



The Fate of Formamide in a Fragmenting Protoplanetary Disk

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Abstract

Recent high-sensitivity observations carried out with the Atacama Large Millimeter Array have revealed the presence of complex organic molecules (COMs) such as methyl cyanide (CH_3CN) and methanol (CH_3OH) in relatively evolved protoplanetary discs. The behavior and abundance of COMs in earlier phases of disk evolution remain unclear. Here, we combine a smoothed particle hydrodynamics simulation of a fragmenting, gravitationally unstable disk with a gas-grain chemical code. We use this to investigate the evolution of formamide (NH_2CHO), a prebiotic species, in both the disk and in the fragments that form within it. Our results show that formamide remains frozen onto grains in the majority of the disks where the temperatures are <100 K, with a predicted solid-phase abundance that matches those observed in comets. Formamide is present in the gas phase in three fragments as a result of the high temperatures (≥ 200 K), but remains in the solid phase in one colder (≤ 150 K) fragment. The timescale over which this occurs is comparable to the dust sedimentation timescales, suggesting that any rocky core that is formed would inherit their formamide content directly from the protosolar nebula.

Key words: astrochemistry – hydrodynamics – planets and satellites: formation – protoplanetary disks

1. Introduction

The origin of the prebiotic content at the surface of the primitive Earth remains unclear. It is postulated that a large quantity of prebiotic material might have been brought to Earth by comets and meteorites during its early formation (Caselli et al. 2012). This hypothesis is supported by the detection of numerous prebiotically relevant molecular species in comets such as Hale-Bopp (Bockelée-Morvan et al. 2000), Lemmon and Lovejoy (Biver et al. 2014), and more recently 67P/Churyumov–Gerasimenko (Goesmann et al. 2015; Altwegg et al. 2017).

Among these species is formamide (NH_2CHO), an important precursor of pre-genetic and pre-metabolic compounds such as nucleic acids, nucleobases, sugars, and amino acids (see Saladino et al. 2012). The reservoir of complex organic molecules (COMs)—such as formamide—in comets is expected to derive from the previous stages of star formation. Formamide has indeed been detected in the hot envelopes around low-mass protostars (e.g., Kahane et al. 2013). The unrivalled sensitivity of the Atacama Large Millimeter Array (ALMA) has opened up the possibility to also study COM chemistry in protoplanetary disks. However, only two COMs have been detected so far in the gas phase in these objects: methyl cyanide (CH_3CN , Öberg et al. 2015) and methanol (CH_3OH , Walsh et al. 2016). A large proportion of these COMs are thought to remain frozen onto dust grains, making their detection in the gas phase challenging (see, e.g., Walsh et al. 2014).

There have been several theoretical studies of the evolution and inheritance of chemistry during the youngest stages of protoplanetary disks. Some have considered the chemical link

between a surrounding envelope and an axisymmetric disk (e.g., Visser et al. 2011; Drozdovskaya et al. 2016; Furuya et al. 2017), finding that chemical composition depends on the route taken by material to the forming disk. Other works have chosen to concentrate on the chemistry of isolated disks, modeling the full three-dimensional hydrodynamic evolution within the disk itself (e.g., Ilee et al. 2011; Evans et al. 2015; Yoneda et al. 2016), finding that dynamically induced changes in the temperature and density of the disk material can significantly influence the chemical composition. To date, only one study has followed the time-dependent chemical evolution of a disk that forms protoplanetary fragments (Ilee et al. 2017), but this work did not consider the formation and evolution of complex molecules. Therefore, the formation and evolution of COMs from the protostellar phase to the birth of a planetary system remains unclear.

In this paper, we investigate the fate of one such COM, formamide, during one of the earliest phases of disk evolution. We consider a massive, self-gravitating protostellar disk that undergoes fragmentation, which we model using smoothed particle hydrodynamics (SPH) coupled with a gas-grain chemical evolution code. We examine the abundance of formamide in the different environments of the disk (e.g., cold outer regions, regions undergoing shocks) and particularly in the four protoplanetary fragments formed during the simulation. We show that our models broadly reproduce observations of COMs in comets, and that formamide may survive around the fragments long enough to be incorporated into protoplanets' rocky cores, implying the possibility of direct inheritance of formamide from the protosolar epoch.

2. Methodology

2.1. Hydrodynamic Modeling

We use the same radiation-hydrodynamic disk simulation as in Ilee et al. (2017), which is also “Run 2” from Hall et al. (2017). Full details of the simulation setup are given in these works, so we only briefly reiterate key aspects here. We simulate a $0.25 M_{\odot}$ disk, with a $1 M_{\odot}$ central protostar. The gas in the disk is represented by 4×10^6 SPH particles distributed between 10 and 100 au, with an initial surface density profile that scales as $\Sigma(r) \propto r^{-1}$ and an initial sound speed profile that scales as $c_s \propto r^{-0.5}$. We include the modified radiation transfer scheme of Forgan et al. (2009) in which the gas is able to cool radiatively according to its local optical depth (calculated from the gravitational potential), and can exchange energy with neighboring fluid elements via flux-limited diffusion.

Four fragments are formed during this simulation (labeled with ascending radial distance from the central star as John, Paul, George and Ringo; see Ilee et al. 2017). John and Paul are tidally disrupted at $t = 3081$ yr and $t = 3514$ yr, respectively, while George and Ringo survive until the end of the simulation. John contains the most mass ($10.3 M_{\text{Jup}}$) but undergoes tidal disruption, George is the most massive surviving fragment ($8.2 M_{\text{Jup}}$), Ringo reaches $5.4 M_{\text{Jup}}$ at the end of the simulation, and Paul is the least massive fragment, reaching $3.7 M_{\text{Jup}}$ before being disrupted.

2.2. Chemical Modeling

The chemical calculations are performed using UCLCHEM (Viti et al. 2004; Holdship et al. 2017), a gas-grain chemical evolution code. The gas-phase reactions are taken from the UMIST database (McElroy et al. 2013), with additional reactions from Quénard et al. (2018). The grain surface reaction network is the one presented in Quénard et al. (2018). Several thermal and non-thermal processes are considered on the grain surface, such as diffusion via thermal hopping and quantum tunneling (Hasegawa et al. 1992), thermal desorption (Hasegawa et al. 1992), cosmic-ray desorption, direct-UV and secondary-UV desorptions (McElroy et al. 2013; Holdship et al. 2017), and chemically reactive desorption (Minissale et al. 2016).

The network contains 364 species (243 in the gas phase and 121 on the grain surface) and 3446 reactions. The starting elemental abundances are the same as those in Quénard et al. (2018). Binding energies of molecules are taken from Wakelam et al. (2017; see Table 1). We note that we do not consider three-body reactions, expected to be important in high-density environments, within our chemical model. However, Ilee et al. (2011) included three-body reactions in their study of the chemical composition of a massive disk, finding that they have a small effect on the molecular abundances (see their Section 3.3).

The chemical modeling is performed in four steps, where the first three initialize the chemical conditions prior to the full protoplanetary disk phase (i.e., the final abundances of phases 1, 2 and 3 are used as initial abundances for phases 2, 3 and 4, respectively).

The first step (Phase 1) corresponds to the diffuse cloud phase, where we follow the chemistry of a low-density cloud ($n_{\text{H}} = 10^2 \text{ cm}^{-3}$ and $A_{\text{V}} = 2$ mag) over 10^6 years at a kinetic temperature $T_{\text{kin}} = 10$ K. Only gas-phase chemistry plays a

Table 1

Binding Energies of Formamide and Related Species Adopted in Our Chemical Model, Taken from Wakelam et al. (2017)

Species	Binding Energy E_{D} (K)
CO	1300
CO ₂	2600
CH ₄	960
HCN	3700
H ₂ O	5600
CH ₃ OH	5000
NH ₃	5500
NH ₂	3200
H ₂ CO	4500
NH ₂ CHO	6300

role at this stage since the A_{V} is low, and all grain surface species are either destroyed or desorbed via non-thermal processes (see the left panel of Figure 1).

The second step (Phase 2) is the pre-stellar phase, where the cloud contracts for $\sim 5.4 \times 10^6$ yr until it reaches a final density of $n_{\text{H}} = 1 \times 10^{11} \text{ cm}^{-3}$ following a freefall collapse parameterization (Rawlings et al. 1992; Holdship et al. 2017). The gas temperature remains at 10 K during this phase. The cloud remains at low $A_{\text{V}} \lesssim 4$ mag with $n_{\text{H}} \lesssim 2 \times 10^3 \text{ cm}^{-3}$ for the first $\sim 4.7 \times 10^6$ yr. During this period, the chemistry does not vary significantly and the abundances of formamide, as well as those of species such as CO, H₂O, H₂CO, and NH₂, remain relatively the same (see the left panel of Figure 1). At approximately 6.3×10^6 yr, almost all molecules experience a steep increase in gas-phase abundance. As discussed in Quénard et al. (2018), the high density allows molecular species to be efficiently produced on grain surfaces, since the high visual extinction prevents UV photons from destroying or desorbing molecules. These newly formed species are then injected into the gas phase via non-thermal desorption processes such as chemical reactive desorption. The abundances of H₂O, H₂CO, NH₂, and NH₂CHO drastically decrease after 6.4×10^6 years, as a result of the severe freezing out expected at the high-density regime in pre-stellar cores.

The third step (Phase 3) is the warming-up (or protostellar/hot corino) phase. While the density is kept constant and the temperature increases from 10 to 300 K following a rate defined by Awad et al. (2010), due to the increasing grain surface temperature, this phase enables the formation of COMs via radical-radical addition. Once the dust temperature exceeds the desorption energies shown in Table 1, molecular species are thermally desorbed into the gas phase. We also emphasize that higher temperatures also help to overcome the activation barrier of the endothermic reactions present in our gas-phase chemical network, which triggers a richer gas-phase chemistry.

The fourth step (Phase 4) is the protoplanetary disk phase, when we run UCLCHEM for a subset of 100,000 particles taken from the SPH simulation. We proceed in this way to make it possible to undertake the calculation of the chemical evolution of the disk in a reasonable time frame using the large network of UCLCHEM. The subset of particles was selected randomly, but the selection was confined to within a smoothing-length of the midplane in order to preserve the radial surface density structure of the disk. The subset is large enough to ensure that all the physical regimes of the disk are probed. For each particle within this subset, the starting abundance is taken to be

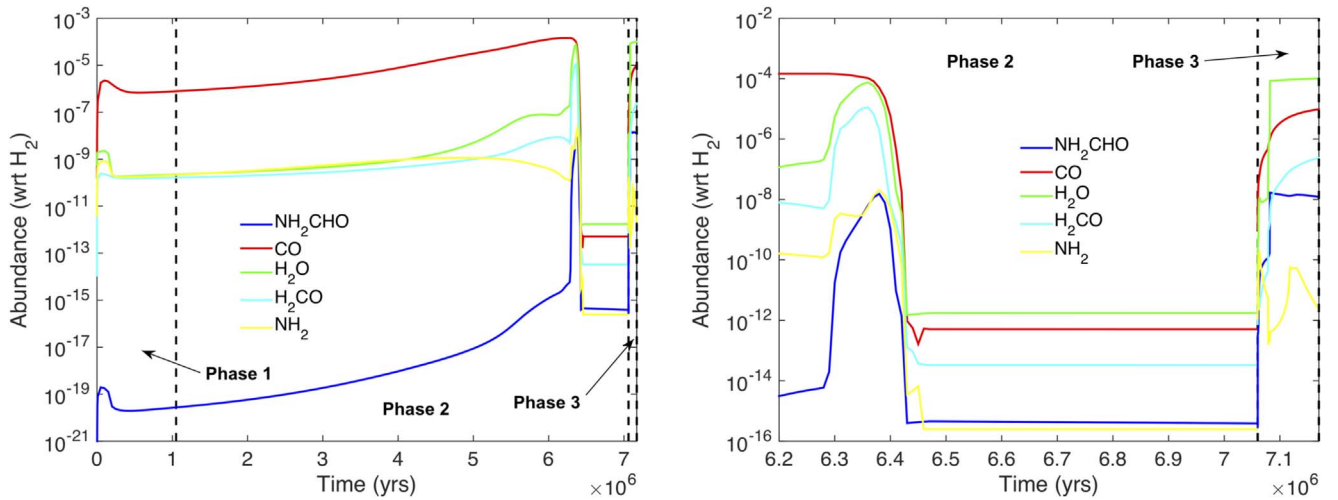


Figure 1. Gas-phase abundance of NH_2CHO (formamide), CO , H_2O , NH_2 , and H_2CO across the pre-disk phases in our chemical modeling. Left panel: abundances across phase 1–3 for the full ~ 7.2 Myr considered. Right panel: a closeup of the final Myr showing phases 2 & 3.

the final abundance of the warming-up (protostellar) phase. Since the subset is drawn from close to the midplane, we set the visual extinction (A_V) in this phase to >100 mag, such that any photo-processing can be considered negligible. For all phases, we assume an external radiation field of $G_0 = 1$ Habing and a standard cosmic-ray ionization rate of $\zeta = 1.3 \times 10^{-17} \text{ s}^{-1}$.

We have benchmarked the chemical abundances obtained in our work with those calculated by Ilee et al. (2017), who used a smaller chemical network (125 species–89 in the gas, 36 on dust grains—and 1334 reactions; see their Section 2.2). We find a good agreement between the two networks, especially for molecules that are key in the production of formamide (e.g., CO , CO_2 , H_2CO , H_2O , NH_2 , and NH_3), with abundances differing by less than a dex. This result is reassuring given the large discrepancy in size between the two networks.

3. Results

3.1. Formamide (NH_2CHO) in the Disk

Figure 2 shows the abundance of formamide in four different time snapshots in the protoplanetary disk. The first snapshot at 204 yr shows the initial high gas-phase abundance of formamide across the disk, inherited from the protostellar (hot corino) phase. Rapidly, around 850 yr, all species are frozen out onto dust grains because of the cold temperatures of the disk ($T \lesssim 20$ K) combined with high densities ($n(\text{H}_2) \gtrsim 10^{10} \text{ cm}^{-3}$). Later, formamide is released back into the gas phase via thermal desorption in regions where $T \gtrsim 100$ K. The magnitude of these effects depends strongly on the position of the particle within the disk (e.g., whether they reside in shocked regions, or within fragments). The total amount of formamide (gas phase plus solid phase) does not vary significantly with time, since the bulk of the formamide has already been produced during the protostellar phase. Therefore, the formamide abundance in the gas phase or solid phase is mainly driven by the temperature, leading either to freeze-out or to desorption.

In the innermost parts of the disk ($r < 15$ au), Figure 2 shows that formamide is in the gas phase. This is due to the high gas kinetic temperatures ($T = 100\text{--}400$ K), which thermally desorb formamide from grains. Note that the presences of UV and/or X-ray radiation fields from the central protostar are

not taken into account here. Therefore, the abundance of formamide at $r < 15$ au might be lower due to the photo-destruction of this molecule. At larger radii, the high visual extinction ($A_V \geq 470$ mag) associated with the large densities in the disk (up to 10^{13} cm^{-3} in the midplane) will shield formamide from any radiation fields (both stellar and external).

The gas-phase formamide abundance in the outermost parts of the disk typically remains lower than 10^{-12} , even reaching $\sim 10^{-18}$ in some regions. As the temperature is also low (close to 10 K), the bulk of the formamide remains frozen onto dust grains (with a constant abundance of 10^{-8}) in the arms and the outer parts of the disk. In Section 4.1, we compare these solid-phase abundances of formamide in the outer disk predicted by our model with those measured observationally in solar system comets.

Figure 3 presents the evolution of the column density of the disk overlaid with contours showing the location of the midplane snowlines for CO , H_2O , and NH_2CHO . The snowlines are defined by the radii at which the gas-phase column density of a species is equal to that of the column density of the corresponding ice-phase species, i.e., $N_{\text{gas}} = N_{\text{ice}}$ (see Section 4.2 of Ilee et al. 2017). The CO snow line is much larger than the one for H_2O because of the larger desorption temperature for the latter. Since formamide is desorbed under similar conditions to water (their binding energies are similar, $E_D = 5600$ K for H_2O versus $E_D = 6300$ K for NH_2CHO , see Table 1), their snowlines are practically the same. We note that the formamide/water snow line is only present for the hottest fragments (George and John; middle panel of Figure 3), because the amount of gas-phase formamide (or water) in Paul and Ringo is not high enough to match the amount of this molecule on the surfaces of dust grains.

3.2. Radial Distribution of Formamide in the Fragments

Figure 4 presents the median abundance of formamide for the four fragments as a function of distance from the fragment’s center for different times. These abundances are calculated by taking into account the median of all abundance values within shells of 0.2 au in width, with 50 shells covering a radial distance of 10 au around each fragment). The time range for each fragment is taken between their formation time and the end of the simulation (or destruction time where appropriate,

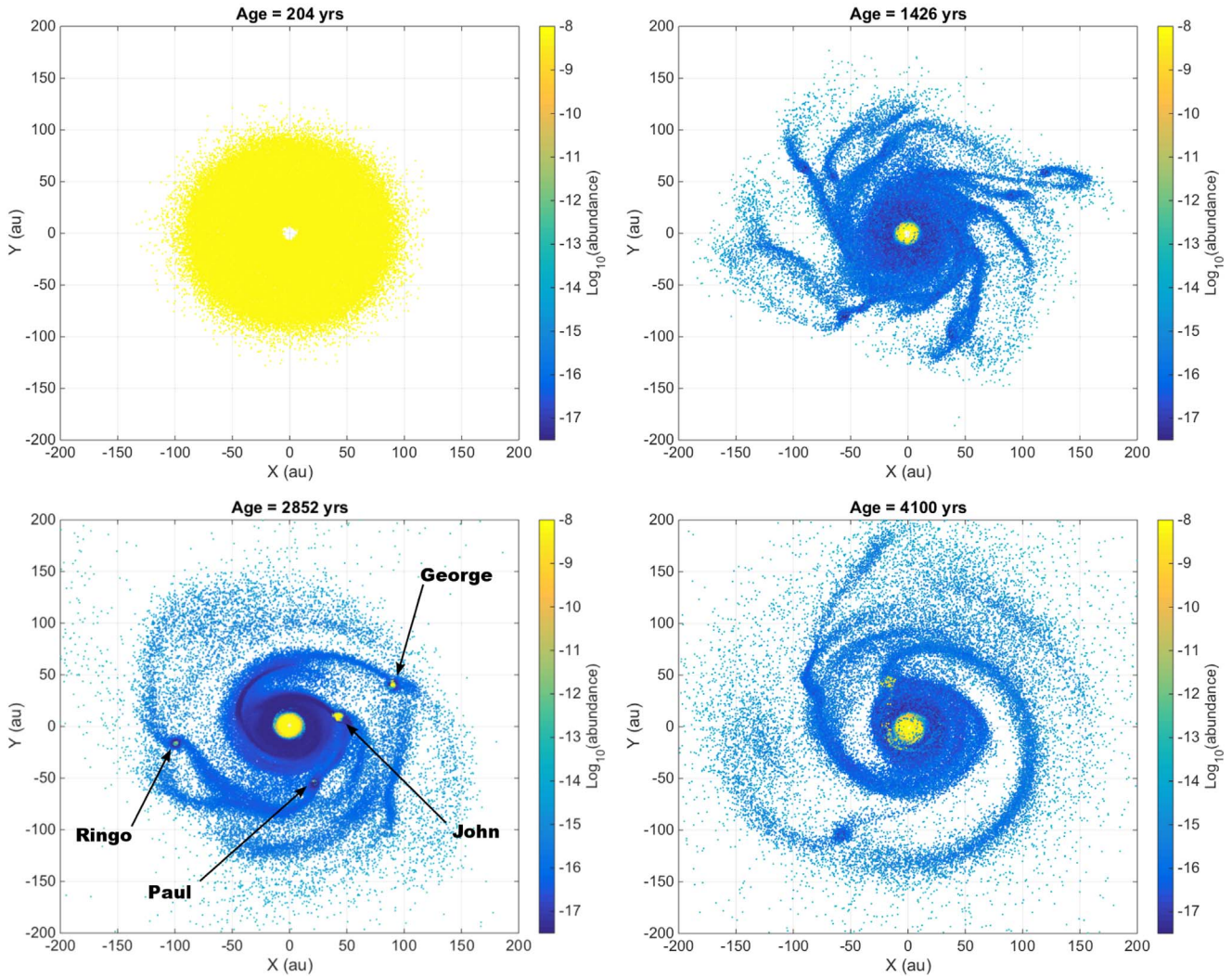


Figure 2. Gas-phase abundance of formamide across the protoplanetary disk for four different snapshots covering 4100 yr of disk evolution. Formamide is kept on the grain surface in the majority of the disk, but with time it goes back to the gas phase in the central part of the disk and around some fragments.

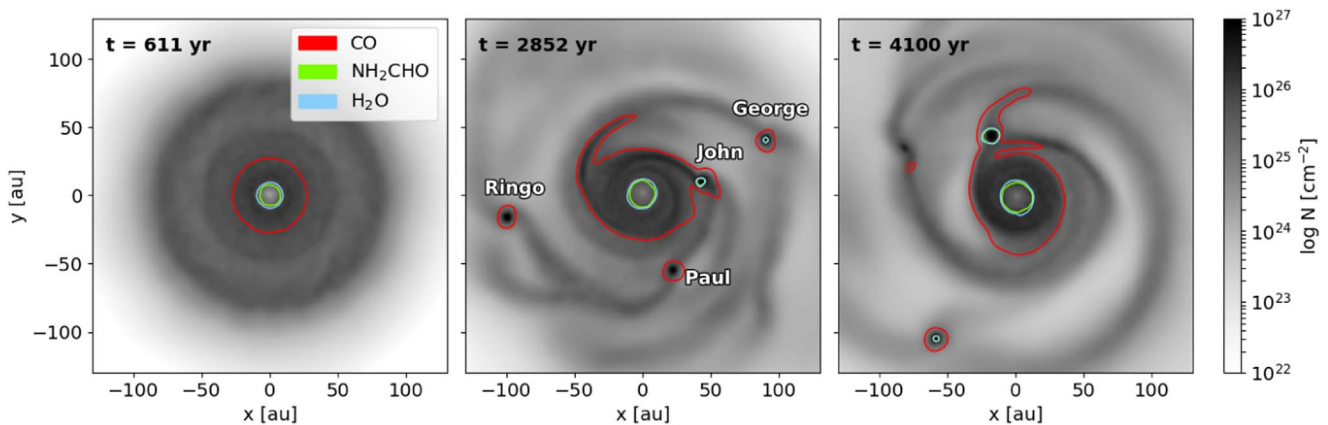


Figure 3. Evolution of the total column density of the disk (gray scale) overlaid with the snowlines of CO (red), NH_2CHO (green), and H_2O (blue). The snowlines quickly deviate from the expected concentric ring structure due to the dynamic evolution of the disk.

i.e., 2037–3081 yr for John; 1630–3514 yr for Paul; 2037–4100 yr for George; and 2445–4100 yr for Ringo—see also Table 2 of Ilee et al. 2017).

For George, the radius of gas-phase formamide increases with time, to be ~ 7 au at the end of the simulation. This is due

to the increasing temperature of the fragment (up to temperatures of 1100 K), caused by its inward migration and increasing size with time. For Ringo, the radius of gas-phase formamide stays rather constant at around ~ 2.5 au. Indeed, Ringo has a relatively constant orbital radius as a function time,

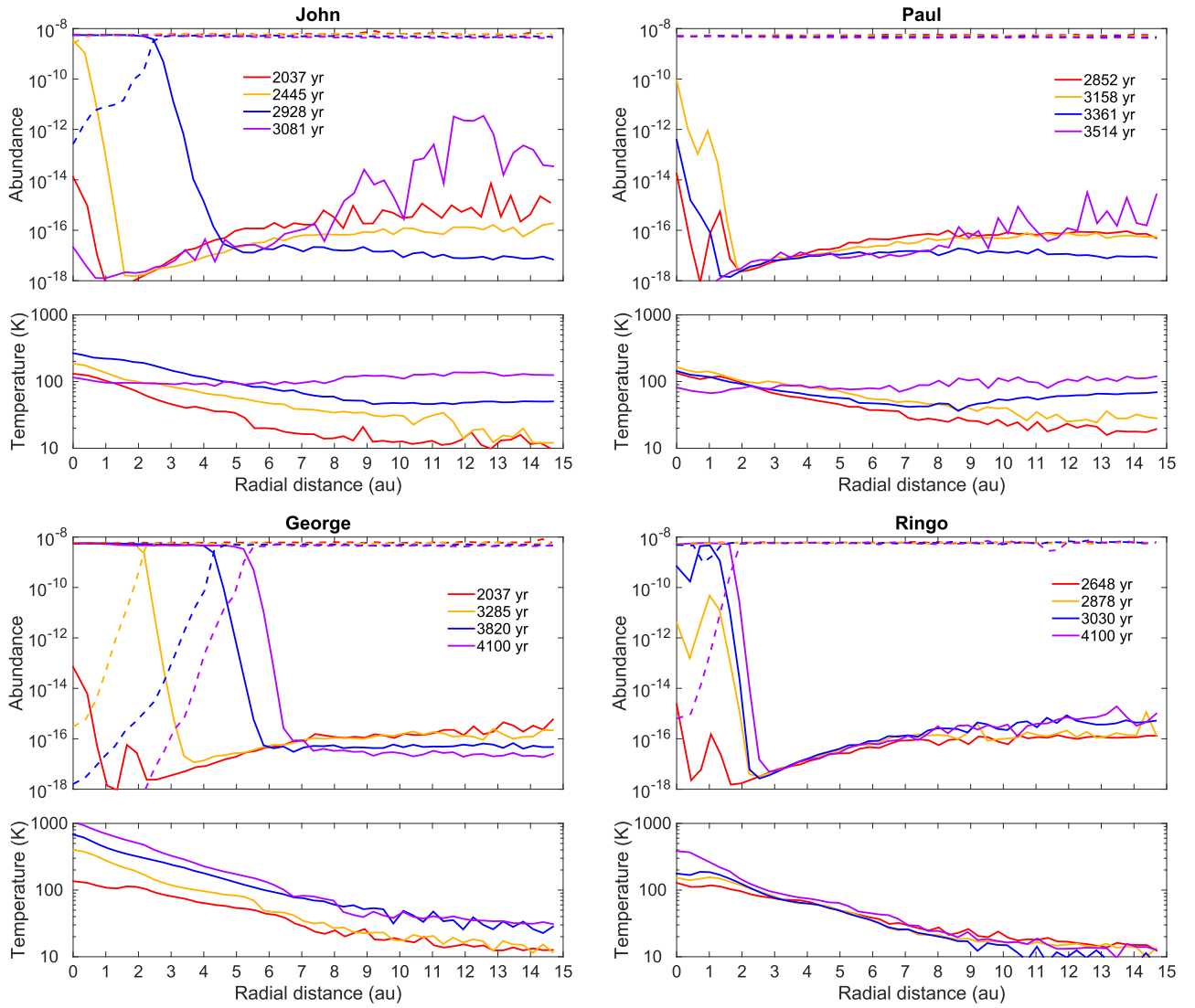


Figure 4. Median abundance of formamide in the gas phase (full lines) and on the grain surface (dashed lines) as a function of the radius of each fragment: John (top left), Paul (top right), George (bottom left), and Ringo (bottom right). The lower panels show the median temperature profile around each fragment. Each color represents a different time, comprised between the formation and destruction time (if any) of each fragment.

thus its radial temperature profile remains the same with time ($T \geq 100$ K for $R \lesssim 2.5$ au). The temperature in Paul is much lower than that of the other fragments (150 K versus 300–1100 K, at most), hence the majority of formamide remains frozen onto dust grains. Indeed the gas-phase abundance in Paul is less than a few 10^{-10} within 0.5 au, corresponding to $<1\%$ of the total formamide abundance of 10^{-8} available on the grains.

4. Discussion

4.1. Comparison with Cometary Abundances of Formamide

In Section 3.1, we found that formamide is fully frozen on the surface of dust grains soon after the start of the simulation in the outermost regions of the disk and outside the fragments, as a result of their low temperatures and high densities ($T \leq 10$ K and $n(\text{H}_2) \geq 10^{10}$ cm $^{-3}$). The solid-phase abundance of formamide in these regions therefore does not change significantly during the remaining evolution of the disk. In a similar way, comets are thought to be formed from

pristine material at the very edge of protoplanetary disks that do not undergo significant chemical processing. Therefore, the solid-phase abundance of formamide in our model at timescales ~ 850 years from the start of the disk simulation should be representative of such a situation. In order to test this, we gathered observationally determined abundances of formamide and related species for three long-period comets—C/1995 “Hale-Bopp,” C/2012 F6 “Lemmon” and C/2013 R1 “Lovejoy” (see Bockelée-Morvan et al. 2000; Biver et al. 2014; Le Roy et al. 2015, and references therein). Table 2 shows these abundances (relative to H_2O) alongside abundances measured in the outer disk regions early in the disk simulation. We find that these abundances agree to within a factor of 5 for all species, and note an excellent agreement between the modeled and observed values of formamide. As such, our results show that even massive, unstable disks that undergo gravitational fragmentation can possess cold outer regions within which the abundance of complex organics can be similar to those observed in long-period comets in the solar system.

Table 2

Solid-phase (Ice) Abundances in the Outer Parts of the Disk Model Compared to Observed Abundances in Long-period Comets

Species	Model (%)	Observed ^a (%)
NH ₂ CHO	0.02	0.01–0.021
NH ₃	1.08	0.61–0.7
H ₂ CO	0.23	0.7–1.1
CO	9.65	4.0–23.0
CO ₂	27.4	6
CH ₄	1.65	0.6–0.67
CH ₃ OH	1.00	1.6–2.6
HCN	0.766	0.14–0.25

Notes. All values are relative to water (in percent).

^a Long-period comets: C/1995 O1 (Hale-Bopp), C/2012 F6 (Lemmon) and C/2013 R1 (Lovejoy). The values are taken from Biver et al. (2014), Bockelée-Morvan et al. (2000), and Le Roy et al. (2015; and references therein).

4.2. Formamide Preservation around “Cold” Fragments

In Section 3.2, we show that most of the formamide content in Paul ($\sim 99\%$) remains on the grain surface during its existence. The “cold” environment around Paul is a direct result of a smaller size and mass (~ 2 au in radius and mass of $3.7 M_{\text{Jup}}$) compared to the other three fragments (with radii between 2.5 and 7 au and masses between 5.4 and $10.3 M_{\text{Jup}}$). Paul does not survive long enough to further contract and heat up its surroundings ($T \leq 150$ K), so water and many other molecules remain on icy grain surfaces. As proposed by Ilee et al. (2017), this effect may have important consequences for the formation of protoplanetary objects if the timescale for dust settling within the fragments is small enough. Indeed, Ilee et al. (2017) inferred that the timescale for centimeter-sized silicate grains to settle to the center of the fragment would be 1000–2000 yr, which is within the ~ 2000 yr of existence of Paul. Such dust settling might even occur more quickly, either due to icy grains possessing a high sticking co-efficient than bare grains (e.g., Ehrenfreund et al. 2003; Wang et al. 2005), or via hydrodynamic instabilities within the fragments themselves (Nayakshin 2018). Both effects would allow a higher proportion of solid material to sediment toward the center of the fragment. If rocky protoplanetary cores are formed within these regions, then formamide (and other grain species, including larger complex organics) would be incorporated into them. Such rocky cores may go on to become terrestrial-like planets through e.g., tidal downsizing (Nayakshin 2017; Forgan et al. 2018), sequestering primordial complex organics within them.

4.3. Formamide Preservation around “Hot” Fragments

As shown in Section 3.2, formamide is thermally desorbed into the gas phase around hot fragments due to high temperatures. In the innermost regions of the fragments where $T \geq 500$ K, formamide is destroyed via gas-phase reactions for timescales longer than those considered in these simulations. While at larger radii, the temperature drops to levels at which formamide can be retained on dust grains (Figure 4). Thus, while forming protoplanetary fragments would be devoid of solid-phase species, the circumfragmentary material around them would not be. Here, we discuss the content of this circumfragmentary material, and the potential makeup of any planetary satellites that may go on to form from it.

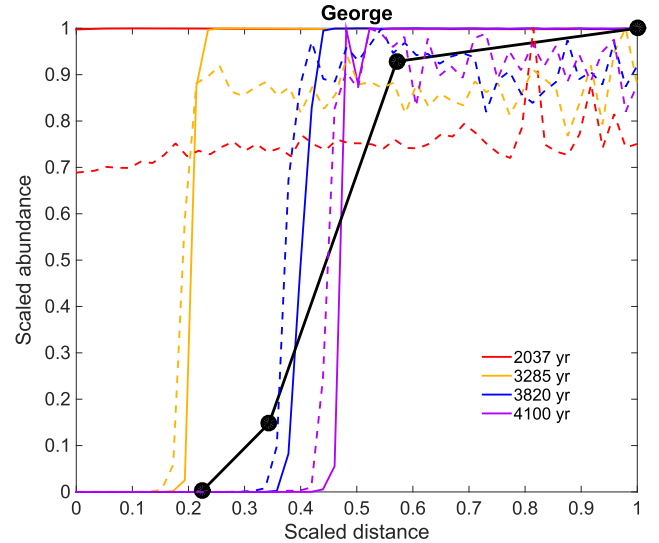


Figure 5. Scaled water (full colored lines) and formamide (dashed colored lines) abundances on the grain surface as a function of the radius of George. The scaled distance used for George is equal to 12 au. Each color represents a different time. The black line show the water content inside the four Galilean moons, as a function of their distance from Jupiter (scaled to the distance of Callisto).

In Figure 5, we present the distribution of water (in solid lines) and formamide (see dashed lines) in solid phase around George, the hottest surviving fragment in our simulations (different colors correspond to different timescales). We also overplot the solid water content measured in the Galilean moons: 46–48 wt% [percentage by mass] for Ganymede, 49–55 wt% for Callisto, 0% for Io, and 6–9 wt% for Europa (Kuskov & Kronrod 2001, 2005). The x-axis has been scaled to the radius of Callisto and to the radius for which the temperature around George equals the temperature of the disk (i.e., at a radius of 12 au). From Figure 5, we find that the fraction of solid water and formamide in our simulations increases with radius in a similar fashion to that observed for the solid water content in the Galilean moons. Indeed, the low temperatures (~ 40 K) and high water abundances found at the outer radius around George are consistent with an icy moon like Callisto being formed at low temperatures, which could have prevented any melting of the ice (while Io, Europa, and Ganymede are completely differentiated bodies, Callisto presents a mantle without any separation between the ice and the rock material; Kuskov & Kronrod 2005). Therefore, formamide could also be retained in satellites formed around protoplanetary fragments from gravitationally unstable disks.

We stress that this comparison is valid regardless of the formation scenario proposed for a Jupiter-like system (e.g., whether the planet is formed via core accretion or gravitational instability; Canup & Ward 2002; Mosqueira & Estrada 2003), because the expected shapes of the radial temperature profile around nascent protoplanets formed via either mechanism are the same (see Cleaves et al. 2015; Szulágyi et al. 2017). As such, water and formamide will always follow a similar radial behavior as shown in Figure 5. Thus, while significant chemical processing can still occur in the later, Class II stages of protoplanetary disk evolution, our results demonstrate that such processing is not required to match the radial abundance profiles of solar system satellites, and similarities can be in place at early times for objects formed via gravitational fragmentation of the disk.

5. Summary

We have studied the formation and evolution of the key prebiotic molecule formamide in the context of a young, massive circumstellar disk that gravitationally fragments into bound protoplanets. We find that formamide is located in the gas phase in the innermost ($r < 15$ au) regions of the disk, and in three of the four fragments. It remains on the grain surfaces in the cold, outer regions of the disk where its abundance is consistent with that observed in long-period comets. In one fragment, formamide remains on dust grains long enough to allow sedimentation, implying that rocky protoplanetary cores formed via, e.g., tidal downsizing, can inherit primordial prebiotic material directly from the protostellar nebula. In the circumfragmentary material, the radial abundance of solid formamide and water is consistent with the behavior of solid water observed across Galilean satellites.

In combination, our results show that when calculated explicitly, the chemical composition of protoplanets and protosatellites formed in gravitationally unstable disks can be similar to those expected from core accretion, and can even be consistent with the composition of objects within our own solar system.

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Software: sphNG (Bate et al. 1995); UCLCHEM (Holdship et al. 2017).


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References

- Altwegg, K., Balsiger, H., Berthelier, J. J., et al. 2017, *MNRAS*, 469, S130
- Awad, Z., Viti, S., Collings, M. P., & Williams, D. A. 2010, *MNRAS*, 407, 2511
- Bate, M. R., Bonnell, I. A., & Price, N. M. 1995, *MNRAS*, 277, 362
- Biver, N., Bockelée-Morvan, D., Debout, V., et al. 2014, *A&A*, 566, L5
- Bockelée-Morvan, D., Lis, D. C., Wink, J. E., et al. 2000, *A&A*, 353, 1101
- Canup, R. M., & Ward, W. R. 2002, *AJ*, 124, 3404
- Caselli, P., Keto, E., Bergin, E. A., et al. 2012, *ApJL*, 759, L37
- Cleeves, L. I., Bergin, E. A., & Harries, T. J. 2015, *ApJ*, 807, 2
- Drozdzovskaya, M. N., Walsh, C., van Dishoeck, E. F., et al. 2016, *MNRAS*, 462, 977
- Ehrenfreund, P., Fraser, H. J., Blum, J., et al. 2003, *P&SS*, 51, 473
- Evans, M. G., Ilee, J. D., Boley, A. C., et al. 2015, *MNRAS*, 453, 1147
- Forgan, D., Rice, K., Stamatellos, D., & Whitworth, A. 2009, *MNRAS*, 394, 882
- Forgan, D. H., Hall, C., Meru, F., & Rice, W. K. M. 2018, *MNRAS*, 474, 5036
- Furuya, K., Drozdovskaya, M. N., Visser, R., et al. 2017, *A&A*, 599, A40
- Goesmann, F., Rosenbauer, H., Bredehöft, J. H., et al. 2015, *Sci*, 349, doi:10.1126/science.aab0689
- Hall, C., Forgan, D., & Rice, K. 2017, *MNRAS*, 470, 2517
- Hasegawa, T. I., Herbst, E., & Leung, C. M. 1992, *ApJS*, 82, 167
- Holdship, J., Viti, S., Jiménez-Serra, I., Makrymallis, A., & Priestley, F. 2017, *AJ*, 154, 38
- Ilee, J. D., Boley, A. C., Caselli, P., et al. 2011, *MNRAS*, 417, 2950
- Ilee, J. D., Forgan, D. H., Evans, M. G., et al. 2017, *MNRAS*, 472, 189
- Kahane, C., Ceccarelli, C., Faure, A., & Caux, E. 2013, *ApJL*, 763, L38
- Kuskov, O. L., & Kronrod, V. A. 2001, *Icar*, 151, 204
- Kuskov, O. L., & Kronrod, V. A. 2005, *SoSyR*, 39, 283
- Le Roy, L., Altwegg, K., Balsiger, H., et al. 2015, *A&A*, 583, A1
- McElroy, D., Walsh, C., Markwick, A. J., et al. 2013, *A&A*, 550, A36
- Minissale, M., Dulieu, F., Cazaux, S., & Hocuk, S. 2016, *A&A*, 585, A24
- Mosqueira, I., & Estrada, P. R. 2003, *Icar*, 163, 198
- Nayakshin, S. 2017, *PASA*, 34, e002
- Nayakshin, S. 2018, arXiv:1808.05151
- Öberg, K. I., Guzmán, V. V., Furuya, K., et al. 2015, *Natur*, 520, 198
- Quénard, D., Jiménez-Serra, I., Viti, S., Holdship, J., & Coutens, A. 2018, *MNRAS*, 474, 2796
- Rawlings, J. M. C., Hartquist, T. W., Menten, K. M., & Williams, D. A. 1992, *MNRAS*, 255, 471
- Saladino, R., Crestini, C., Pino, S., Costanzo, G., & Di Mauro, E. 2012, *PhLRv*, 9, 84
- Szulágyi, J., Mayer, L., & Quinn, T. 2017, *MNRAS*, 464, 3158
- Visser, R., Doty, S. D., & van Dishoeck, E. F. 2011, *A&A*, 534, A132
- Viti, S., Collings, M. P., Dever, J. W., McCoustra, M. R. S., & Williams, D. A. 2004, *MNRAS*, 354, 1141
- Wakelam, V., Loison, J.-C., Mereau, R., & Ruaud, M. 2017, *MolAs*, 6, 22
- Walsh, C., Loomis, R. A., Öberg, K. I., et al. 2016, *ApJL*, 823, L10
- Walsh, C., Millar, T. J., Nomura, H., et al. 2014, *A&A*, 563, A33
- Wang, H., Bell, R. C., Iedema, M. J., Tsekouras, A. A., & Cowin, J. P. 2005, *ApJ*, 620, 1027
- Yoneda, H., Tsukamoto, Y., Furuya, K., & Aikawa, Y. 2016, *ApJ*, 833, 105