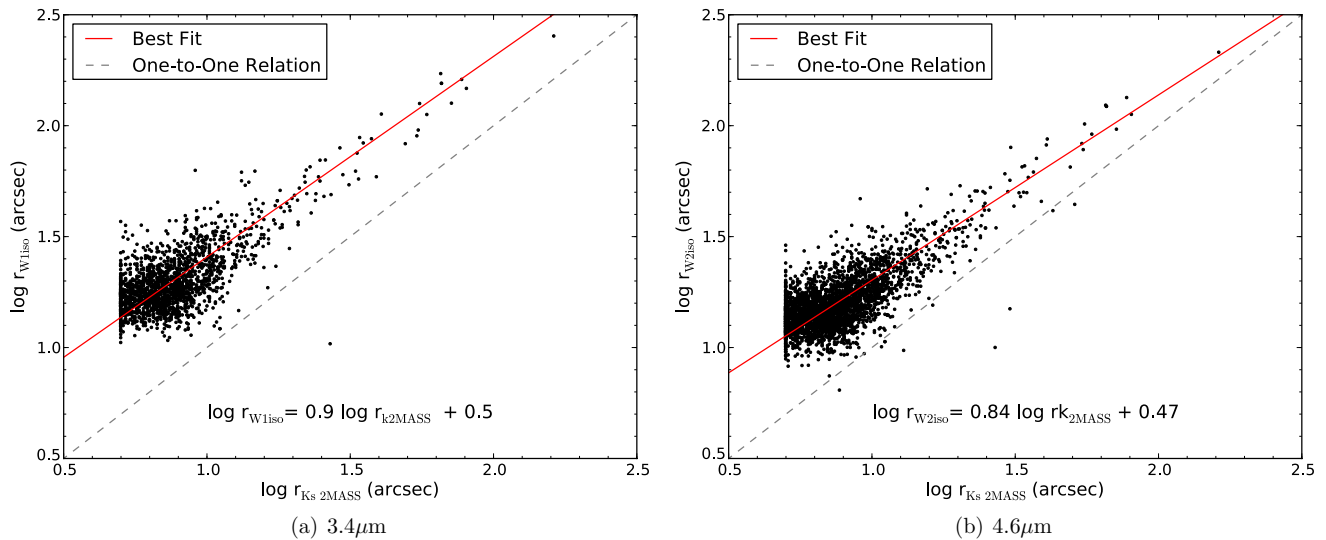


Figure 21. Comparison of isophotal and w^*gmag radii for resolved sources in G12 and G15; the dashed line indicates a one-to-one relation. (A color version of this figure is available in the online journal.)

We explore the relationship between the isophotal photometry and the w^*gmag photometry for sources in the 2MASS XSC. Since the w^*gmags are measured using elliptical apertures with radii scaled to twice the 2MASS radii, this provides an indication of the additional sensitivity. We note that w^*gmags can be contaminated by nearby objects since no attempt is made to remove neighboring sources. In Figure 20 we show the difference between w^*gmag and W^*iso photometry for resolved galaxies in G12 and G15. This shows that for fainter sources w^*gmag is brighter than W^*iso , likely due to contamination.

Finally, in Figure 21 the *WISE* isophotal and 2MASS isophotal radii of resolved sources which clearly shows that the 1σ isophotal radii are systematically larger than the 2MASS radii by a factor of 2–2.5. The largest offset occurs for the most compact 2MASS sources probably due to the increased sensitivity of *WISE* in the W1 and W2 bands resolving more of the galaxy relative to 2MASS.

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