

The VANDELS ESO public spectroscopic survey: observations and first data release

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ABSTRACT

This paper describes the observations and the first data release (DR1) of the ESO public spectroscopic survey "VANDELS, a deep VIMOS survey of the CANDELS CDFS and UDS fields".

VANDELS' main targets are star-forming galaxies at redshift $2.4 < z < 5.5$, an epoch when the Universe was less than 20% of its current age, and massive passive galaxies in the range $1 < z < 2.5$. By adopting a strategy of ultra-long exposure times, ranging from a minimum of 20 hours to a maximum of 80 hours per source, VANDELS is specifically designed to be the deepest ever spectroscopic survey of the high-redshift Universe. Exploiting the red sensitivity of the refurbished VIMOS spectrograph, the survey is obtaining ultra-deep optical spectroscopy covering the wavelength range 4800-10000 Å with sufficient signal-to-noise to investigate the astrophysics of high-redshift galaxy evolution via detailed absorption line studies of well defined samples of high-redshift galaxies.

The VANDELS-DR1 is the release of all medium-resolution spectroscopic data obtained during the first season of observations, on a 0.2 square degree area centered around the CANDELS-CDFS and CANDELS-UDS areas. It includes data for all galaxies for which the total (or half of the total) scheduled integration time was completed. The data release contains 879 individual objects, approximately half in each of the two fields, which have a measured redshift, with the highest reliable redshifts reaching $z_{spec} \sim 6$. In the data release we include fully wavelength and flux-calibrated 1D spectra, the associated error spectrum and sky spectrum and the associated wavelength-calibrated 2D spectra. We also provide a catalogue with the essential galaxy parameters, including spectroscopic redshifts and redshift quality flags measured by the collaboration. In this paper we present the survey layout and observations, the data reduction and redshift measurement procedure and the general properties of the VANDELS-DR1 sample. In particular we discuss the spectroscopic redshift distribution, the accuracy of the photometric redshifts for each individual target category and we provide some examples of data products for the various target types and the different quality flags.

All VANDELS-DR1 data are publicly available and can be retrieved from the ESO archive. Two further data releases are foreseen in the next two years with a final data release currently scheduled for June 2020 which will include improved re-reduction of the entire spectroscopic data set.

1. Introduction

Making significant progress in our understanding of galaxy formation and evolution requires observations of substantial samples of galaxies, over a large enough volume and a large range of mass and redshift. Only with representative samples and reliable observations we are able to test models of galaxy formation in a rigorous way. In particular, spectroscopic surveys can play a key role, since they provide, besides accurate information on the galaxies' redshift, a wealth of other important observable properties, such as emission and absorption line features, measures of internal motions, spectral indices and spectral breaks. These in turn allow us to characterise the intrinsic physical properties of galaxies, the nature of their stellar populations including the chemical composition, the presence of non thermal sources, the

ionizing radiation field and the evolution of all these properties with cosmic time (Ellis et al. 2017).

In recent years a series of spectroscopic campaigns have been carried out, starting from the low redshift Universe where the Sloan Digital Sky Survey (SDSS) observed more than a million galaxies at redshifts $z \sim 0.1-0.7$ (Reid et al. 2016; Alam et al. 2015). At redshifts $z \sim 1$ a series of spectroscopic surveys have sampled large volumes of the Universe, by observing many tens of thousands galaxies: in this redshift range, the Very Large Telescope (VLT) and the Visible MultiObject Spectrograph (VIMOS) have played a major role with the the VIMOS Very Deep Survey (VVDS Le Fèvre et al. 2005), the COSMOS spectroscopic survey (zCOSMOS Lilly et al. 2007) and the VIMOS Public Extragalactic Redshift Survey (VIPERS Guzzo et al. 2014). These surveys contributed substantially to improv-

ing our understanding of galaxy evolution also in relation to the environment. At even higher redshift ($z > 2$) spectroscopic surveys have been necessarily more limited with at most several thousand galaxies identified at $z \sim 2-4$ (e.g. KBSS-MOSFIRE, Steidel et al. 2014 and VUDS, Le Fèvre et al. 2015), which become less than few hundreds as we approach the reionization epoch at $z \geq 5$ (De Barros et al. 2017; Shibuya et al. 2018; Pentericci et al. 2018 subm). The main aim of these surveys has been the redshift identification of increasingly distant (and faint) galaxies, particularly star forming objects. A somewhat complementary approach was employed by the Galaxy Mass Assembly ultra-deep Spectroscopic Survey (GMASS, Kurk et al. 2013; Cimatti et al. 2008), which investigated the physical and evolutionary processes of galaxy mass assembly in the redshift range of $1.5 < z < 3$, by obtaining ultra deep optical spectra (up to 32 hours), which allowed a detailed spectral study of a small sample of passive galaxies at $z \sim 1.5$ and several tens of star forming galaxies up to $z \sim 3$.

VANDELS, a VIMOS survey of the CANDELS CDFS and UDS fields, is an ESO public spectroscopic survey designed to complement and extend the work of these previous campaigns by focusing on ultra-long exposures of a relatively small number of galaxies, pre-selected to lie at high redshift. VANDELS started observations in August 2015 and it was completed in February 2018. Exploiting the red sensitivity of the refurbished VIMOS spectrograph, and ultra-long integration times of up to 80 hours on source, the survey is obtaining ultra-deep optical spectroscopy of around 2100 galaxies in the redshift interval $1.0 < z < 7.0$, including star forming galaxies at redshift $z > 2.4$, an epoch when the Universe was less than 20% of its current age and massive passive galaxies in the range $1 < z < 2.5$. VANDELS has observed in the wavelength range 4800-10000 Å with intermediate resolution. VANDELS' prime motivation is to move beyond simple redshift acquisition, by obtaining spectra with high enough signal-to-noise to allow detailed absorption and emission line studies in individual spectra, derive metallicities and velocity offsets, and finally derive improved constraints on physical parameters, such as stellar mass and star formation rates. These informations will enable a detailed investigation of the physics of galaxies in the early Universe. By targeting two extragalactic survey fields with superb multi-wavelength imaging data, including the best optical+nearIR+Spitzer imaging, VANDELS will produce a unique legacy dataset for exploring the physics underpinning high-redshift galaxy evolution.

In this paper, after a brief summary of the target selection (Section 2), we present the survey layout and observations (Section 3), the data reduction process (Section 4), the redshift measurement procedure (Section 5) and, finally, the content of the first data release (DR1) (Section 6) with the general description of the data publicly available, the properties of the galaxies in the release and some examples of data products for different types of galaxies. In a companion paper (McLure et al. 2018), we present complementary information including the main scientific motivations of the survey, the assembly of the photometric catalogs, the determination of the photometric redshifts and the selection for the different category of targets that were observed in the survey.

We refer to total magnitudes in the AB system (Oke & Gunn 1983). When quoting absolute quantities we assume a cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_\Lambda = 0.7$, and $\Omega_m = 0.3$.

2. The VANDELS public spectroscopic survey

We provide a brief summary of the survey below, but the reader is referred to McLure et al. (2018) for a detailed description of the target selection and survey definition.

The VANDELS survey targets two fields, the UKIDSS Ultra Deep Survey (UDS: 02:17:38, -05:11:55) and the Chandra Deep Field South (CDFS: 03:32:30, -27:48:28). These fields were selected since their central areas have the best available HST multi-wavelength data from the CANDELS treasury survey (Koeke-moer et al. 2011; Grogin et al. 2011), as well as a wealth of ancillary data including ultra-deep IRAC photometry.

For the CANDELS/HST areas (CDFS-HST and UDS-HST) we adopted the photometric catalogues, based on H_{160} -band detections, provided by the CANDELS team (Galametz et al. 2013; Guo et al. 2013) to select our targets. These catalogues provide PSF homogenised photometry for the available ACS and WFC3/IR imaging, in addition to spatial-resolution matched photometry from Spitzer IRAC and key ground-based imaging data sets.

Because of the large field of view of VIMOS, the spectroscopic observations cover also areas that are outside the original CANDELS footprints: for these wider-field areas (CDFS-GROUND and UDS-GROUND) there were no near-IR selected, photometric catalogues therefore new multi-wavelength photometric catalogues were assembled using the publicly available imaging. The assembly of the new catalogues and the multiband data employed, is described in details in McLure et al. (2018).

For the two central regions covered by deep HST imaging, we adopted the photometric redshift solutions by the CANDELS survey team (Santini et al. 2015; Guo et al. 2013), which were derived by optimally combining several independent estimates produced by different photometric redshift codes by CANDELS team members, as described in details in Dahlen et al. (2013). For the wider areas outside the CANDELS footprint, new photometric redshifts were generated by our team, based on the new photometric catalogues described above. Fourteen independent photometric redshifts estimates were generated by eleven individual team members, based on different codes and methods. They were then combined taking the median value of z_{phot} for each galaxy and producing an official VANDELS z_{phot} .

Target selection was then carried out using the CANDELS and the new VANDELS photometric catalogs and photometric redshifts, respectively for the areas with deep HST imaging and for the wider areas. Here we briefly recall the main categories of objects that were selected:

- 1) bright star forming galaxies (SFG): this sample consists of star forming galaxies with $i \leq 25$ with a best photometric redshift in the range $2.4 \leq z_{\text{phot}} \leq 5.5$ so that the main absorption features necessary to investigate the metallicity (e.g. as in Sommariva et al. 2012 and Rix et al. 2004) fall in the observed spectral range. In practice the redshift range of the resulting sample observed is limited to $z_{\text{phot}} < 5$. The galaxies were required to have $sSFR > 0.1 \text{ Gyr}^{-1}$ where $sSFR$ is the specific star formation rate derived by the SED fitting described by McLure et al. (2018), assuming the best photometric redshift.
- 2) passive galaxies (PASS): this sample consists of UVJ selected passive galaxies (see Williams et al. 2009 and McLure et al. 2018 for more details on the definition) which have a photometric redshift in the range $1 \leq z_{\text{phot}} \leq 2.5$ with $H \leq 22.5$ and $i \leq 25$. The magnitude constraints is equivalent to a minimum total stellar mass of $10^{10} M_\odot$. In practise the redshift range of the resulting sample observed is limited to $z \sim 2$.
- 3) fainter star forming galaxies that we call Lyman Break Galax-

ies (LBG) to differentiate them from group 1: this sample consists of galaxies with a photometric redshift in the range $3 \leq z_{phot} \leq 5.5$ which have $H \leq 27$ and $i \leq 27.5$ ($i \leq 26$ in the wider regions) and galaxies with $5.5 \leq z_{phot} \leq 7$ which have $H \leq 27$ and $z \leq 26$ ($z \leq 26$ and $z \leq 25$ in the wider regions, respectively for the UDS and CDFS). The redshift range is such that the Ly α emission line or the Lyman break fall within the observed spectral range. As for the bright star forming galaxies, the targets were required to have $sSFR > 0.1 Gyr^{-1}$.

In the above cases the i-band (or z-band for objects with $z_{phot} > 5.5$) constraints are imposed in order to ensure that the final 1d spectra have a minimum S/N in the observed spectral range around 6000-7000Å, specifically S/N per resolution element > 15 for objects with $i < 24.5$ and S/N per resolution element ~ 10 for the faintest $i = 25$ objects. For the faint star-forming galaxies, the VANDELS strategy is designed to provide a consistent Ly α emission line detection limit of $\sim 2 \times 10^{-18}$ erg $s^{-1} cm^{-2}$ at 5σ .

In addition, a small sample of Herschel detected sources with $z_{phot} > 2.4$ and $i < 27.5$ were selected in both fields (1% of the sample), and in the CDFS we further added another 2% of targets selected as AGN, which were selected thanks to the availability of the Chandra 4Msec (Xue et al. 2011) observations as in Hsu et al. (2014), or via IRAC power-law and 24 μm detection as in Chang et al. (2017), with the further constraints of having $z_{phot} > 2.4$ and $i < 27.5$. For these, the photometric redshifts present in our catalog where used.

Following the above selection criteria a total of 9656, objects were available in the final catalogs, approximately half for the UDS and CDFS fields respectively.

3. Survey layout and observations

Our survey is conducted with VIMOS (Le Fèvre et al. 2003) which is mounted on the ESO-VLT unit number 3 “Melipal” and has a pixel scale of 0.205"/pixel with a total field of view of $4 \times 7' \times 8'$. We used the medium resolution grism which gives a wavelength coverage from 4800 to 10000 Å and provides a resolution of 580 with a dispersion of 2.5Å/pix. Spectral multiplexing (i.e. the ability of having more than one slit along the dispersion direction) is possible only when sources are positioned at the very upper or bottom edge of the FOV. As detailed in McLure et al. (2018), the VANDELS survey targets a total of eight VIMOS pointings, four pointings in the UDS field and four pointings in the CDFS field. Each VANDELS pointing has four associated masks, each of which is observed for 20-hours of on-source integration time. The survey utilizes a nested slit allocation policy, such that the brightest objects within a given pointing appear on a single mask (receiving 20-hours of integration), fainter objects appear on two masks (receiving 40-hours of integration) and the faintest objects appear on all four masks (receiving 80-hours of integration). In this way we can reach a very similar S/N on the continuum (for the passive and star forming galaxies) or emission line flux limit (for the faint star forming galaxies) for all our sources. In the next subsections we will describe the target allocation procedure and the procedure to make the VIMOS masks.

3.1. Slit allocation

Because of the nested strategy that VANDELS employs, all target allocation was done simultaneously for all masks at the beginning of the survey, in order to maximise the total number of

sources observed. Extensive simulations were run using the Slit Positioning Optimization Code (SPOC, Bottini et al. 2005). The primary goal of the simulation was to maximize the total number of total slits allocated to VANDELS targets, keeping particular attention to the two categories of bright star-forming galaxies and massive passive galaxies, which are the samples with the lowest surface densities. The only other constraint applied during the simulation work was the requirement to allocate the slits to objects requiring 20, 40 and 80 hours of integration in approximately a 1:2:1 ratio. We did not apply any additional prioritization (e.g. in terms of redshift or source brightness) during the slit allocation. Before the simulations it was decided to keep a minimum slit length of 7" (28 pixels) to facilitate sky background subtraction and that the targets should be positioned at least 8 pixels from one edge, to take into account the nodding strategy that is employed in the observations. Targets were treated as point sources.

Given the uneven distribution in ra-dec of our targets, due also to the fact that only the central area is covered by the deep CANDELS data (representing about 45% of the total area), the original goal stated in the proposal to observe a total of 1280 objects per field proved to be too ambitious, and a more realistic set of numbers was about 20% lower for all categories. We notice that the total area covered by the VIMOS pointings in the CDFS is 20% smaller than in the UDS, 360 sq arcmins instead of 460 sq arcmins in UDS, due to a different choice of pointings centers. The CDFS also has a surface density of PASS targets which is lower by $\sim 15\%$ compared to the UDS, due to cosmic variance. Finally, in the CDFS field we have the further set of AGN selected targets which is absent in the UDS, but this group comprises only 2% of the catalog.

In a first pass we only took the SFG+PASS+HERSCHEL targets as input (adding the AGN for the CDFS field) to check what is the maximum number of such targets we could observe. Then a series of runs were done using this sample as a forced set of targets and adding the list of LBG galaxies, trying to maximize the total number of targets observed, while avoiding to penalize too much the PASS+SFG targets. The best solution was obtained with a total of 1078 galaxies for the UDS field (1028 for CDFS), i.e. 693, 224, 151 and 10 respectively for LBG, SFG, PASS and Herschel for the UDS field (656, 200, 117, 9 and 46 for LBG, SFG, PASS, HERSCHEL and AGN for the CDFS field). A final further optimization was run using this set of 1079(1025) targets as input catalog to position slits but this time allowing completed targets to still be in the observable pool of objects for subsequent pointings: in this way several tens of targets get a higher S/N than initially allocated although we lose the completion of 1 target for UDS (gain 3 targets for CDFS). We finally remark that although the HST regions represent about 45% of the total survey area, the higher surface density of faint $z > 3$ galaxies allowed us to maximize the number of slits in these central HST areas. The HST selected galaxies are approximately 54% and 48% of the total targets respectively for the CDFS and UDS fields.

3.2. Mask preparation

Mask preparation was done using the VIMOS Mask Preparation Software (VMMPS, Bottini et al. 2005), distributed by ESO. VMMPS requires the acquisition of a direct image (pre-imaging), from which a list of visible sources is extracted. The pre-imaging was obtained in October 2014 in service mode, in the R band. The catalogue is then cross-matched with the catalogue of sources contributed by the user, so to match the astrono-

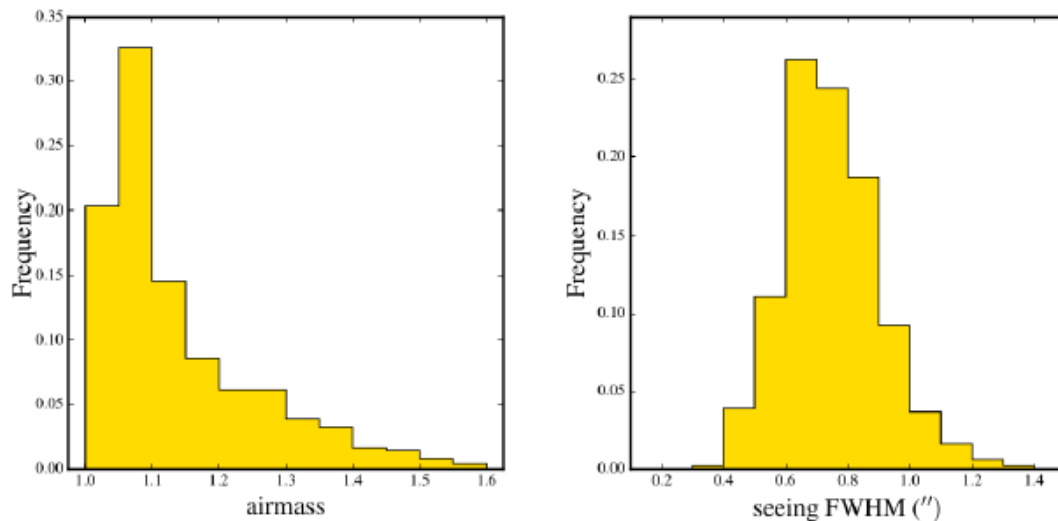


Fig. 1. The observing conditions under which the data released in DR1 were obtained: left is the airmass of individual exposures, mostly constrained below 1.5; right is the seeing, measured directly on the spectra, mostly constrained below 1''

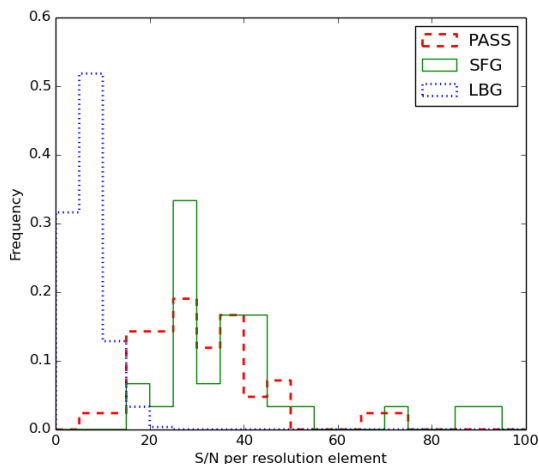


Fig. 2. The distribution of the median S/N per resolution element, of all completed DR1 spectra, measured in the range 6000-7400 Å and separated for the three main VANDELS classes. The red dashed histogram refers to PASS galaxies, the blue dotted histogram to LBGs and the green histogram to SFGs. Each histogram is normalised separately.

entry of the targets catalogue to the instrument coordinate system. After this, the cross-matched position of each VANDELS target was overplotted on the pre-image and visually inspected. If there was a mis-match in RA or Dec of one pixel or more, the target position was modified. This check has been done directly on the targets which were bright enough to be visible in the preimaging data, while for the faint targets the position of close-by brighter sources was checked with respect to their pre-image counterparts, and the possible mismatch used to modify the target position. The same procedure was applied to the two reference stars, chosen in each VIMOS quadrant.

As a second step, VMMPS assigns the slits to the input targets. We imposed a slit length of at least 28 pixels which means that the target center must always be at least 14 pixels away from both slit ends, while its maximum length is normally maximised

by VMMPS during the optimisation process, to allow for a better sky subtraction. Given the simulations described in Sect. 3.1, for each mask, we had an input catalogue of targets which were supposed to perfectly match the VIMOS multiplexing, given the slit constraints and the chosen set up. For this reason, it was expected that VMMPS could place a slit on all the input targets. This was not always the case for the following reasons: i) slightly different field of view with respect to the simulation, which in some cases made us miss the target closest to the left border of the CCD; ii) the initial astrometry match, which in crowded regions can modify the target position enough not to match anymore the constraints on slit length; iii) the presence of the two reference stars in each quadrant (not considered during the simulations), which reduced the area on the CCD available to slits. For all the targets which were not automatically assigned a slit by VMMPS, we verified if they could be fit in a slit by loosening the constraints on the position of the target within the slit, still keeping a minimum distance of the target from the slit borders of at least 10 pixels. This way we were able to fit in a slit almost all the input targets. On average, we lost 1-2 targets per quadrant, mainly because of the need for the reference stars.

3.3. Observations

All VANDELS observations were obtained in visitor mode during a period which ran from August 2015 to February 2018. Spectra belonging to the Data Release 1 were observed during the first observational season which ran from August 2015 to February 2016. The individual OBs were designed to deliver a total of one-hour of on-source integration time. Each OB consisted of three integrations of 1200s, obtained in a three-point dither pattern, with dither offsets of 0.82 arcseconds (dither positions 0,-0.82'',+0.82'') corresponding to 4 pixels. This was done to remove most of the small-scale detector pattern and facilitate sky-subtraction. For the same reason we have tried to obtain an equal number of frames at each of the three positions, even if on individual nights it was not always possible to complete the OBs.

One arc and one flat exposures were obtained for calibration after the execution of OB: it was possible to perform one cali-

bration every two OBs and save a little bit of time skipping the second calibration. In practise the observing sequence was set to be: science-calib-science. This was true for all the observations done while the target was rising/setting. If in one OB the target was rising and in the following one the target was setting we have obtained calibration at the end of both OBs. Finally, a spectrophotometric standard star was observed at least once every 7 nights and at least once per run during twilight, if possible under photometric conditions. Because of our wide wavelength coverage, we used bright late-type (F-G) stars such as LTT9239 and LTT3864 that cover up to $1\mu\text{m}$.

The nominal observing conditions required for a single exposure to be validated were: moon illumination ≤ 0.5 , moon angular distance above 90° , seeing $\leq 1.0''$ FWHM, as measured directly on the spectra of the brightest objects which were visible in the single exposures, airmass ≤ 1.5 and clear weather conditions. An exposure was still validated if one (and only one) of the above conditions was not met, but the discrepancy was less than 20% (e.g. seeing $\leq 1.2''$), while all other conditions were satisfied. The visiting observers judged directly from the quality of the spectra if the weather could be considered clear/photometric/thin (in some cases overriding the official ESO conditions) from the S/N of the brightest objects in the masks which were visible in 20 minute exposures.

In Fig. 1 we show histograms of airmass and FWHM for all individual exposures obtained during season one. The median seeing of the observations is $0.7''$ and only a tiny fraction of the data is obtained with seeing $\geq 1''$; similarly a negligible fraction of the data has been obtained at airmass >1.5 . In Figure 2 we show the distribution of signal-to-noise ratio (S/N) for the completed spectra released in DR1. The S/N is determined from the error spectrum as the median value in the range 6000-7400 Å, which is essentially free from bright skylines, per spectral resolution element (one spectral resolution element is equal to ~ 4 pixels). The distribution of S/N obtained agrees very well with the S/N predictions we made in the original proposal.

4. Data reduction

The VANDELS data was reduced using the fully-automated pipeline Easy-Life which starts from the raw data and produces wavelength-and flux- calibrated spectra. The pipeline is an updated version of the algorithms and dataflow from the original VIPGI system, fully described in Scodreggio et al. (2005). In Figure 3 we show a flow diagram illustrating the key features of the data reduction pipeline. The first step in the reduction of VIMOS science data is the canonical preliminary reduction of the CCD frames, which includes pre-scan level and average bias frame subtraction, trimming of the frame to eliminate pre-scan and over-scan areas, interpolation to remove bad CCD pixels and flat fielding. After the preliminary reduction step, subsequent data reduction steps are carried out on all MOS slits individually. For each individual spectroscopic exposure the wavelength calibration derived from the arc exposures is checked against the positions of bright sky lines and the local inverse dispersion solution modified to account for any discrepancies. The next steps in the data reduction procedure are the object detection and sky subtraction for each MOS slit spectrum within each individual 1200s science exposure. Initially, the slit spectrum is collapsed along the wavelength axis, following the geometrical shape defined by the local curvature model, to produce a slit cross-dispersion profile. A robust determination of the average signal level and r.m.s in this profile is then obtained using an iterative σ -clipping procedure, and objects are detected as groups

of contiguous pixels above a given detection threshold. Before wavelength calibration is applied, a median estimate of the sky spectrum, derived using all the pixels that are devoid of object signal, is subtracted from each slit. The sky spectrum is estimated separately for each individual science exposure due to the significant variation in OH line strength over the timescale of a typical spectroscopic exposure. The sky- subtracted slit spectra are then two-dimensionally extracted using the tracing provided by the slit curvature model, and resampled to a common linear wavelength scale. Only after this point are the single exposures of a pointing combined together. First, the N two-dimensionally extracted spectra for each slit are median combined (with object pixels masked), without taking into account the jitter off-sets, to produce a two-dimensional sky-subtraction residual map. The residual map is then subtracted from all the N two-dimensional single-exposure slit spectra, improving the sky-subtraction and removing any residual fringing. At this point a second combination is carried out, this time taking into account the jitter offsets among the N individual two-dimensional slit spectra (as determined during the previous object detection procedure).

The single-exposure, residual-map subtracted, spectra are off-set to compensate for the effect of the jitter, and a final average two-dimensional spectrum for each slit is obtained. The object detection process is repeated on the combined two-dimensional spectra to produce the final catalogue of detected spectra, and a one-dimensional spectrum is extracted for each detected object, using the Horne optimal extraction procedure Horne (1986). Finally, spectra are flux calibrated using a simple polynomial fit to the instrument response curve derived from observations of spectrophotometric standard stars, and corrected for telluric absorption features. The last correction is based on a template absorption spectrum derived for each combined jitter sequence from the data themselves. The final flux calibration was performed correcting the spectra for both atmospheric and galactic extinction and then normalizing them to the i-band photometry available for each target. This procedure was already successfully employed by the VIMOS Ultra Deep Survey (Le Fèvre et al. 2015). We plan to improve the calibration procedure employing additional broad band filters: the final data release will feature a re-reduction of the entire spectroscopic data set, incorporating this and possibly other improvements.

4.1. Problems with blue-end calibration

During final testing of the flux calibration of the DR1 spectra, it became clear that the extreme blue end of the spectra (i.e. $\lambda \leq 5600\text{Å}$) suffer from a systematic drop in flux when compared to the available broad band photometry. The underlying cause for this loss of blue flux is still under investigation. For the purposes of the first data release we have implemented an empirically derived correction to the spectra at these blue wavelengths which accounts for the flux loss on average. The empirical correction, which has been applied to all of the DR1 spectra, is designed to ensure that the final spectra of bright star-forming galaxies in the redshift interval $2.4 < z < 3.0$ display the expected power-law continuum slopes in the rest-frame wavelength range ($1300\text{Å} < \lambda < 2400\text{Å}$), which are independently confirmed from the available photometry. At the time of the data release, we believe that the spectra including the correction for blue flux loss represent our best calibration of the VANDELS spectra. However, for completeness, we are also making available the spectra without the blue flux correction. To show the effect of this blue correction, in Figure 4 we present two exam-

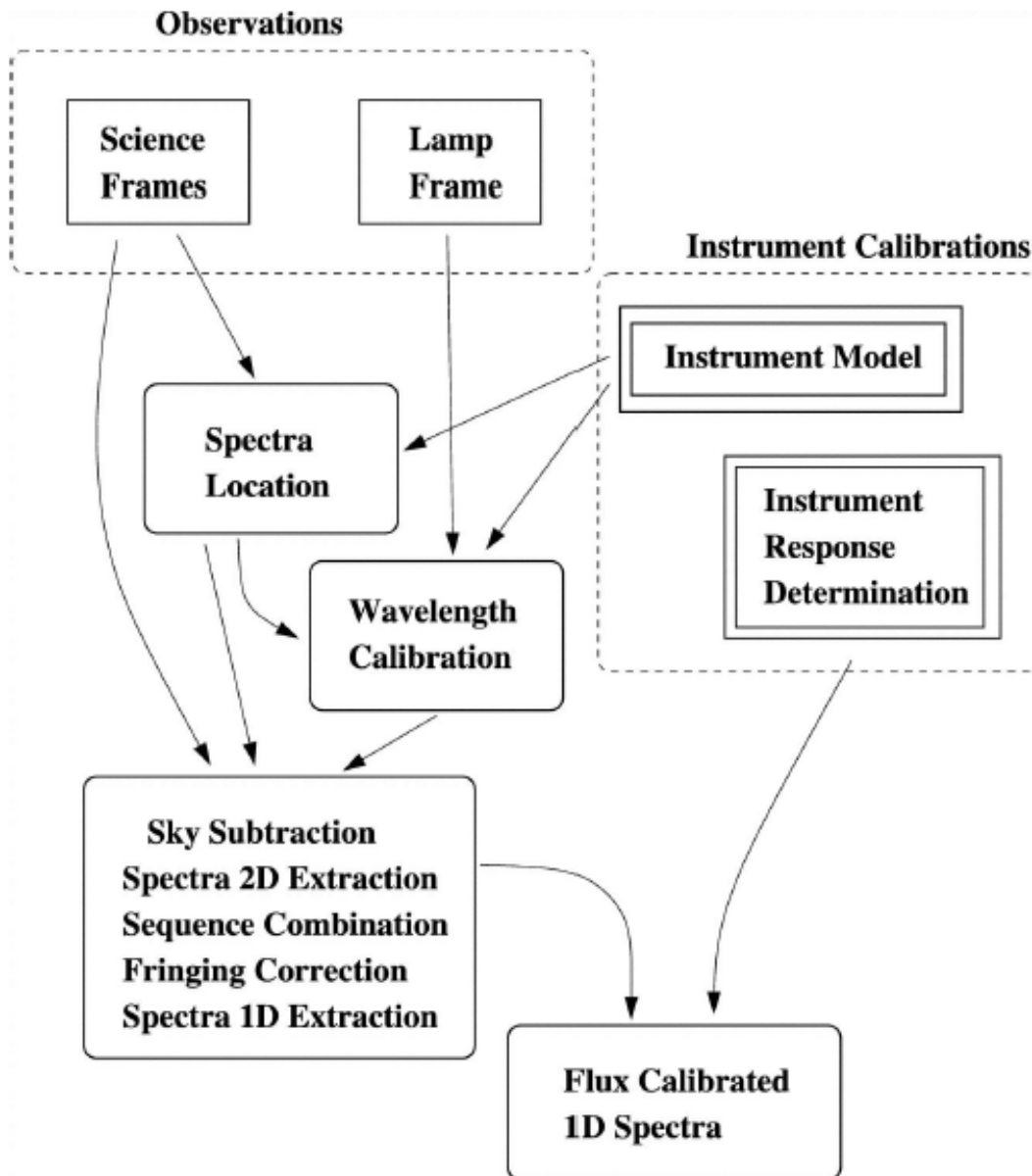


Fig. 3. A flow diagram illustrating the various steps of the data reduction pipeline

ples of spectra belonging to the SFG and PASS samples, respectively, and we compare the original and corrected flux. The effect is only noticeable in the first few hundreds Å of the spectral range, and it is most important for the star forming galaxies.

5. Redshift measurement

Spectroscopic redshifts and associated quality flags have been determined for all objects using the Pandora software package, within the EZ environment (Garilli et al. 2010). This software was used successfully in previous VIMOS survey and can simultaneously display the 1D extracted spectrum, the 2D linearly re-sampled spectrum, the 1D sky spectrum and the noise. It is also possible to inspect the image thumbnail of the object with the exact position of the slit, to check for the presence of other sources in the same slit and the pixels over which the source was extracted. The core algorithm for redshift determination is the cor-

relation with available galaxy spectral templates. A key element for the cross-correlation engine to deliver a robust measurement is the availability of reference templates that cover a wide range of galaxy and star types and a wide range of rest-frame wavelengths. To determine the VANDELS spectroscopic redshifts we adopted templates derived from previous VIMOS observations for the VVDS (Le Fèvre et al. 2013) and zCOSMOS surveys (Lilly et al. 2007), with and without Ly α emission. Alternatively, it is also possible to determine the redshift by estimating manually the center of one or more emission and absorption lines. In several cases, it was necessary to manually perform some cleaning of the spectra, i.e. removing obvious noise residuals at the location of strong sky lines, or the zero-order projection.

Each target was assigned to two measurers from the VANDELS team who independently determined the redshift and located the main spectral features (in emission or absorption). Each measurer also assigned a spectroscopic quality flag to the

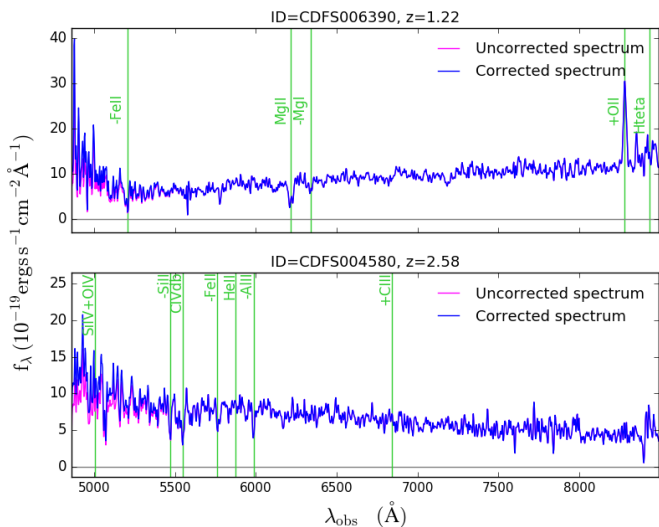


Fig. 4. The effect of the blue flux correction: in blue the final corrected spectrum and in magenta the original flux calibrated spectrum for CDFS006390 at $z=1.22$ (top panel) which was selected in the PASS sample, and CDFS004580 at $z=2.58$ which was selected in the SFG sample. We also indicate with green vertical lines the main spectral features identified.

target: these quality flags have been allocated according to the original system devised for the VVDS and are related to the confidence of the spectral measurement. The reliability flag may take the following values:

0: No redshift could be assigned (redshift is set to Nan)

1: 50% probability to be correct. Some low S/N lines are identified but there is a weak to moderate match with templates.

2: 70-80% probability to be correct. There are several matching absorption lines and a general good match with templates.

3: 95-100% probability to be correct. The spectrum shows multiple strong absorption and/or emission lines giving a consistent redshift and has a moderate S/N. There is a strong cross-correlation signal with an excellent continuum match to templates.

4: 100% probability to be correct. The spectrum shows multiple strong absorption and/or emission lines giving a consistent redshift, and has a high S/N. There is a strong cross-correlation signal with an excellent continuum match to templates

9: spectrum with a single emission line. The redshift given is the most probable given the observed continuum and the shape of the emission line, and it has a $>80\%$ probability to be correct.

We emphasize that the quality flag only reflects the accuracy of the redshift determination and in principle is not related to the S/N of the spectrum although almost all the QF=4 spectra have very high S/N. The quality flags for AGN spectra are preceded by a leading 1 (e.g. 12, 14 etc), the quality flags for spectra which were not primary targets (i.e. were serendipitously observed in a certain slit) are preceded by an additional 2; finally the quality flags for spectra deemed to be problematic are preceded by an additional -1. In the DR1 only the AGN flags are present, since the secondary targets will be released only at the end of the survey. Following their independent redshift determinations, the two measurers were required to compare their redshifts and flags and to reconcile any differences. As a final step, all spectra were also independently re-checked by the two PIs and any remaining discrepancies in the redshifts and quality flags were again reconciled. This final pass was especially necessary to homogenize

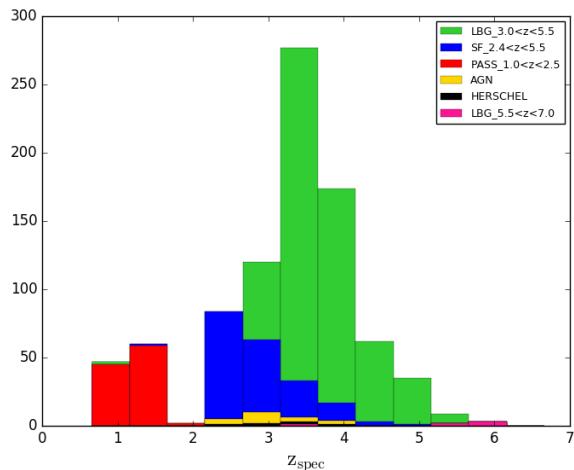


Fig. 5. The spectroscopic redshift distribution of all DR1 targets, divided by the original target classification

the quality flags as much as possible, given the different expertise and ability of the various redshift measurers. Based on repeated measurements, the typical accuracy of the spectroscopic redshift measurements is estimated to be ± 0.0005 ($\sim 150 \text{ km s}^{-1}$).

There are 42 objects in the Data Release 1 with a previously published spectroscopic redshift. These were included in the survey since the expected integration time (mostly 80 or 40 hours) would result in spectra with a much higher S/N compared to those already available. Of these 42 sources, one has no redshift assigned in VANDELS (but it is a source with half of the final scheduled integration) and one (AGN type) has a discrepant redshift, 3.442 with flag 1 in VANDELS and 2.448 from (Trump et al. 2013), also with a low quality flag indicating that the redshift was based on the match with the photometric redshift. For the other 40 objects, including all quality flags both for VANDELS redshifts and for the previous redshifts, $\Delta z = (z_{\text{VANDELS}} - z_{\text{old}})/(1 + z_{\text{VANDELS}})$ has a mean value of -0.0003 and an rms of 0.0016 , slightly higher than the redshift uncertainty quoted above.

6. VANDELS DR1-data

The data release 1 (DR1) is available in the ESO archive and consists of all spectra obtained during the first VANDELS observing season, which ran from August 2015 until February 2016. The data were acquired during runs 194.A-2003(E-K). The data release includes the spectra for all galaxies for which the final scheduled integration time was completed during season one (356 objects including 6 that actually received more than the scheduled time). In addition, the data release also includes the spectra for 523 galaxies for which the scheduled integration time was 50% complete at the end of season one (i.e. 20 out of 40 scheduled hours and 40 out of 80 scheduled hours). The total number of spectra released is 879 (415 in CDFS and 464 in UDS).

For each target the following data files are being released:

1) the one-dimensional spectrum in FITS format, containing the following arrays

WAVE: wavelength in Angstroms (in air)

FLUX: 1D spectrum blue-corrected flux in $\text{erg cm}^{-2} \text{s}^{-1} \text{angstrom}^{-1}$

ERR: noise estimate in $\text{erg cm}^{-2} \text{s}^{-1} \text{angstrom}^{-1}$

Table 1. Detailed list of entries in the catalogues associated to the data release 1

Name	Description	Data type	Unit
ID	The object ID	Char	
alpha	RA (J2000)	Double	deg
delta	Dec (J2000)	Double	deg
i_{AB}	The i-band magnitude	float	mag
$i - FILTER$	The i-band filter	string	
z_{AB}	The z-band magnitude	float	mag
$z - FILTER$	The z-band filter	string	
H_{AB}	The H-band magnitude	float	mag
$H - FILTER$	The H-band filter	string	
$t_{schedtime}$	The scheduled integration time	int	s
z_{phot}	The photometric redshift	float	
z_{spec}	Spectroscopic redshift	float	
z_{flag}	The quality flag for the redshift	float	
FILENAME	The fits filename of the spectrum	string	
$t_{exptime}$	The current integration time	float	s

UNCORR-FLUX: 1D spectrum flux uncorrected for blue flux loss (see details in subsection 4.1)

SKY: the subtracted sky in counts

2) the two-dimensional resampled and sky subtracted (but not flux calibrated and atmospheric extinction corrected) spectrum in FITS format.

3) a catalogue with essential galaxy parameters, listed in Table 1 and presented in Table 2 which include i_{AB} , z_{AB} and H_{AB} magnitudes with the relevant filters (see below), the scheduled and current integration times, and the spectroscopic redshift and quality flag determined as in Section 5. As already mentioned in Section 2, the UDS and CDFS fields are covered by different sets of observations, with 45% of the area covered by deep HST imaging (the CANDELS footprint) and the rest covered only by ground based imaging (the wider surrounding areas). For this reason the i_{AB} , z_{AB} and H_{AB} magnitudes listed in the release catalogues are generated from different filters. For each object, the origin of the i_{AB} , z_{AB} and H_{AB} photometry is listed in the catalogue in the columns with the filter name. Here we list the match between the catalogue photometry and the filters used:

- i_{AB} magnitudes refer to the SUBARU i' -filter for the UDS-GROUND and UDS-HST targets, to the F775W filter for CDFS-HST targets and to the SUBARU IA738 filter for the CDFS-GROUND targets.
- z_{AB} magnitudes refer to the SUBARU z' -filter for the UDS-GROUND and UDS-HST targets and to the F850LP filter for the CDFS-GROUND and CDFS-HST targets
- H_{AB} magnitudes refer to the F160W filter for the UDS-HST and CDFS-HST targets, the WFCAM H-filter for the UDS-GROUND targets and the VISTA H-filter for the CDFS-GROUND targets

In the catalogue we report both the total requested integration time ($t_{schedtime}$) and the current total integration time ($t_{exptime}$). Clearly the objects for which the observations are not completed will have $t_{exptime} < t_{schedtime}$. In some cases objects have $t_{exptime} > t_{schedtime}$: this is because, due to changing observing conditions, some masks in both the UDS and CDFS fields have received slightly more than their nominally scheduled 20 hours of on-source integration. In addition, to optimize slit allocation, six objects in this data release requiring 20 hours of on-source integration were placed on two VIMOS masks, and therefore actually received 40 hours of on-source integration.

7. Galaxies in the Data Release 1

7.1. General properties and redshift distribution

The targets in the DR1 belong to the following category: 177 objects are selected as SFG with $2.4 < z_{phot} < 5.5$, 106 objects are PASS with $1.0 < z_{phot} < 2.5$, 566 objects are LBG with $3.0 < z_{phot} < 5.5$ and 6 are LBG with $5.5 < z_{phot} < 7.0$. In addition we have 18 targets selected as radio or X-ray AGN (all in the CDFS field) and 6 Herschel selected sources (3 in each field). These numbers, with the subdivision into fields are reported in Table 3.

In Figure 5 we present the redshift distribution of all DR1 spectra, separated by the original target classification (i.e. the AGN class includes the targets originally selected as AGN and not the flag 14 objects). The quality flag statistics for galaxies in DR1 are reported in Table 4 and are as follows: 299 galaxies i.e. 34% of the sample, received a flag 4 or 14; 184 i.e. 21%, received a flag 3; 156 i.e., 18%, received a flag 2; 180 i.e. 20% received a flag 1; and 54 i.e. 6%, received a flag 9. For 6 objects no redshift was assigned and therefore they are labelled as flag 0. From the same Table we also see that for the two main classes of targets (SFG and PASS) the large majority of galaxies has flag ≥ 3 , while AGN and Herschel selected sources have all lower quality flags. 4 targets have received a flag 14 indicating that their spectra show AGN features: of these only one was originally selected as AGN, while two were selected as LBGs and one as SFG. In the other selected AGN, no prominent AGN features were identified even when the redshift was determined.

7.2. Photometric redshift accuracy

In Figures 6 and 7 we show the plots of the photometric redshifts vs spectroscopic redshifts for all galaxies that are part of the DR1 set, for which a redshift could be determined (873 objects), separated by flag type and target category. We also show in Figure 9 the same plot for objects that received the full final integration time (20, 40 or 80 hours) and those that received only half of the final allocation. In Table 5 we report the number of outliers and the resulting scatter between the photometric redshifts and the spectroscopic redshifts, where the photometric redshift is the CANDELS one for the CANDELS footprint areas, and the VANDELS one (McLure et al. 2018) for the outer areas. This Table presents both the full scatter $\sigma_F = rms(\Delta z)$ where $\Delta z = (z_{spec} - z_{phot}) / (1 + z_{spec})$, and σ_O derived after exclud-

ID	alpha	delta	i_{AB}	i_FILTER	z_{AB}	z_FILTER	H_{AB}	H_FILTER	$t_{schedule}$	z_{phot}	z_{spec}	z_{flg}	FILENAME	$t_{exptime}$
VANDELS_UDS_000090	34.4182011	-5.2764185	25.41889	SUBARU_i'	25.59937	SUBARU_z'	24.97066	HST_F160W	144000	3.35	3.239	3.0	VANDELS_UDS_000090.fits	71999.88
VANDELS_UDS_000129	34.4883644	-5.2760994	25.73422	SUBARU_i'	25.85772	SUBARU_z'	25.77788	HST_F160W	144000	3.44	3.639	4.0	VANDELS_UDS_000129.fits	72599.72
VANDELS_UDS_021074	34.383926	-5.1597713	25.45468	SUBARU_i'	25.53815	SUBARU_z'	24.77221	HST_F160W	144000	3.01	3.3044	1.0	VANDELS_UDS_021074.fits	72600.95
VANDELS_UDS_389704	34.4310368	-5.08566940	24.65722	SUBARU_i'	24.57985	SUBARU_z'	24.17536	WFCAM_LH	288000	3.58	3.5899	4.0	VANDELS_UDS_389704.fits	143999.7
VANDELS_CDFS_000541	53.1066882	-27.93019	26.26715	HST_F775W	24.67075	HST_F850LP	24.71833	HST_F160W	72000	5.764	5.784	4.0	VANDELS_CDFS_000541.fits	78001.62
VANDELS_CDFS_021735	53.0628228	-27.7264618	23.77141	HST_F775W	23.00993	HST_F850LP	21.08368	HST_F160W	144000	1.526	1.6087	2.0	VANDELS_CDFS_021735.fits	72001.39
VANDELS_CDFS_246958	53.2954379	-27.681040	25.44884	SUBARU_IA738	25.17301	HST_F850LP	26.77869	VISTA_H	144000	3.2	3.428	9.0	VANDELS_CDFS_246958.fits	149999.6
VANDELS_CDFS_129520	53.2628908	-27.7349146	24.60112	SUBARU_IA738	22.93470	HST_F850LP	20.77475	VISTA_H	288000	1.59	1.6074	3.0	VANDELS_CDFS_129520.fits	149999.6

Table 2. Data release 1 catalogue: random examples of galaxies in the catalogue of the first data release. the entries are explained in Table 1. The data release is available at <http://www.eso.org/sci/publications/announcements/sciann17068.html> and the catalogue is available at <https://www.eso.org/qi/catalogQuery/index/212>

Table 3. Galaxies in DR1 divided by target type

Sample	N_{UDS}	N_{CDFs}	N_{tot}
SFG $2.4 < z_{phot} < 5.5$	110	67	177
LBG $3.0 < z_{phot} < 5.5$	285	281	566
PASS $1.0 < z_{phot} < 2.5$	64	42	106
LBG $5.5 < z_{phot} < 7.0$	2	4	6
AGN	0	18	18
HERSCHEL	3	3	6
TOTAL	464	415	879

Table 4. Galaxies in DR1 divided by flag

Flag	SFG	LBG	PASS	AGN	HERSCH	Total
Flag 0	0	6	0	0	0	6
Flag 1	10	152	3	10	5	180
Flag 2	19	127	5	5	0	156
Flag 3	29	137	18	0	0	184
Flag 4(14)	119	99	80	1	0	299
Flag 9	0	51	0	2	1	54
TOTAL	177	572	106	18	6	879

ing the catastrophic outliers. We define a catastrophic outlier as a galaxy for which $|\Delta z| > 0.15$. We remark that σ_o gives a non-optimal representation of the scatter since a few objects (i.e., the outliers) can drive the scatter to large values. We also quantify any systematic bias between photometric and spectroscopic redshifts, by $b = \text{mean}(\Delta z)$ after excluding the outliers. From the Table we see that the bias is much smaller than the scatter.

In total there are 18 outliers, thus the resulting rate is 2.1% which is in perfect agreement with the rate estimated from the accuracy of the VANDELS photometric redshifts, even if these were actually validated on a sample with brighter average magnitude (see McLure et al. 2018 for details). The number of outliers is higher for Flag 1 and 9 objects as expected, while it is extremely low for Flags 2,3,4. This indicates that our flag system is actually conservative and probably the flags underestimate the reliability of the redshift. We notice that 8 of the 18 outliers are actually AGN or HERSCHEL selected sources. In particular, for AGN the disagreement between photometric and spectroscopic redshift might be partly due to the fact that only templates of normal galaxies were used to determine the photometric redshifts (Salvato et al. 2011, 2009).

If we further restrict the sample to the three main categories of targets of the VANDELS proposal, the outliers rate is actually just 1.2% which is extremely low. The Table shows that the agreement between z_{phot} and z_{spec} is excellent, for all quality flags with σ_o clearly improving for higher flags. The class of targets with the highest overall accuracy is the PASS one: this is easily explained both by the relative brightness of this sample (the median H-band magnitude is $H = 21.4$) and by the fact that the 4000 Å break for galaxies in this redshift range is very clear in the SED. In the same Table we report the accuracy of the photometric redshifts divided by H-magnitude bins, after removing the HERSCHEL and AGN targets to have an unbiased view of the three main VANDELS groups. We find that σ_o decreases mildly as a function of the H-band magnitude, while the percentage of outliers increasing slightly at fainter magnitudes. This was also found by Dahlen et al. (2013) for the CANDELS photometric redshifts, although in that case the fraction of outliers at mag-

nitude fainter than 25 was much higher, $> 10\%$ compared to our 2.3% even in the faintest bin. We also investigate the photometric redshift accuracy as a function of redshift, again after taking out the HERSCHEL and AGN targets. Figure 8 indicates that there is a redshift trend in the photometric redshift accuracy, since the scatter σ_o increases very smoothly from $\sigma_o = 0.016$ at $z \sim 1 - 1.3$ to $\sigma_o = 0.044$ at $z \sim 3.5 - 4$ and then decreases beyond that redshift.

We finally compare the accuracy of photometric redshifts in the CANDELS footprint and in the outer areas. The largest number of outliers is actually found in the CANDELS areas (16) but this is due mostly to the fact that all the Herschel targets and almost all of the AGN (i.e., the classes with the majority of interlopers) are selected in the CANDELS areas, and partly to the fact that the CANDELS targets are on average 0.4 magnitudes fainter than those in the wider areas. However the new VANDELS redshifts (i.e. those in the outer areas) also have a smaller bias and scatter. A possible explanation for the increased success rate of the VANDELS photometric redshifts is that all potential targets were visually inspected before the final assembly of the catalogue, to reject obviously spurious and artefacts.

7.3. Examples of data products

To illustrate the content of the VANDELS-DR1, we present some examples of data products for different types of targets and different quality flag. In Figures 10,11,12,13 we show the 2-dimensional, 1-dimensional, sky and noise spectra of four targets, respectively selected as PASS, LBG, SFG and AGN. All these galaxies have quality flags 4, since the redshift is determined with great accuracy from the many emission and/or absorption lines visible in the spectra. To illustrate the meaning of the different quality flags, in Figure 14 we present the spectra of five galaxies, all selected as LBGs, and approximately at the same redshift ($z \sim 4$), but with quality flags 4, 3, 2, 1 and 9 respectively. In the Flag 4 spectrum we can clearly identify two emission lines (the Ly α and HeII lines) and several absorption lines such as the SiII, OI, CII, CIV etc. The drop in the continuum flux below the Ly α line is also very clear. In the Flag 3 spectrum we identify the Ly α , several absorption lines (e.g. the SiII and CIV) and the drop in the continuum flux. In the Flag 2 spectrum (which has been smoothed in the Figure for clarity), we identify the SiII, CII and SiIV lines with good confidence and the drop in the continuum flux is also clear. In the Flag 1 only the SiIV line is clearly identified, but there is a general good agreement when cross correlating this spectrum with templates at the assigned redshift. Finally in the Flag 9 case, there is only one bright emission line that is detected in the spectrum. The line does not show a prominent asymmetry (which would clearly identify it as Ly α) and no continuum is detected blueward of the line, therefore there remains some ambiguity in the redshift identification.

8. Summary

We have presented the first Data Release of the VANDELS public spectroscopic survey. VANDELS, a deep survey of the CANDELS CDFS and UDS fields is an ESO public survey carried out with VIMOS and has obtained more than 2000 ultradeep medium resolution spectra of galaxies in the wavelength range 4800-10000 Å. The DR1 is the release of all spectra obtained during the first season of observations, and it includes all targets for which either the scheduled integration time or half of the

Table 5. Accuracy of photometric redshifts

Sample	$bias_z$	σ_F	σ_O	%out	N_{tot}	N_{out}
All Flags	-0.002	0.106	0.036	2.1	873	18
Flag 1	0.005	0.061	0.047	4.4	180	8
Flag 2	-0.004	0.054	0.035	0.6	156	1
Flag 3	0.001	0.184	0.031	1.6	184	3
Flag 4	0.005	0.039	0.029	0.3	295	1
Flag 14	0.039	0.021	0.021	0.0	4	0
Flag 9	0.016	0.187	0.039	9.3	54	5
SFG $2.4 < z_{phot} < 5.5$	-0.008	0.044	0.032	0.6	177	1
PASS $1.0 < z_{phot} < 2.5$	0.011	0.053	0.019	0.9	106	1
LBG $3.0 < z_{phot} < 5.5$	0.004	0.123	0.038	1.3	560	7
LBG $5.5 < z_{phot} < 7.0$	0.007	0.198	0.015	16.7	6	1
AGN	0.029	0.086	0.030	27.8	18	5
HERSCHEL	0.059	0.116	0.069	50.0	6	3
CANDELS footprint	0.002	0.132	0.040	3.6	439	16
Wide areas	0.004	0.070	0.032	0.5	434	2
H<21	0.008	0.018	0.018	0	41	0
H=[21-22]	0.012	0.074	0.019	2.0	50	1
H=[22-23]	-0.002	0.033	0.033	0.0	45	0
H=[23-24]	-0.005	0.123	0.037	1.6	129	2
H=[24-25]	0.001	0.165	0.038	0.9	222	2
H=[25-26]	0.001	0.047	0.035	1.2	254	3
H>26	0.014	0.073	0.042	2.3	87	2

total time was completed. The release includes spectra for 879 objects, 464 in the UDS and 415 in the CDFS. Together with the spectra, we release also the spectroscopic redshifts measured by the collaboration, with a quality flag that assesses their reliability. In this paper we present the statistics of the redshift quality and discuss the excellent accuracy of the VANDELS photometric redshifts, with an outlier rate as low as 2.1% and overall accuracy of $\sigma_o = 0.035$, which improves if we restrict the statistics to the main VANDELS target categories. We also present some examples of data products, to illustrate the content of the release. All spectra and information are available in the ESO archive. The second Data Release will include all spectra completed during the second observational season, and will be available at the end of Summer 2018. A final release is expected for June 2020 and will include improved re-reduction of the entire spectroscopic data, a series of galaxies' physical properties (stellar masses, SFRs, dust attenuation etc) derived by the collaboration through SED fitting, and measurements of absorption and emission lines identified in the VANDELS spectra.

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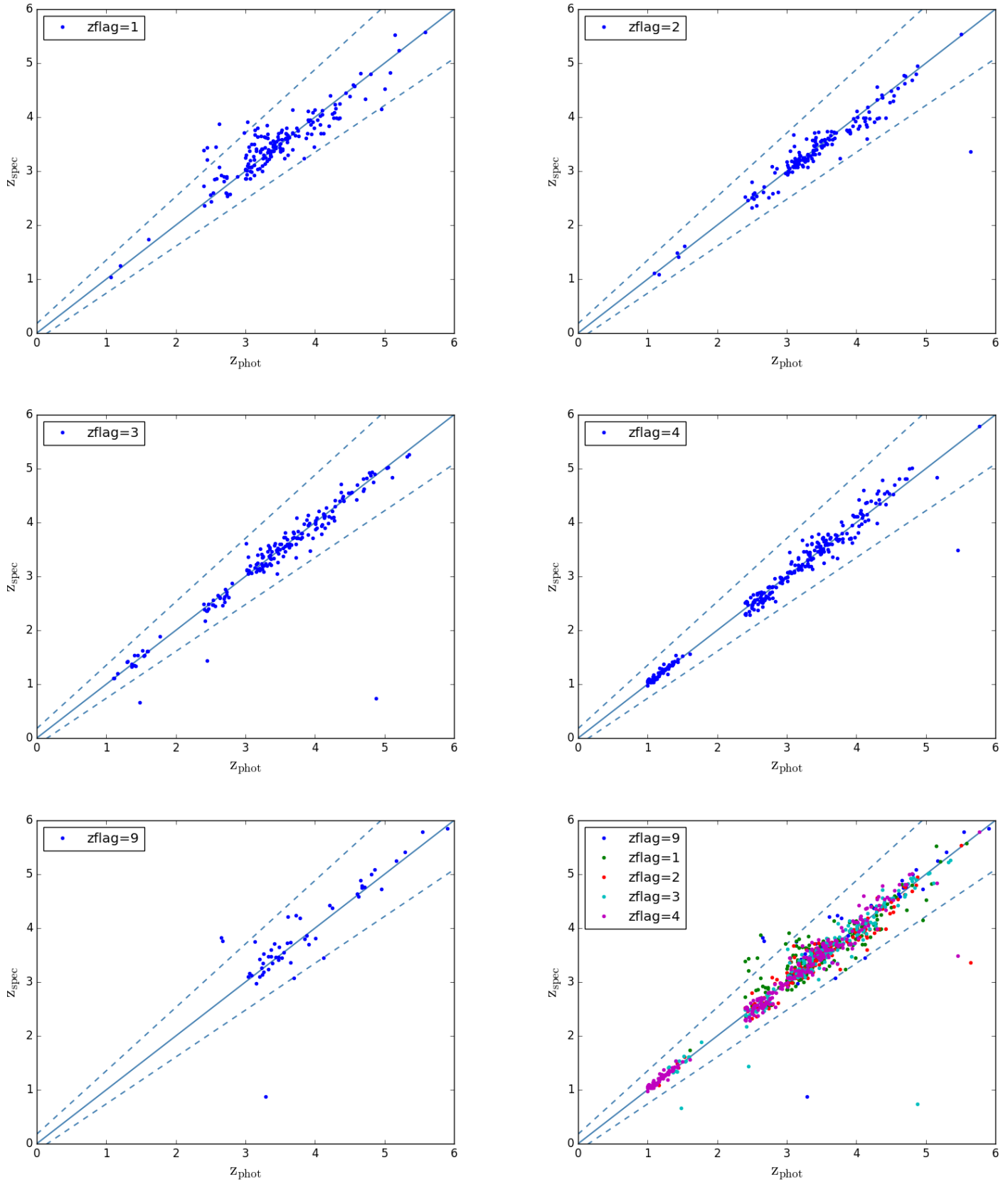


Fig. 6. The z_{phot} vs z_{spec} for all galaxies in the DR1 separated by flag type. The galaxies falling outside the dashed lines are the catastrophic outliers that have $|dz| > 0.15$.

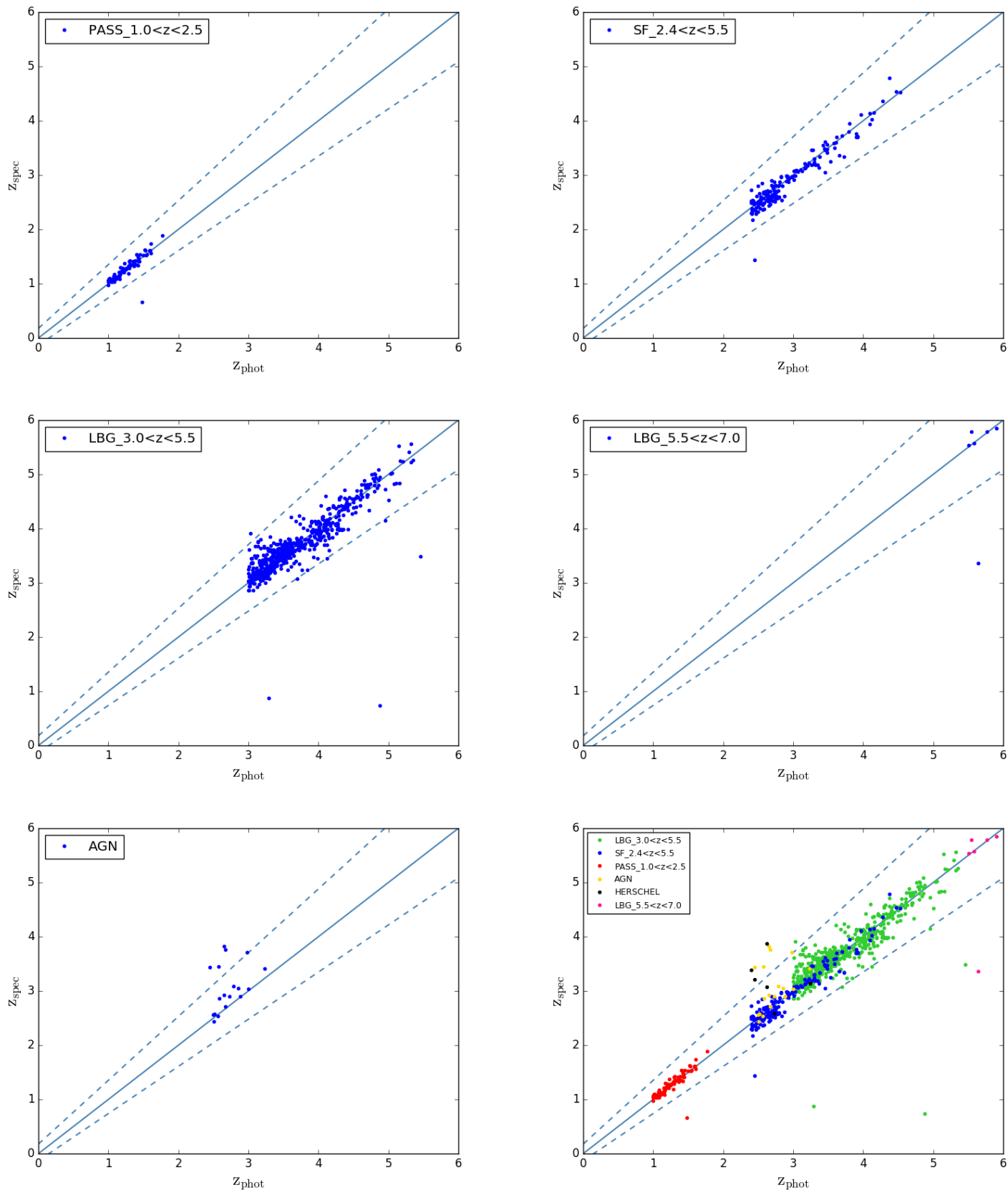


Fig. 7. The z_{phot} vs z_{spec} for all galaxies separated by selection class type. The galaxies falling outside the dashed lines are the catastrophic outliers that have $|dz| > 0.15$.

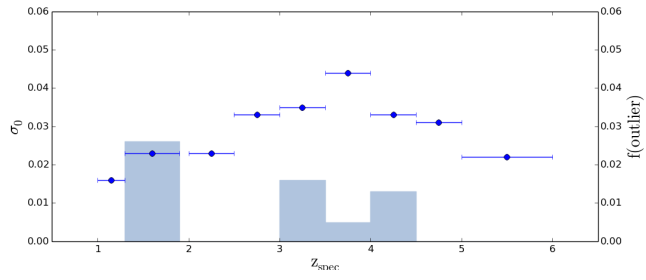


Fig. 8. Redshift dependence of the photometric redshift scatter and outlier fraction when comparing the photometric redshift with the spectroscopic redshift. The blue dots show the scatter σ_O (scaling on left-hand y-axis). The histograms show the fraction of outliers (scaling on right-hand y-axis). The sample does not include HERSCHEL and AGN selected targets.

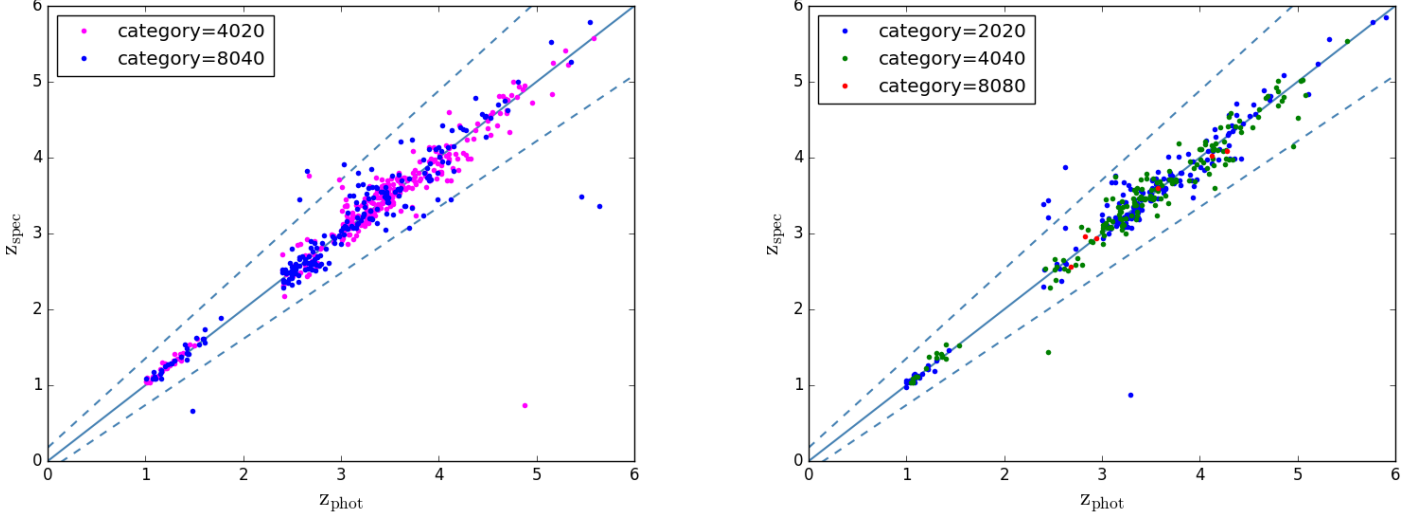


Fig. 9. The z_{phot} vs z_{spec} for all galaxies color-coded by integration time: left are the galaxies which received half of the final integration time, and right the galaxies which received the full (20, 40 or 80 hours) integration time. The galaxies falling outside the dashed lines are the catastrophic outliers that have $|dz| > 0.15$.

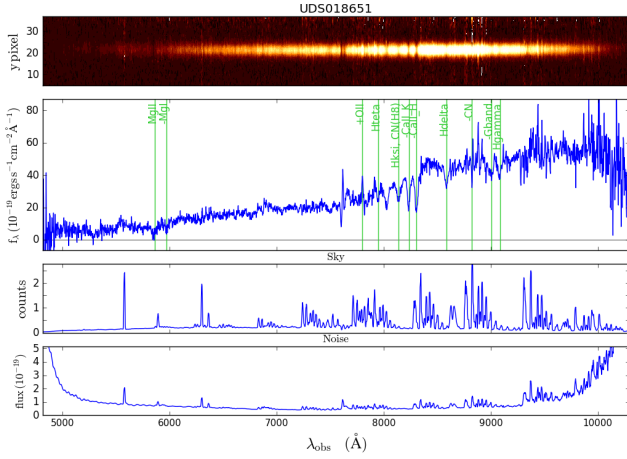


Fig. 10. From top to bottom, the 2-dimensional spectrum, 1-dimensional extracted spectrum, sky counts and rms noise of target VANDELS-UDS-018651. This object was selected as a PASS galaxy and has a redshift 1.093 with Flag 4. The main spectral features identified are indicated in the 1-D spectrum.

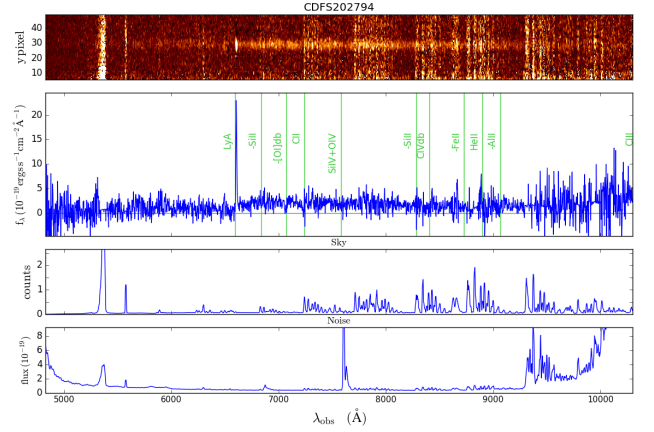


Fig. 11. From top to bottom, the 2-dimensional spectrum, 1-dimensional extracted spectrum, sky counts and rms noise of target VANDELS-CDFS-202794. This object has an H-band magnitude of 24.1 and was selected in the category LBG. It was assigned a $z_{\text{spec}} = 4.4266$ and Flag 4. The main spectral features identified (emission lines and interstellar absorption lines) are indicated in the 1-D spectrum.

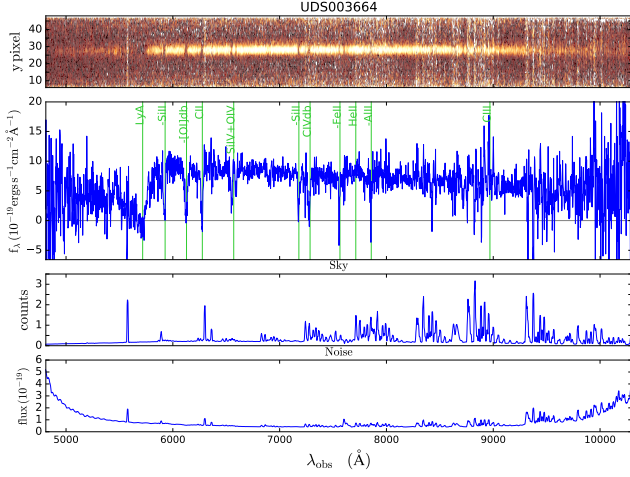


Fig. 12. From top to bottom, the 2-dimensional spectrum, 1-dimensional extracted spectrum, sky counts and rms noise of target VANDELS-UDS-003664. This object has an H-band magnitude of 23.1 and was selected in the category SFG. It was assigned a $z_{spec} = 3.703$ and Flag 4. The main spectral features identified (the Ly α absorption and several interstellar absorption lines) are indicated in the 1-D spectrum.

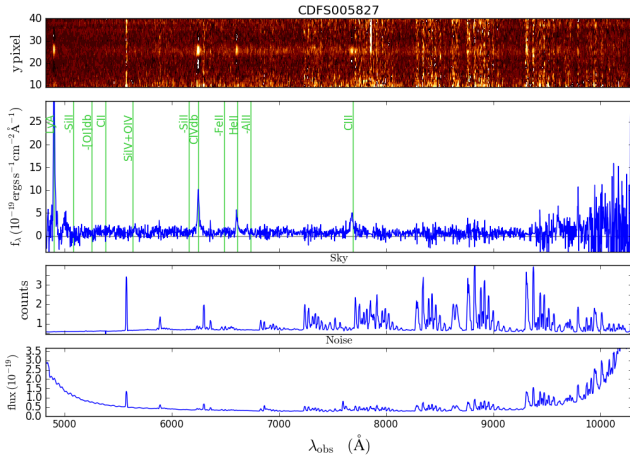


Fig. 13. From top to bottom, the 2-dimensional spectrum, 1-dimensional extracted spectrum, sky counts and rms noise of target VANDELS-CDFS-005827. This object was selected as an AGN target and indeed it is identified as an AGN (quality flag 14) at $z=3.0328$. The main spectral features identified (the Ly α , CIV, HeII and CIII] emission lines) are indicated in the 1-D spectrum.

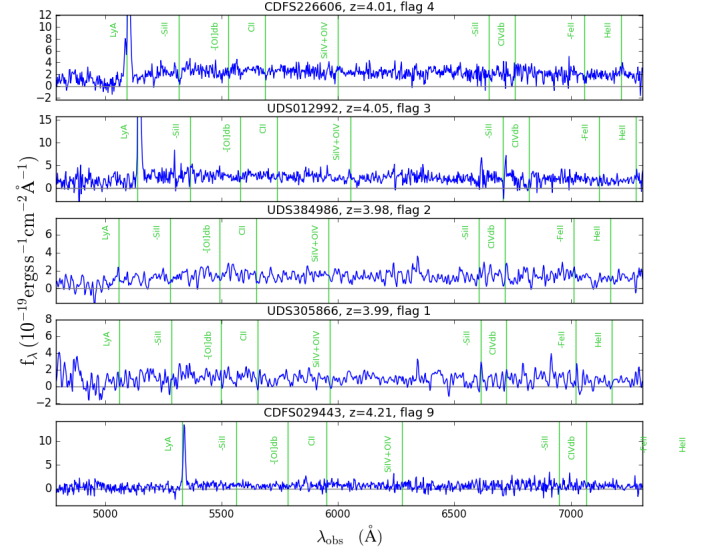


Fig. 14. From top to bottom, 1-dimensional extracted spectra of five VANDELS LBGs at redshift $z \sim 4$ with different quality flag: VANDES-CDFS-226606 with QF=4, VANDELS-UDS-012992 with QF=3, VANDELS-UDS-384986 with QF=2, VANDELS-UDS-305866 with QF=1 and VANDELS-CDFS-029443 with QF=9. The spectra of the QF=1 and 2 galaxies have been slightly smoothed to better show the absorption lines.